**Title:** Gauss-Bonnet Gravity in 4D and the connection to Horndeski's Theory

**Speakers:** Pedro Fernandes

Collection/Series: 50 Years of Horndeski Gravity: Exploring Modified Gravity

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#### **Abstract:**

In this talk I will review the topic of 4D Einstein-Gauss-Bonnet gravity, which has been the subject of considerable interest over the past years. I will discuss the mathematical complexities involved in implementing this idea, and review recent attempts at constructing well-defined, self-consistent theories that enact it, and their relation to Horndeski gravity. I then move on to consider the interesting phenomenology that results from these theories.

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# Gauss-Bonnet Gravity in 4D and the connection to Horndeski's Theory

Pedro G. S. Fernandes

CP3-Origins, University of Southern Denmark 50 Years of Horndeski Gravity: Exploring Modified Gravity July 2024

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#### Lovelock's Theorem

The only second-order, local gravitational field equations derivable from an action containing solely the 4-dimensional metric tensor (plus related tensors) are the Einstein field equations with a cosmological constant.

#### Our options for modifying gravity:

- 1. New field content
- 2. Higher dimensions
- 3. Derivatives of higher order in the field equations
- 4. Non-locality
- 5. Non-derivable from an action principle

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# The Gauss-Bonnet Term

$$\mathcal{G} = R^2 - 4R_{\mu\nu}R^{\mu\nu} + R_{\mu\nu\alpha\beta}R^{\mu\nu\alpha\beta}$$

- ► Natural generalization of Einstein's gravity in higher-dimensions (EGB gravity)
- ▶ Part of Horndeski's gravity  $(F(\phi)\mathcal{G})$
- ► Motivated by quantum gravity (string theory low-energy EFTs, trace anomaly...)

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#### The Gauss-Bonnet Term in 4D

Gauss-Bonnet action in *D*-dimensions

$$S_{\mathcal{G}} = \alpha \int d^D x \sqrt{-g} \mathcal{G}$$

In D = 4 we have that

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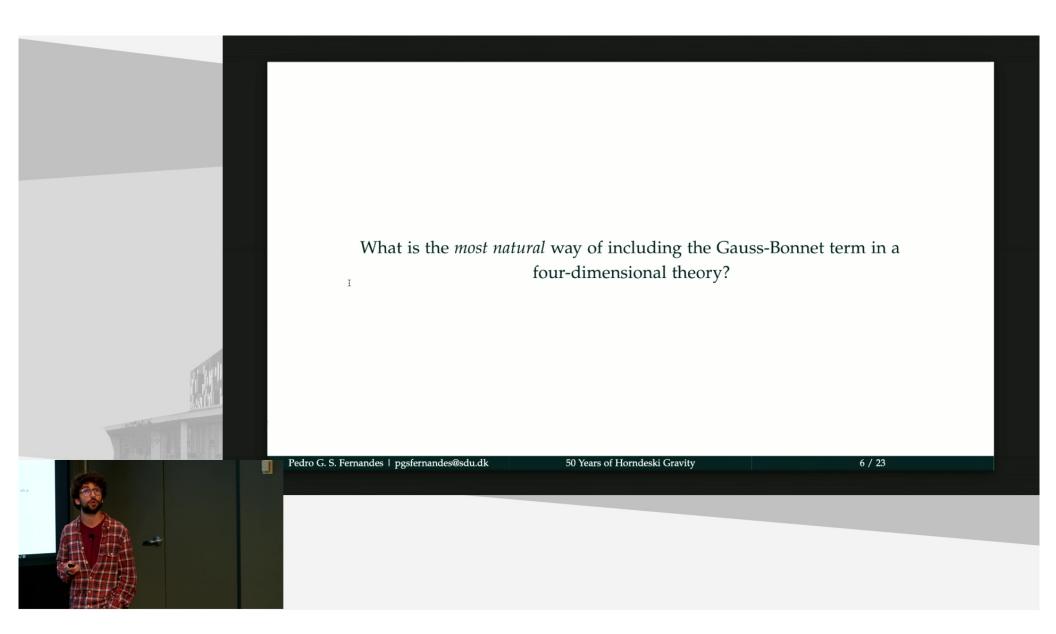
$$rac{\delta}{\delta g^{\mu
u}} S_{\mathcal{G}} \equiv H_{\mu
u}$$

is identically vanishing (Euler characteristic  $\chi \propto S_{\mathcal{G}} \rightarrow H_{\mu\nu} = 0$ ). Manifest in the trace:

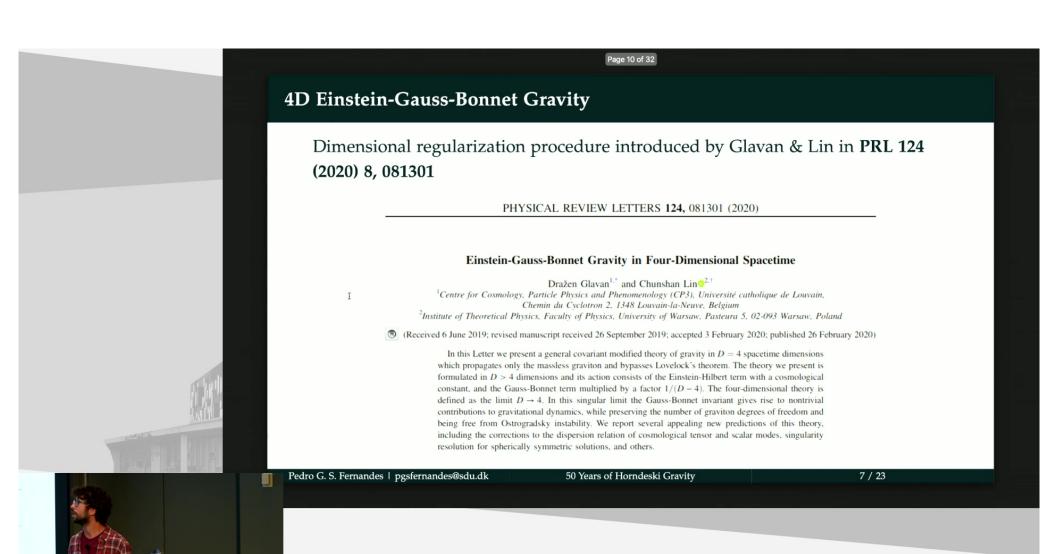
$$g^{\mu\nu}\frac{\delta}{\delta g^{\mu\nu}}S_{\mathcal{G}}=g^{\mu\nu}H_{\mu\nu}=\alpha\frac{(D-4)}{2}\mathcal{G}$$

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# 4D Einstein-Gauss-Bonnet Gravity

1. Start with the Einstein-Gauss-Bonnet action in *D* dimensions and introduce a divergent factor for the Gauss-Bonnet sector

$$S = \lim_{D \to 4} \left[ \frac{1}{16\pi} \int d^{\mathbf{D}} x \sqrt{-g} \left( R + \frac{\alpha}{D - 4} \mathcal{G} \right) + S_M \right]$$

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### **4D Einstein Gauss-Bonnet Gravity**

3. The trace equation, for instance, takes the finite form

$$\lim_{D \to 4} \left[ g^{\mu\nu} \left( G_{\mu\nu} + \frac{\alpha}{D - 4} H_{\mu\nu} \right) = 8\pi T \right] \Leftrightarrow$$

$$\lim_{D \to 4} \left[ \frac{(D - 2)}{2} R + \frac{\alpha}{(D - 4)} \frac{(D - 4)}{2} \mathcal{G} = -8\pi T \right] \Leftrightarrow$$

$$R + \frac{\alpha}{2} \mathcal{G} = -8\pi T,$$



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#### 4D Einstein Gauss-Bonnet Gravity

3. The trace equation, for instance, takes the finite form

$$\lim_{D o 4}\left[g^{\mu
u}\left(G_{\mu
u}+rac{lpha}{D-4}H_{\mu
u}
ight)=8\pi T
ight]\Leftrightarrow \ \lim_{D o 4}\left[rac{(D-2)}{2}R+rac{lpha}{(D-4)}rac{(D-4)}{2}\mathcal{G}=-8\pi T
ight]\Leftrightarrow \ R+rac{lpha}{2}\mathcal{G}=-8\pi T,$$

4. The equations of motion for certain line elements, such as static spherically symmetric or FLRW do not diverge in the 4D limit. Solutions similar to the higher-dimensional EGB ones.

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## 4D Einstein Gauss-Bonnet Gravity – Black holes and Friedmann equations

Gauss-Bonnet corrected black holes:

$$ds^2 = -f(r)dt^2 + \frac{dr^2}{f(r)} + r^2 \left(d\theta^2 + \sin^2\theta d\varphi^2\right),$$

$$f(r) = 1 + \frac{r^2}{2\alpha} \left(1 - \sqrt{1 + \frac{8M\alpha}{r^3}}\right)$$

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Gauss-Bonnet corrected Friedmann equations:

$$H^2 + \alpha H^4 = \frac{8\pi G}{3}\rho + \frac{\Lambda}{3},$$
  
 $\left(1 + 2\alpha H^2\right)\dot{H} = -4\pi G\left(\rho + p\right).$ 

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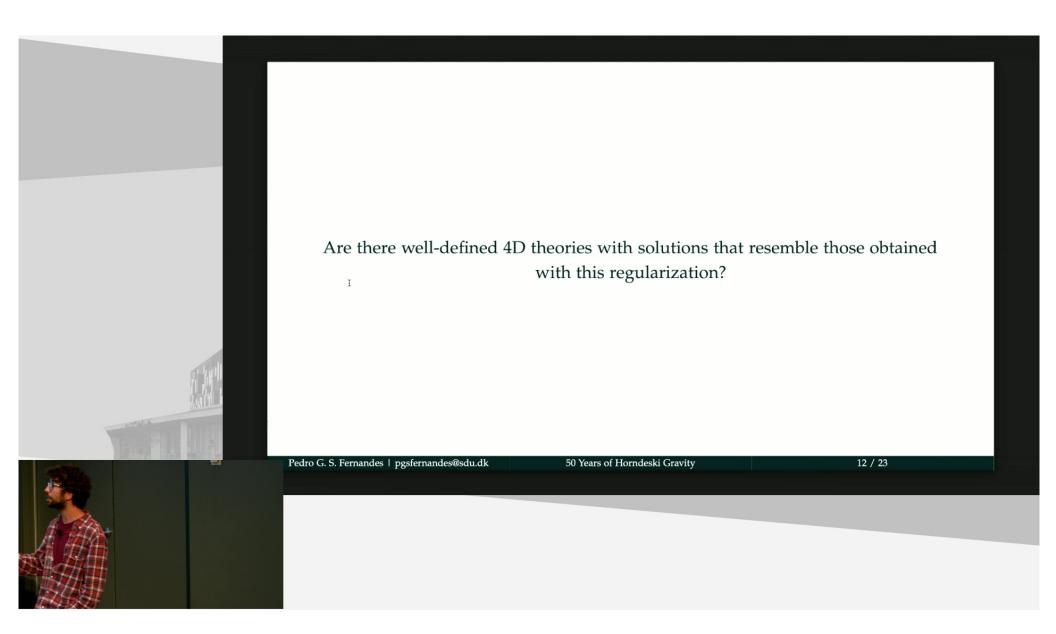
# **4D Einstein Gauss-Bonnet Gravity – Shortcomings**

- ► The process relies on the specification of the geometry of the extra dimensional space before taking the 4D limit. There are countless ways to do this.
- ► In spacetimes that lack explicit symmetries, the equations of motion (other than the trace equation) are not well-defined in the 4D limit

$$\lim_{D\to 4} \frac{H_{\mu\nu}}{D-4} = \text{finite term} + \lim_{D\to 4} \frac{1}{D-4} \left( C_{\mu\alpha\beta\rho} C_{\nu}^{\ \alpha\beta\rho} - \frac{1}{4} g_{\mu\nu} C^2 \right)$$

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#### Dimensional Regularization of the Gauss-Bonnet term (2004.08362)

- ▶ Problematic terms depend on the Weyl tensor
- ► The Weyl tensor is conformally invariant, i.e.,  $C^{\mu}_{\alpha\beta\rho} = \tilde{C}^{\mu}_{\alpha\beta\rho}$  given two conformally related metrics  $g_{\mu\nu} = e^{-2\phi}\tilde{g}_{\mu\nu}$
- ► We will then remove the divergences of the theory by adding to our action the counterterm

$$S = \lim_{D \to 4} \int d^D x \sqrt{-g} \left( R + \frac{\alpha}{D - 4} \mathcal{G} \right) - \lim_{D \to 4} \int d^D x \frac{\alpha}{D - 4} \sqrt{-\tilde{g}} \tilde{\mathcal{G}}$$

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# Dimensional Regularization of the Gauss-Bonnet term (2004.08362)

The limit is well-defined and results in

$$S = \int d^4x \sqrt{-g} \left[ R - \alpha \left( \phi \mathcal{G} - 4G^{\mu\nu} \nabla_{\mu} \phi \nabla_{\nu} \phi - 4 \Box \phi (\nabla \phi)^2 - 2(\nabla \phi)^4 \right) \right]$$

The theory belongs to the shift-symmetric ( $\phi o \phi + c$ ) Horndeski class of theories

$$G_2 = 8\alpha X^2$$
,  $G_3 = 8\alpha X$ ,  $G_4 = 1 + 4\alpha X$ ,  $G_5 = 4\alpha \log X$ 

Remarkably, a linear combination of the trace and scalar field equations leads to

$$R + \frac{\alpha}{2}\mathcal{G} = -8\pi T$$

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In this well-defined regularized theory we can recover the same static black holes and Friedmann equations as in the original regularization procedure!

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# Dimensional Regularization of the Gauss-Bonnet term

Regularized Kaluza-Klein reduction [Lu and Pang, 2003.11552]

$$S = \int d^4x \sqrt{-g} \left[ R - \alpha \left( \phi \mathcal{G} - 4G^{\mu\nu} \nabla_{\mu} \phi \nabla_{\nu} \phi - 4\Box \phi (\nabla \phi)^2 - 2(\nabla \phi)^4 + 2\lambda e^{2\phi} \left[ R + 6(\nabla \phi)^2 + 3\lambda e^{2\phi} \right] \right) \right]$$

The same combination of the trace and scalar field equations leads to the purely geometrical equation. Solutions also present.

See also [Aoki, Gorji, Mukohyama, 2005.03859, 2020] for diffeomorphism breaking regularization

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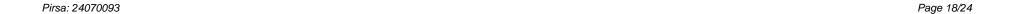
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What is the connection between these theories (with highly non-trivial structure), and why does a special combination of the field equations completely decouples from the scalar field?



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### Conformally coupled scalar field (2105.04687)

PHYSICAL REVIEW D 103, 104065 (2021)

#### Gravity with a generalized conformal scalar field: Theory and solutions

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(Received 2 December 2020; accepted 10 May 2021; published 28 May 2021)

We naturally extend general relativity with a conformally coupled scalar field by only requiring conformal invariance of the scalar field equation of motion and not of the action. The classically extended theory incorporates a scalar-Gauss-Bonnet sector and has second-order equations of motion, belonging to the Horndeski class. Remarkably, the theory features a purely geometrical field equation that allows for closed-form black hole solutions and cosmologies to be easily found. These solutions permit investigations of in-vogue scalar-Gauss-Bonnet corrections to the gravitational action without the need of resorting to approximations or numerical methods.

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# Conformally coupled scalar field (2105.04687)

► The answer is related to the conformal symmetry of the scalar field equation of motion:

$$\frac{\delta S}{\delta \phi}$$
 invariant under  $g_{\mu \nu} o g_{\mu \nu} e^{2\sigma}$ ,  $\phi o \phi - \sigma$ 

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#### Conformally coupled scalar field (2105.04687)

► The answer is related to the conformal symmetry of the scalar field equation of motion:

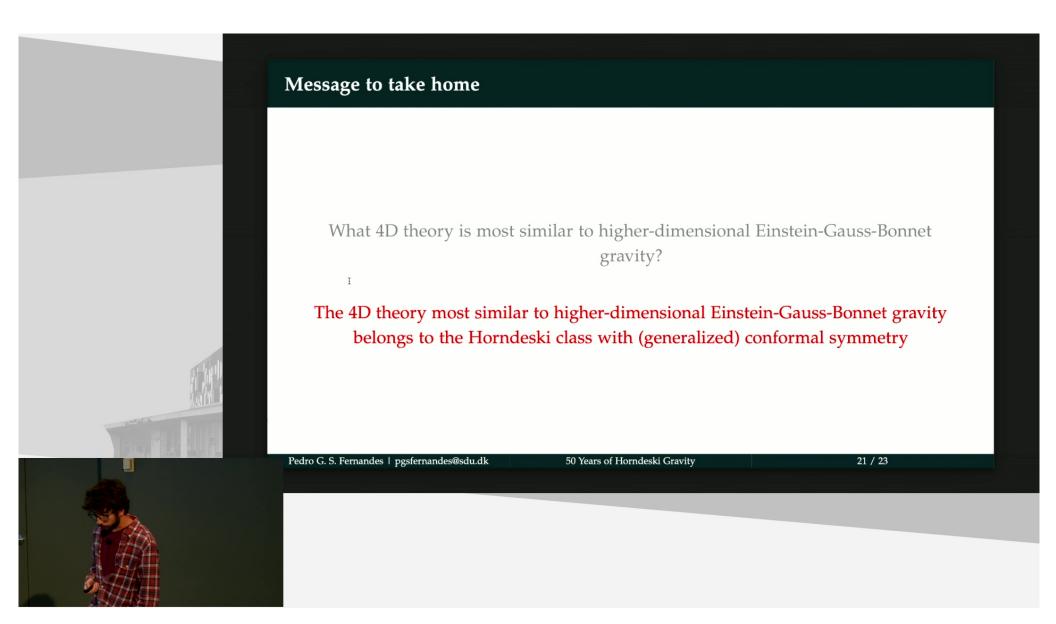
$$\frac{\delta S}{\delta \phi}$$
 invariant under  $g_{\mu\nu} \to g_{\mu\nu} e^{2\sigma}$ ,  $\phi \to \phi - \sigma$ 

- Because of this symmetry, we have nice integrability properties and solutions that resemble those of higher-dimensional EGB gravity.
   Both counterterm regularized and the KK theory have this symmetry.
- Most general theory with this symmetry and second-order equations of motion

$$G_2 = -2\Lambda - 2\gamma e^{4\phi} + 12\beta e^{2\phi}X + 8\alpha X^2, \quad G_3 = 8\alpha X$$
  
 $G_4 = 1 - \beta e^{2\phi} + 4\alpha X, \quad G_5 = 4\alpha \log X$ 

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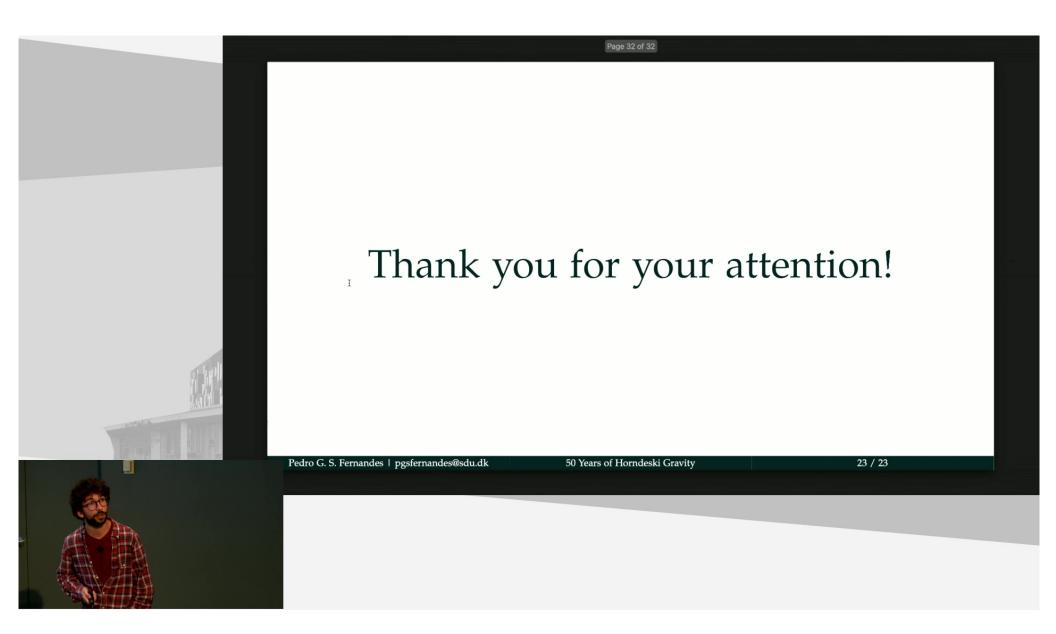
If you're interested in learning more please check the review "The 4D Einstein–Gauss–Bonnet theory of gravity: a review", Fernandes et. al, arXiv:2202.13908

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