Title: General Relativity for Cosmology Lecture - 112823

Speakers: Achim Kempf

Collection: General Relativity for Cosmology

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GR for Cosmology, Achim Kempf — Lecture 22

A singularity theorem:

Assume that:

- $(\mathcal{M}, g)$ is a globally hyperbolic spacetime
- The energy-momentum tensor of matter obeys the **Strong energy condition**:

$$\left(T_{\mu\nu} - \frac{1}{2} T^{g}{}_{g}\, g_{\mu\nu}\right)\xi^\mu \xi^\nu \geq 0 \quad \forall \text{ timelike } \xi.$$

Notice: Since the Einstein equation can be brought in the form $kR_{\mu\nu} = T_{\mu\nu} - \frac{1}{2}T g_{\mu\nu}$, the strong energy condition is a condition on the Ricci tensor too. This will be the use of the strong energy condition.

- There exists a C^2 spacelike Cauchy surface Σ, on which the trace of the extrinsic curvature, K, is bounded from above by a negative constant C:

$$K(p) \leq C < 0 \quad \text{for all } p \in \Sigma$$

Then: No past-directed timelike curve from a spacelike hypersurface Σ can have eigentime, i.e., proper length, larger than $\frac{3}{C}$.

I.e.: All past-directed timelike geodesics are incomplete.

because all past-directed paths end on it.

$\Rightarrow$ There is a cosmological singularity in the finite past!

Extrinsic curvature?

later more on this

- The extrinsic curvature of a spacelike hypersurface describes how much curvature there is in between the spacelike hypersurface and the time dimension.

Intuitively: it is the rate of the expansion of spacetime, more precisely its negative, the rate of contraction.

The strong energy condition?

Recall:

- The "weak energy condition":

$$T_{\mu\nu} v^\mu v^\nu \geq 0 \quad \text{for all timelike } v: g(v,v) < 0$$

Meaning? For an observer with unit tangent v the local energy density is: $T_{\mu\nu}v^\mu v^\nu \geq 0$

- The "dominant energy condition":

$$T_{\mu\nu} v^\mu v^\nu \geq 0 \quad \text{and} \quad K_\mu K^\mu \leq 0$$

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Strong energy condition :

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$$\left(T_{\mu\nu} - \tfrac{1}{2} T^{s}{}_{s}\, g_{\mu\nu}\right) \xi^\mu \xi^\nu \geq 0 \quad \forall \text{ timelike } \xi.$$

- There exists a C^2 spacelike Cauchy surface Σ, on which the trace of the extrinsic curvature, K, is bounded from above by a negative constant C:

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- The extrinsic curvature of a spacelike hypersurface describes how much curvature there is in between the spacelike hypersurface and the time dimension.

Intuitively: it is the rate of the expansion of spacetime, more precisely its negative, the rate of contraction.

Thus: Assuming $K(p) \leq C < 0$ $\forall p \in \Sigma$ means that spacetime has a finite minimum expansion rate everywhere on Σ.

$\rightsquigarrow$ We'll define expansion below in detail.

The strong energy condition?

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Meaning? For an observer with unit tangent v the local energy density is: $T_{\mu\nu} v^\mu v^\nu \geq 0$

- The "dominant energy condition":

$$\underbrace{T_{\mu\nu} v^\mu v^\nu \geq 0}_{\text{weak energy condition}} \quad \text{and} \quad K_\mu K^\mu \leq 0$$

i.e. $T_{\mu\nu} v^\nu$ is non-space-like.

where v is any timelike vector and $K_\mu := T_{\mu\nu} v^\nu$

Meaning? The local energy-momentum flow vector K may not be conserved but has to be non-space-like: Flow should be into the future $\leftarrow$ need for causality.

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Thus: Assuming $K(p) \leq C < 0$ $\forall p \in \Sigma$ meant that spacetime has a finite minimum expansion rate everywhere on Σ.

⇝ We'll define expansion below in detail.

□ The "strong energy condition"

Matter is said to obey the strong energy condition iff:

$$\left(T_{\mu\nu} - \frac{1}{2} T^{\rho}{}_{\rho} g_{\mu\nu}\right) \xi^{\mu} \xi^{\nu} \geq 0 \ \forall \text{ timelike } \xi.$$

□ Intuition? Excludes matter that causes accelerated expansion. (as we will discuss below)

□ Plausible? Yes, obeyed by known matter. (but not by dark energy)

□ Relationship? Independent of weak and dominant energy condition.

Meaning: for an observer with unit tangent v the local energy density is: $T_{\mu\nu} v^{\mu} v^{\nu} \geq 0$

□ The "dominant energy condition":

$\underbrace{T_{\mu\nu} v^{\mu} v^{\nu} \geq 0}_{\text{weak energy condition}}$ and $K_{\mu} K^{\mu} \leq 0$ (i.e. $T_{\mu\nu} v^{\nu}$ is non-space-like.)

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Meaning? The local energy-momentum flow vector K may not be conserved but has to be non-space-like: Flow should be into the future ← need for causality.

Concretely: For known matter, $T_{\mu\nu}$ is diagonalizable to obtain:

$$T_{\mu\nu} = \begin{pmatrix} \rho & & & 0 \\ & p_1 & & \\ & & p_2 & \\ 0 & & & p_3 \end{pmatrix}$$

energy density observed by comoving observer; principal pressures

The energy conditions then read:

□ Weak: $\rho \geq 0$ and $\rho + p_i \geq 0$ for $i \in \{1,2,3\}$

□ Dominant: $\rho \geq |p_i|$ for $i \in \{1,2,3\}$

Exercise: Show this → □ Strong: $\rho + \sum_{i=1}^{3} p_i \geq 0$ and $\rho + p_i \geq 0$ for $i \in \{1,2,3\}$ (Note: could possibly be also negative.)

Recall: A cosmological constant Λ can be viewed as a contribution to $T_{\mu\nu}$.

Indeed, there is no big bang singularity, e.g., if $w = -1$ $\forall t$, i.e., in de Sitter spacetime inflation $a(t) = e^{Ht}$

Exercise: Show that the strong energy condition is violated in cosmology iff $w < -\frac{1}{3}$, i.e., iff the expansion is accelerating: $\ddot{a}(t) >$

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Exercise: Show that the strong energy condition is violated in cosmology iff $w < -\frac{1}{3}$, i.e., iff the expansion is accelerating: $\ddot{a}(t) > 0$.

Given, in particular, the strong energy condition, one can show that geodesics meet a divergence of a quantity called expansion, θ, in finite proper time:

important notion also e.g. in study of grav. collapse of stars.

The "expansion", θ:

- Consider a "congruence of timelike geodesics"
e.g., freely falling dust.
through Σ, i.e., a smooth family of timelike geodesics, exactly one through each $p \in \Sigma$. If parametrized by proper time, their tangent vector field ξ, namely

$$\xi := \frac{d}{d\tau} \leftarrow \text{proper time}$$

will obey: $g(\xi,\xi) = -1\ \forall p$.

- Consider now a one-parameter subfamily of these geodesics:

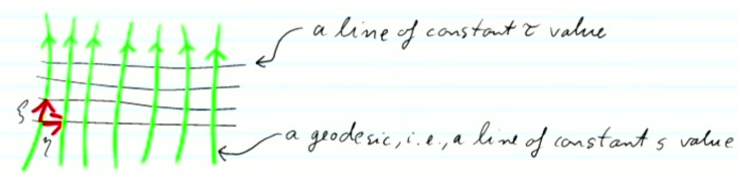
$$\gamma(\tau, s)$$

s: parameter of family of neighboring geodesics.

a "connecting vector field"

Then, we define the deviation vector:

$$\eta := \frac{d}{ds}$$



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The "expansion", θ:

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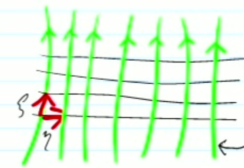

- How does η change along a geodesic?

  τ, s are Riemann normal coordinates for a geodesic traveller.

  $\Rightarrow \frac{d}{d\tau}\frac{d}{ds} = \frac{d}{ds}\frac{d}{d\tau}$, i.e., $[\xi, \eta] = 0$

- Since the torsion vanishes: $0 = T(\xi,\eta) = \nabla_\xi \eta - \nabla_\eta \xi - [\xi,\eta]$

$$\Rightarrow \quad \nabla_\xi \eta = \nabla_\eta \xi$$

$$\Rightarrow \quad \xi^\mu \nabla_{e_\mu} \eta^\nu e_\nu = \eta^\alpha \nabla_{e_\alpha} \xi^\beta e_\beta$$

$$\Rightarrow \quad \xi^\mu \eta^\nu{}_{;\mu} e_\nu = \eta^\alpha \xi^\beta{}_{;\alpha} e_\beta$$

$$\Rightarrow \quad \xi^\mu \eta^\nu{}_{;\mu} = \eta^\alpha \xi^\nu{}_{;\alpha} = \eta^\alpha B^\nu{}_\alpha \text{ for } \boxed{B^\nu{}_\mu := \xi^\nu{}_{;\mu}}$$

$\Rightarrow$ Along the geodesic's direction, ξ, the deviation vector η^μ changes its direction and length by $B^\nu{}_\mu \eta^\mu$.

- The tensor $B^\nu{}_\mu$ can be decomposed covariantly and uniquely into:

$$B_{\mu\nu} = \omega_{\mu\nu} + \sigma_{\mu\nu} + t_{\mu\nu}$$

Cosmic ballet tensor field. ($\omega_{\mu\nu}$: antisymmetric; $\sigma_{\mu\nu}$: symmetric and trace $= 0$; $t_{\mu\nu}$: rest)

(all 3 terms are tensors because the split is covariant)

We have: $\omega_{\mu\nu} = \frac{1}{2}(B_{\mu\nu} - B_{\nu\mu})$, clearly.

But $\sigma_{\mu\nu}, t_{\mu\nu} = ?$

In preparation: define the projector $h_{\mu\nu}$ onto $(\mathbb{R}\xi)^\perp$ i.e. onto the spatial components: (ξ is timelike)

$$h_{\mu\nu} := g_{\mu\nu} + \xi_\mu \xi_\nu$$

Check: is $h_{\mu\nu} w^\nu$ really always $\perp$ to ξ?

Indeed: $\xi^\mu h_{\mu\nu} w^\nu = (\xi, w) + \underbrace{(\xi,\xi)}_{=-1}(\xi, w) = 0$

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$$\Rightarrow \quad \nabla_\xi \eta = \nabla_\eta \xi$$

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Define: The "expansion", θ, is defined as the magnitude of the spatial part of B:

$$\theta := B^{\mu\nu} h_{\mu\nu}$$

Claim: $Tr(B) = \theta$

Indeed: $\theta = B^{\mu\nu} h_{\mu\nu} = B^{\mu\nu} g_{\mu\nu} + \xi^\mu \xi_\nu B_\mu{}^\nu$
$= Tr(B) + \xi^\mu \underbrace{\xi_\nu \nabla_\mu \xi^\nu}$

$\left(= 0 \text{ because } \nabla_\xi \xi = 0 \text{ for geodesics.}\right)$

Therefore: $\sigma_{\mu\nu} = \frac{1}{2}(B_{\mu\nu} + B_{\nu\mu}) - \frac{1}{3}\theta h_{\mu\nu}$

$\left(\text{because: } Tr(h_{\mu\nu}) = g^{\mu\nu} h_{\mu\nu} = g^{\mu\nu}(g_{\mu\nu} + \xi_\mu \xi_\nu) = 4 - 1\right)$

↳ the part of $B_{\mu\nu}$ which is symmetric and traceless.

and:

$t_{\mu\nu} = \frac{1}{3}\theta h_{\mu\nu}$ ← the "rest term".

▫ Interpretation:

a) $\omega_{\mu\nu}$ is antisymmetric: $\omega_{\mu\nu} = -\omega_{\nu\mu}$

$\Rightarrow$ it generates Lorentz transformation for η.

but all η are $\perp$ to the time direction

$\Rightarrow$ $\omega_{\mu\nu}$ generates spatial rotations of neighboring geodesics around another. So, $\omega_{\mu\nu}$ is called

$\omega =$ "Twists tensor"

One can prove: (nontrivial)

If one chooses the congruence of geodesics $\perp$ to Σ then $\omega_{\mu\nu} = 0$.

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$\Longrightarrow \quad \nabla_\xi \eta = \nabla_\eta \xi$

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▯ The tensor $B^\nu{}_\mu$ can be decomposed covariantly and uniquely into:

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$\uparrow$ is timelike

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Therefore: $\sigma_{\mu\nu} = \frac{1}{2}(B_{\mu\nu} + B_{\nu\mu}) - \frac{1}{3}\theta h_{\mu\nu}$ $\left(\text{because: } Tr(h_{\mu\nu}) = g^{\mu\nu} h_{\mu\nu} = g^{\mu\nu}(g_{\mu\nu} + \xi_\mu \xi_\nu) = 4 - 1 \right)$

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Therefore:

$$\sigma_{\mu\nu} = \frac{1}{2}(B_{\mu\nu} + B_{\nu\mu}) - \frac{1}{3}\theta h_{\mu\nu}$$

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the part of $B_{\mu\nu}$ which is symmetric and traceless.

and:

$$t_{\mu\nu} = \frac{1}{3}\theta h_{\mu\nu}$$ ← the "rest term".

⇒ $\omega_{\mu\nu}$ generates spatial rotations of neighboring geodesics around another. So, $\omega_{\mu\nu}$ is called

$\omega =$ "**Twists tensor**"

One can prove: (nontrivial)

If one chooses the congruence of geodesics $\perp$ to Σ then $\omega_{\mu\nu} = 0$.

b.) $\sigma_{\mu\nu}$ is symmetric, $\sigma_{\mu\nu} = \sigma_{\nu\mu}$. (i.e. hermitean)

Consider "diagonalized", by suitable choice of cd basis.

⇒ $\sigma_{\mu\nu}$ changes the relative lengths of the basis vectors, by multiplying them with its eigenvalues.

i.e. points on a sphere will under geodesic flow → become points on an ellipsoid.

Note: Since $Tr(\sigma) = 0$ we have $\det(e^{\epsilon\sigma}) = 1$

infinitesimal transport along geodesics

finite transport

⇒ The volume spanned by basis vectors stays the same under the action of σ.

⇝ **Definition:** $\sigma_{\mu\nu} =:$ "**Shear tensor**" 

c.) While the twist and shear tensors are both traceless and therefore volume-preserving, we see that the trace part, θ, i.e., more precisely

$t_{\mu\nu} = \frac{1}{3}\theta h_{\mu\nu} =:$ "**Expansion tensor**"

recall: is projector on spatial part.

is indeed generating the spatial expansion or contraction of nearby geodesics!

Evolution of θ along a geodesic?

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$t_{\mu\nu} = \frac{1}{3}\,\theta\, h_{\mu\nu} =:$ "Expansion tensor"

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Evolution of θ along a geodesic?

Recall:

Given, in particular, the strong energy condition, our singularity theorem claimed that geodesics meet a divergence of a quantity called expansion, θ, in finite proper time in the past and this will mean a big bang singularity:

important notion also e.g. in study of grav. collapse of stars.

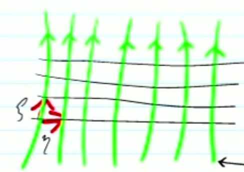
The "expansion", θ:

- Consider a "congruence of timelike geodesics"
  e.g., freely falling dust.
  through Σ, i.e., a smooth family of timelike geodesics, exactly one through each $p \in \Sigma$: (Σ is a Cauchy surface)

- We consider a one-parameter sub-family of these geodesics:

$$\gamma(\tau, s)$$

eigentime; parameter of family of neighboring geodesics.



- Then, we define the deviation vector to a neighboring geodesic:

$$\eta := \frac{d}{ds}$$

- The singularity theorem claims that this happened in the past:

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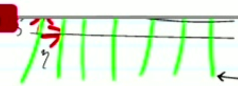

← a geodesic, i.e., a line of constant s value

- Then, we define the deviation vector to a neighboring geodesic:

$$\eta := \frac{d}{ds}$$

- The singularity theorem claims that this happened in the past:

big bang singularity

How does η change along a past-directed timelike geodesic with tangent ξ?

We showed:

$$\boxed{\xi^\mu \eta^\nu{}_{;\mu} = \eta^\alpha B^\nu{}_\alpha \quad \text{where} \quad B^\nu{}_\mu := \xi^\nu{}_{;\mu}}$$

$\Rightarrow$ Along the geodesic, ξ, the deviation vector η^ν changes its direction and length by $B^\nu{}_\mu \eta^\mu$.

- The tensor $B^\nu{}_\mu$ can be decomposed covariantly and uniquely:

$$B_{\mu\nu} = \underbrace{\omega_{\mu\nu}}_{\text{antisymmetric}} + \overbrace{\sigma_{\mu\nu}}^{\text{symmetric and trace}=0} + \underbrace{t_{\mu\nu}}_{\text{rest}}$$

Explicitly:

Volume preserving →

$$\omega_{\mu\nu} = \frac{1}{2}(B_{\mu\nu} - B_{\nu\mu}) \qquad \text{Twist:}$$

$$\sigma_{\mu\nu} = \frac{1}{2}(B_{\mu\nu} + B_{\nu\mu}) - \frac{1}{3}\theta h_{\mu\nu} \qquad \text{Shear:}$$

Volume changing:

$$t_{\mu\nu} = \frac{1}{3}\theta h_{\mu\nu} \qquad \text{Expansion:}$$

Here, we defined: $\theta := B^{\mu\nu} g_{\mu\nu}$ and $h_{\mu\nu} := g_{\mu\nu} + \xi_\mu \xi_\nu$

I.e., the Expansion, θ, is the trace of B, which we showed is also equal to the magnitude of the spatial part of B: $\theta = B^{\mu\nu} h_{\mu\nu}$.

Key question: What is the dynamics of θ?

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The Raychaudhuri equation

For the derivation, we will use:

A) Definition of B is: $B_{\mu\nu} := \xi_{\mu;\nu}$

B) The curvature tensor obeys the Ricci equation:

$$\boxed{\xi^a{}_{;bc} - \xi^a{}_{;cb} = R^a{}_{bcd}\,\xi^d}$$

c) ξ is tangent to a geodesic, i.e., it obeys: $\nabla_\xi \xi = 0$

i.e.: $0 = \nabla_{\xi} \xi^b e_b = \xi^a \nabla_{e_a} \xi^b e_b = \xi^a \xi^b{}_{;a} e_b$

True for all e_a, thus: $\boxed{\xi^a \xi^b{}_{;a} = 0}$

Now calculate the rate of change of B along the geodesic:

$$\underbrace{\xi^c B_{ab;c}}_{\nabla_\xi B} \overset{(A)}{=} \xi^c \xi_{a;bc}$$

$$\overset{(B)}{=} \xi^c \xi_{a;cb} + \xi^c R_{abcd}\,\xi^d$$

$$\underset{\text{rule}}{\overset{\text{Leibniz}}{=}} \underbrace{(\xi^c \xi_{a;c})_{;b}}_{=0} - \xi^c{}_{;b}\,\xi_{a;c} + R_{abcd}\,\xi^c \xi^d$$

$$\overset{(C)}{=} -\xi^c{}_{;b}\,\xi_{a;c} + R_{abcd}\,\xi^c \xi^d$$

$$\overset{(A)}{=} -B^c{}_b\,B_{ac} + R_{abcd}\,\xi^c \xi^d$$

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$\forall$ big bang singularity

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Volume preserving →
$$\sigma_{\mu\nu} = \frac{1}{2}(B_{\mu\nu} + B_{\nu\mu}) - \frac{1}{3}\theta h_{\mu\nu} \qquad \text{Shear: } \circ \to 0$$

Volume changing:
$$t_{\mu\nu} = \frac{1}{3}\theta h_{\mu\nu} \qquad \text{Expansion: } \circ \to \bigcirc$$

Here, we defined: $\theta := B^{\mu\nu} g_{\mu\nu}$ and $h_{\mu\nu} := g_{\mu\nu} + \xi_\mu \xi_\nu$

I.e., the Expansion, θ, is the trace of B, which we showed is also equal to the magnitude of the spatial part of B: $\theta = B^{\mu\nu} h_{\mu\nu}$.

Key question: $\boxed{\text{What is the dynamics of } \theta\text{?}}$

The Raychaudhuri equation

For the derivation, we will use:

A) Definition of B is: $B_{\mu\nu} := \xi_{\mu;\nu}$

B) The curvature tensor obeys the Ricci equation:

Now calculate the rate of change of B along the geodesic:

$$\underbrace{\xi^c B_{ab;c}}_{\nabla_\xi B} \overset{(A)}{=} \xi^c \xi_{a;bc}$$

$$\overset{(B)}{=} \xi^c \xi_{a;cb} + \xi^c R_{abcd} \xi^d$$

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A) Definition of B is: $B_{\mu\nu} := \xi_{\mu;\nu}$

B) The curvature tensor obeys the Ricci equation:

$$\boxed{\xi^a{}_{;bc} - \xi^a{}_{;cb} = R^a{}_{bcd}\,\xi^d}$$

c) ξ is tangent to a geodesic, i.e., it obeys: $\nabla_\xi \xi = 0$

i.e.: $0 = \nabla_{\xi^a e_a} \xi^b e_b = \xi^a \nabla_{e_a} \xi^b e_b = \xi^a \xi^b{}_{;a} e_b$

True for all e_a, thus: $\boxed{\xi^a \xi^b{}_{;a} = 0}$

$\nabla_\xi B$

$$\overset{(B)}{=} \xi^c \xi_{a;cb} + \xi^c R_{abcd}\,\xi^d$$

$$\overset{\text{Leibnitz rule}}{=} \underbrace{(\xi^c \xi_{a;c})_{;b}}_{=0} - \xi^c{}_{;b}\,\xi_{a;c} + R_{abcd}\,\xi^c \xi^d$$

$$\overset{(c)}{=} -\xi^c{}_{;b}\,\xi_{a;c} + R_{abcd}\,\xi^c\xi^d$$

$$\overset{(A)}{=} -B^c{}_b B_{ac} + R_{abcd}\,\xi^c\xi^d$$

In summary, we derived:

$$\xi^c B_{ab;c} = -B^c{}_b B_{ac} + R_{abcd}\,\xi^c\xi^d \qquad (*)$$

The trace of $(*)$ will be the Raychaudhuri equation.

But first, we recall:

- $\xi = \frac{d}{d\tau}$
- $\mathrm{Tr}\,B = B_{\mu\nu} g^{\mu\nu} = \theta$

$\Rightarrow$ Trace(LHS) of $(*)$ reads $\frac{d}{d\tau}\theta$!

Now on the RHS of $(*)$ use the decomposition

$B_{\mu\nu} = \omega_{\mu\nu} + \sigma_{\mu\nu} + \frac{1}{3}\theta h_{\mu\nu}$ to express $B^c{}_b B_{ac}$:

$$B^c{}_b B_{ac} = \omega^c{}_b\left(\underline{\omega_{ac}} + \sigma_{ac} + \tfrac{1}{3}\theta h_{ac}\right)$$
$$+ \sigma^c{}_b\left(\omega_{ac} + \underline{\sigma_{ac}} + \tfrac{1}{3}\theta h_{ac}\right)$$
$$+ \tfrac{1}{3}\theta h^c{}_b\left(\omega_{ac} + \sigma_{ac} + \underline{\tfrac{1}{3}\theta h_{ac}}\right)$$

When taking the trace, $g^{ab} B^c{}_b B_{ac}$, only the diagonal terms survive:

$$\mathrm{Tr}(BB) = g^{ab} B^c{}_b B_{ac} = \omega_{ab}\omega^{ab} + \sigma_{ab}\sigma^{ab} + \tfrac{1}{9}\theta^2 \overbrace{h_{ab}h^{ab}}^{=3}$$

Exercise: show it is 3

The Raychaudhuri equation is then the trace of Eq. $(*)$:

recall: Ricci tensor is $R_{cd} = R_{cda}{}^a$

$$\frac{d\theta}{d\tau} = -\frac{1}{3}\theta^2 - \underbrace{\sigma_{ab}\sigma^{ab}}_{\text{always positive}} - \underbrace{\omega_{ab}\omega^{ab}}_{\text{always positive (and vanishes if}} - \underbrace{R_{cd}\,\xi^c\xi^d}_{\text{pos. or neg?}}$$

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?

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?

Dynamics

- Assume that

$$R_{\mu\nu}\xi^\mu\xi^\nu \geq 0 \quad \text{for all timelike } \xi$$

i.e., using the Einstein equation

$$R_{\mu\nu} = 8\pi G\left(T_{\mu\nu} - \frac{1}{2} g_{\mu\nu} T^a{}_a\right)$$

we are assuming that

$$T_{\mu\nu}\xi^\mu\xi^\nu - \frac{1}{2}\overbrace{\xi^\mu\xi_\mu}^{=-1} T \geq 0 \quad \text{whenever } \xi^\mu\xi_\mu < 0$$

i.e. the Strong Energy Condition.

Thus, assuming the strong energy condition:

$$\frac{d\theta}{d\tau} + \frac{1}{3}\theta^2 \leq 0$$

$$\text{i.e., } -\frac{1}{\theta^2}\frac{d\theta}{d\tau} - \frac{1}{3} \geq 0$$

$$\text{i.e., } \boxed{\frac{d}{d\tau}\theta^{-1} \geq \frac{1}{3}} \qquad (+)$$

Consider the cases when the geodesics are initially all either

a.) diverging, i.e., $\theta(\tau_0) > 0$ (expanding universe) or

b.) converging, i.e., $\theta(\tau_0) < 0$ (contracting universe)

(This is reformulating the theorem's assumption that the extrinsic curvature (i.e. the expansion or contraction at some time exceeds a certain finite value everywhere)

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i.e., $\boxed{\frac{d}{d\tau}\theta^{-1} \geq \frac{1}{3}}$ (+)

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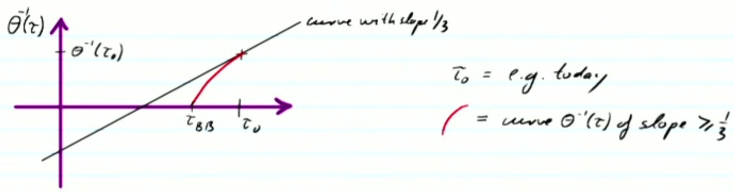
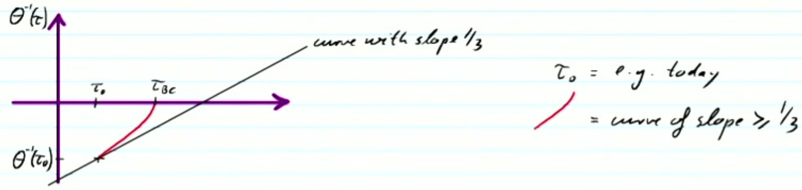
(This is reformulating the theorem's assumption that the extrinsic curvature (i.e. the expansion or contraction at some time exceeds a certain finite value everywhere)

a.)



We see that $\theta^{-1}(\tau)$ must have hit $\theta^{-1}(\tau)=0$ at a finite time τ_{BB} (Big Bang).

b.)



We see that $\theta^{-1}(\tau)$ will hit $\theta^{-1}(\tau)=0$ at a finite time τ_{BC} (Big Crunch)

Conclusion:

Eq. (+) implies that $\theta^{-1}(\tau)$ must go through 0, i.e.:

a.) for sufficiently early τ, have $\theta \to +\infty$, i.e.: **Big Bang**

b.) for sufficiently late τ, have $\theta \to -\infty$, i.e.: **Big Crunch**

Note: This type of reasoning leads also to further cosmological singularity theorems.

E.g., another cosmological singularity theorem does not assume global hyperbolicity, and its conclusion is weaker:

There is at least one incomplete timelike geodesic.

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