

Title: Spectral Search for Echoes

Speakers: Bob Holdom

Collection: Echoes in Southern Ontario

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Spectral Search for Echoes

Bob Holdom

PI Feb 25 2020

obtain the spectrum

- gravitational waves on a truncated Kerr background
- a complex amplitude describing two polarizations

$$h(t, \Omega) = h_+(t, \Omega) - i h_\times(t, \Omega) \approx h(t) Y_{22}(\Omega) \quad \ell = m = 2 \text{ dominates}$$

- $h(t)$ is related to ψ_ω , a quantity in frequency space
- ψ_ω can be obtained from a solution to the Sasaki-Nakamura (SN) equation

$$\frac{d^2\psi}{dx^2} - \mathcal{F}(\omega, x) \frac{d\psi}{dx} - \mathcal{U}(\omega, x)\psi = S(\omega, x)$$

- $\psi(\omega, x) \rightarrow \psi_\omega e^{i\omega x}$ for $x \rightarrow \infty$

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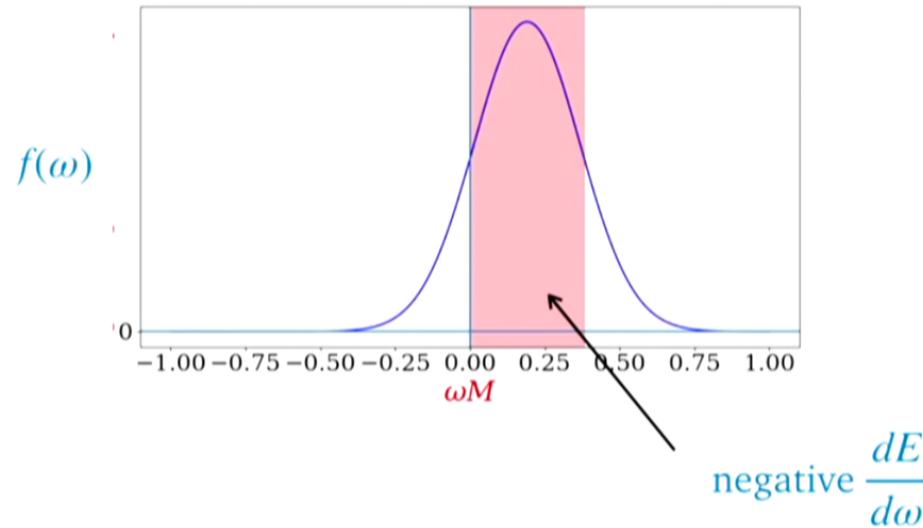
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initial condition from pulse starting on inside

- express initial condition in terms of an amplitude $f(\omega)$ that determines

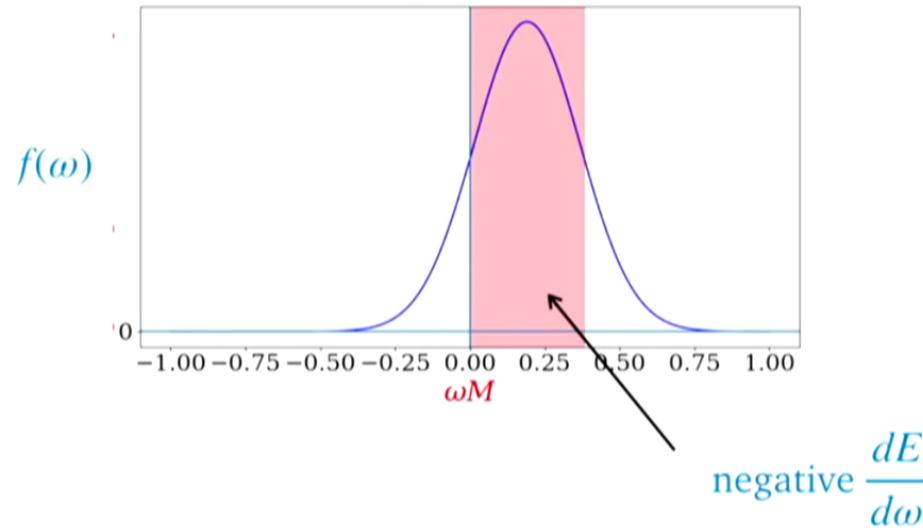
spectral flux density $\frac{dE}{d\omega} = 8\omega(\omega - \omega_0) |f(\omega)|^2$

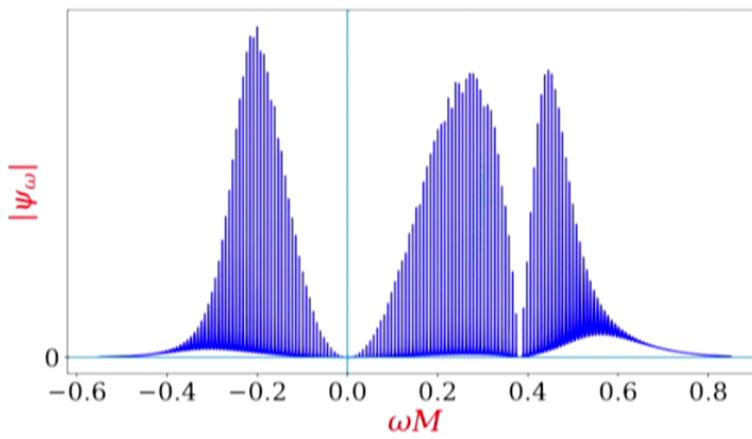


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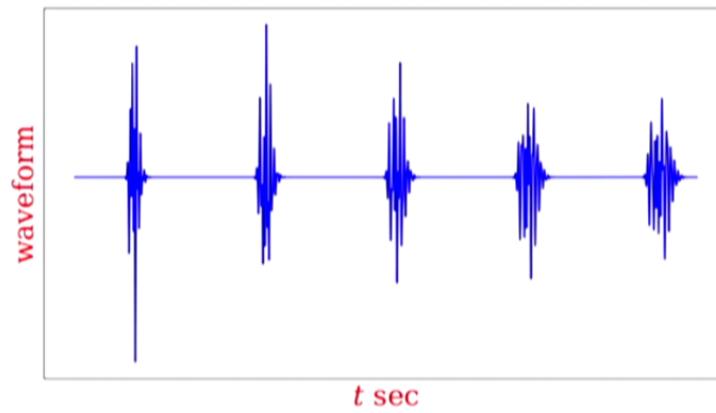
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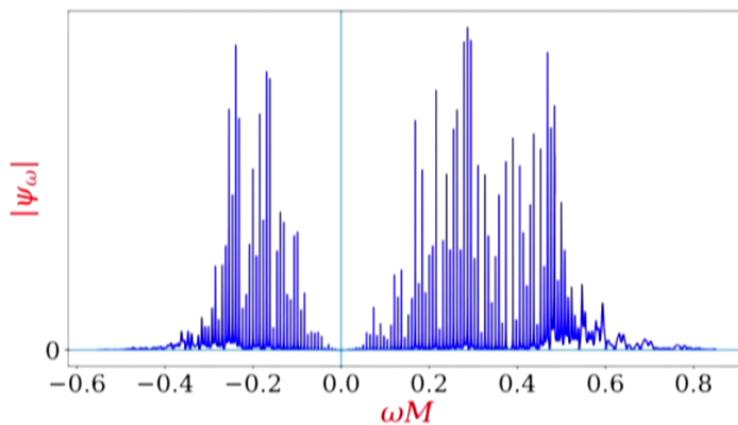




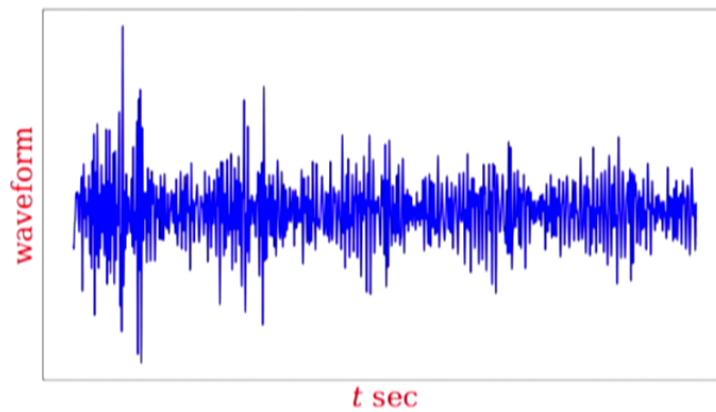
- single pulse in the cavity
- a zero is at $\omega_0 \propto \chi$ (the spin)



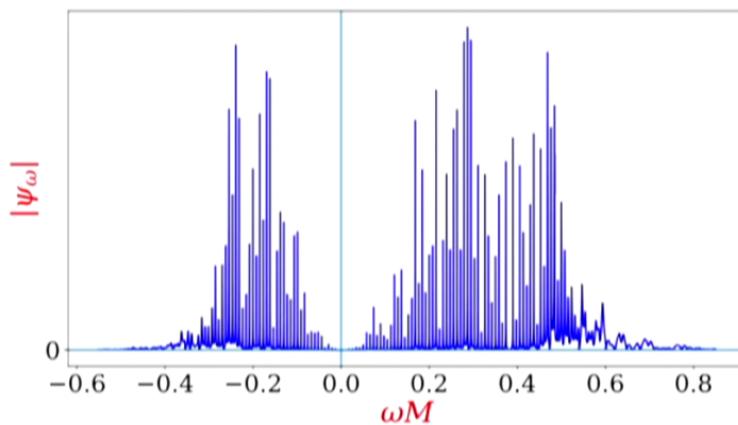
- irregular echoes



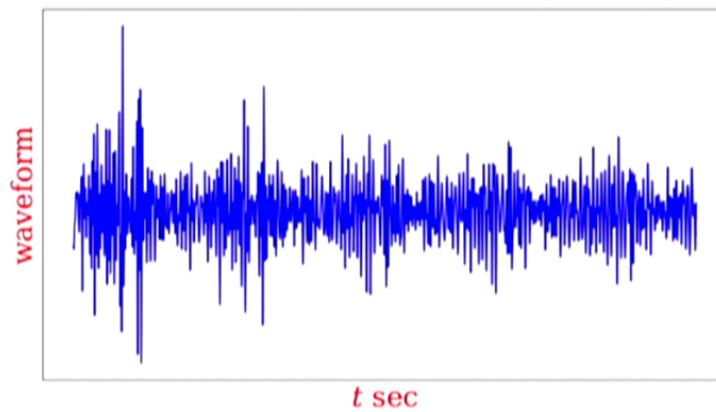
- 100 random pulses as an initial condition
- resonance structure is still present



- echo structure is gone



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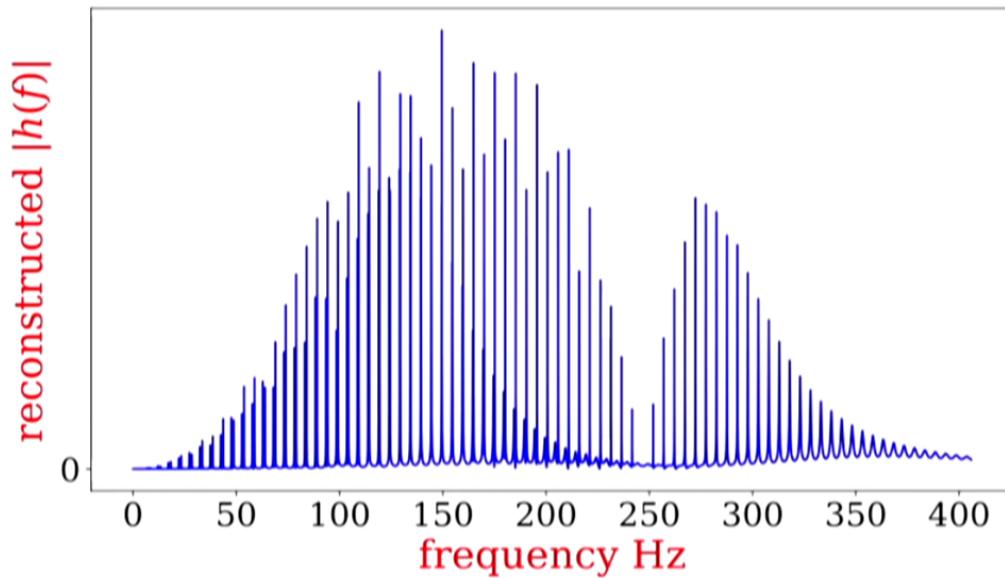


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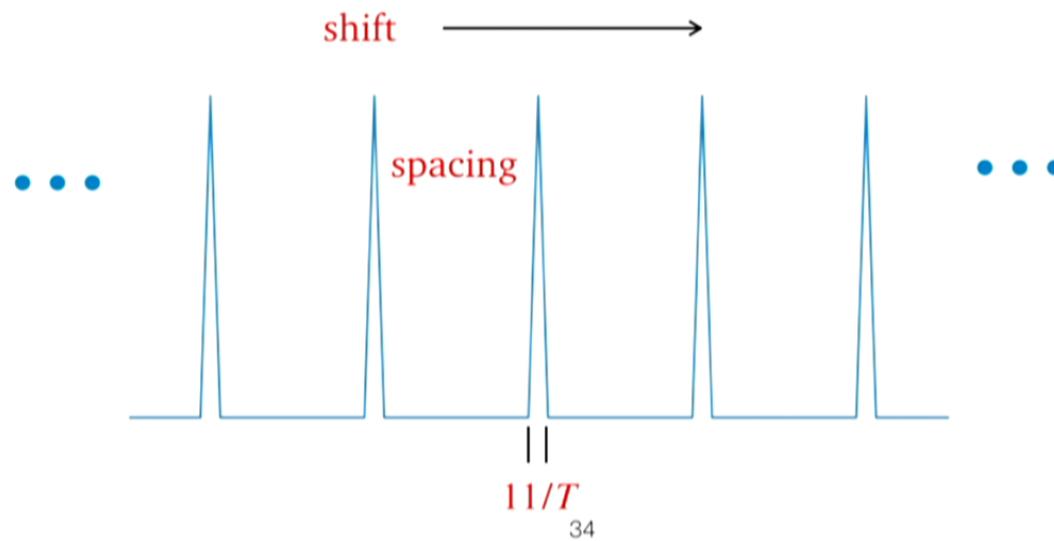
the “reconstructed” spectrum

- a LIGO detector projects the complex waveform into a set of real numbers, the “strain data”
- we take the **real part** of the signal waveform $h(t)$ to model this projection
- to search for a resonance pattern:
 - choose the time duration $T = N_E \Delta t$ of the data to analyze
 - take the FFT
 - take the absolute value
- carrying out these steps on the **real part** of $h(t)$ gives a “reconstructed” $|h(f)|$

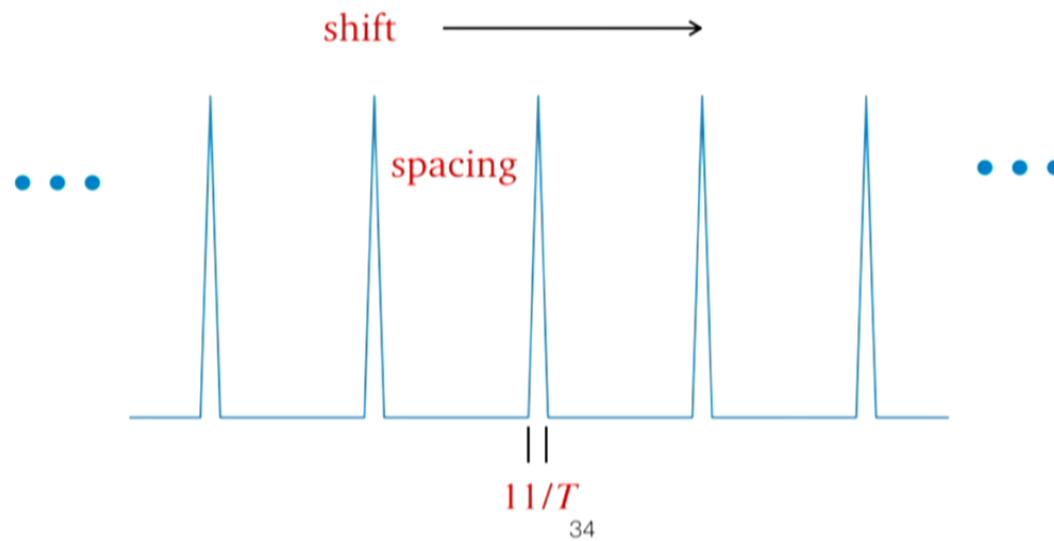
- an example for $N_E = 180$, $\chi = 2/3$, $M = 50M_\odot$, $\Delta t/M = 800$
- two component structure due to the real projection
- lower frequency structure comes from the negative frequency part of ψ_o



- take $|\text{FFT}(\text{data of duration } T \text{ after merger})|$ and choose a bandpass $f_{\min} < f < f_{\max}$
- multiply by a uniform comb structure characterized by a **spacing** between teeth and an overall **shift**
- take these amplitudes from the two detectors and search for a correlated enhancement for some spacing and shift



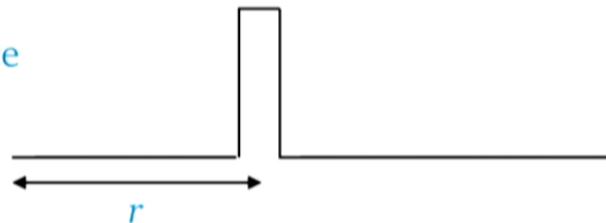
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in more detail ...

$$U = \text{mean} [|\text{FFT}(data)| \times \text{comb}(shift)]$$

- U is vector in shift space
- signal will produce larger values around some shift
- $V(r)$ is another vector in shift space
 - an idealized bump at shift r

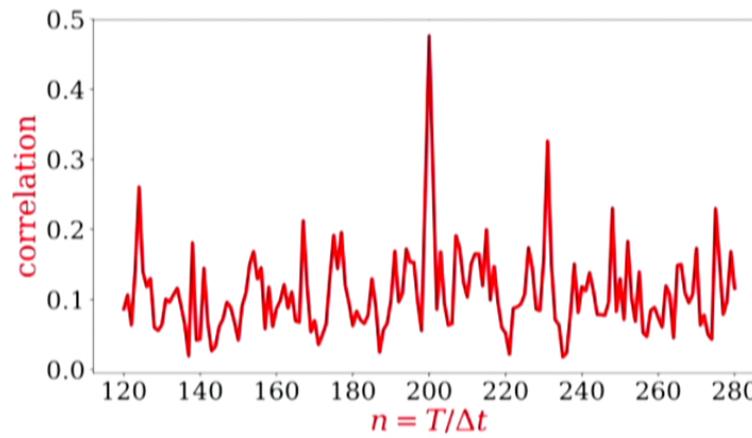


$$C = \max_r \text{Correlation}(U_H, V(r)) \times \text{Correlation}(U_L, V(r))$$

- C is our correlation between detectors, H and L , and it is still a function of the comb spacing $\Delta f = 1/\Delta t$

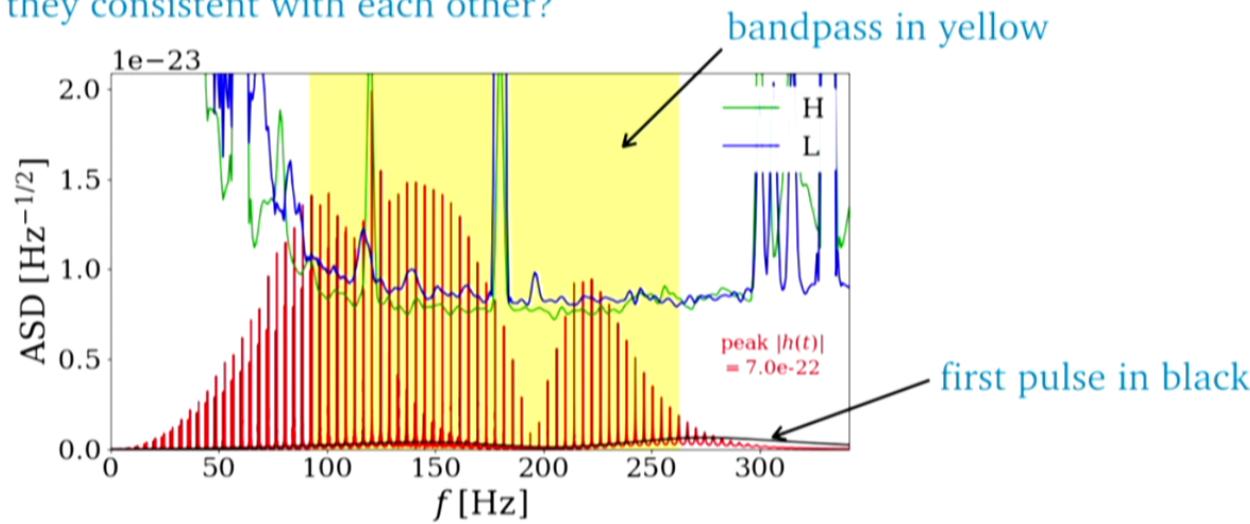
the “signal plot”

- consider discrete set of comb spacings Δf for integers $n = \Delta f/(1/T) = T/\Delta t$
- plot the correlation as a function of n and look for a peak
- can arrange for a peak to sit at $n = N_E$, some number of echoes
- peak should be enhanced by tuning Δt



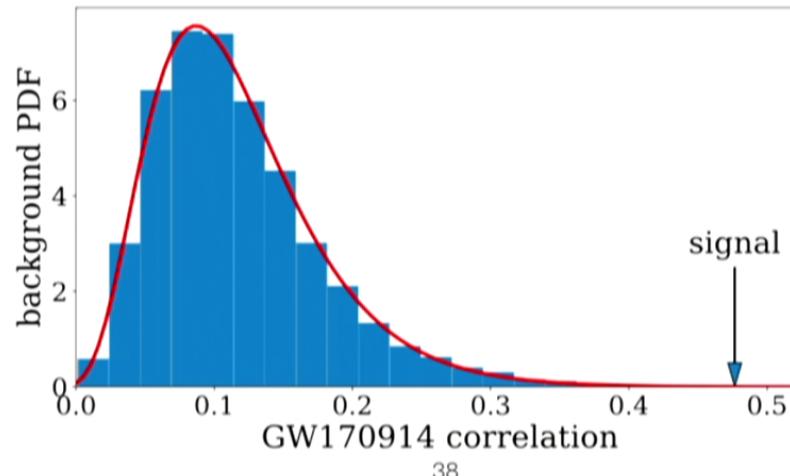
data driven bandpass

- optimize the signal (the peak height):
 - vary the number of echoes N_E (some multiple of 10)
 - vary the frequency bandpass $f_{\min} < f < f_{\max}$
- will compare this data driven bandpass to spectrum predicted by $M, \chi, \Delta t, N_E$
 - are they consistent with each other?



distribution of background correlations

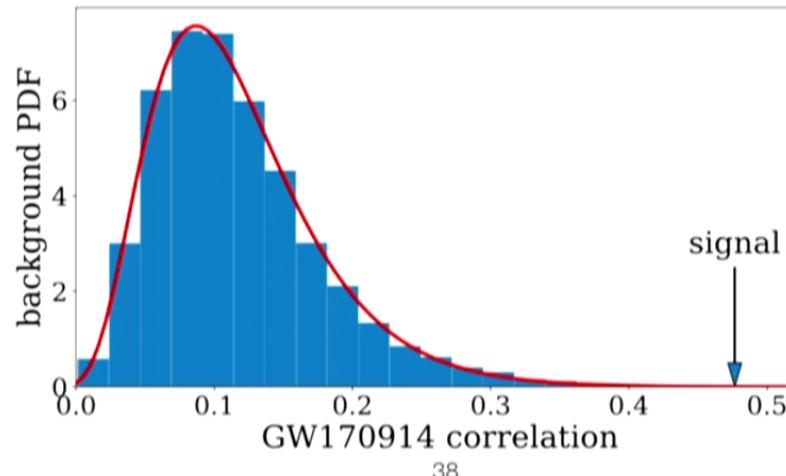
- generate many background versions of the signal plot
 - can check they do not display excess correlation at the signal location
- collect all the correlation values and fit to a PDF (generalized gamma)
 - $P_{gg}(x) \propto x^\alpha e^{-\beta x}$ (4 parameters: α, β , scale parameter, location parameter)
 - maximum likelihood fit using ~ 10000 to 20000 values
 - add histogram to show fit



38

distribution of background correlations

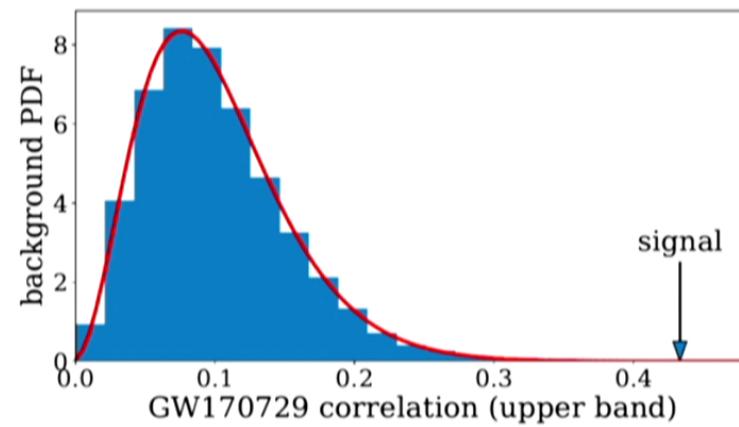
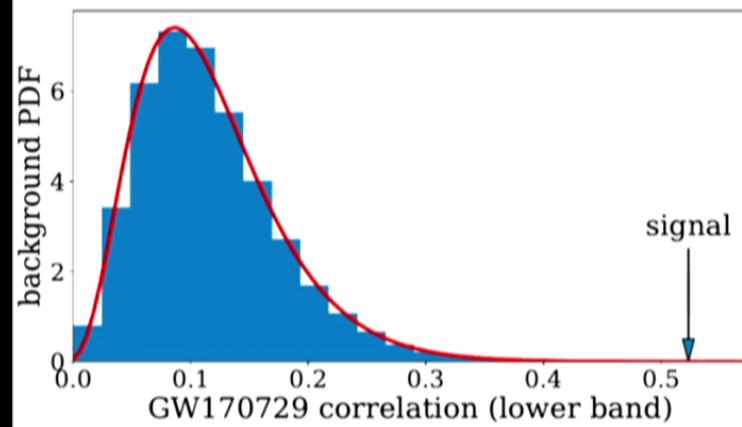
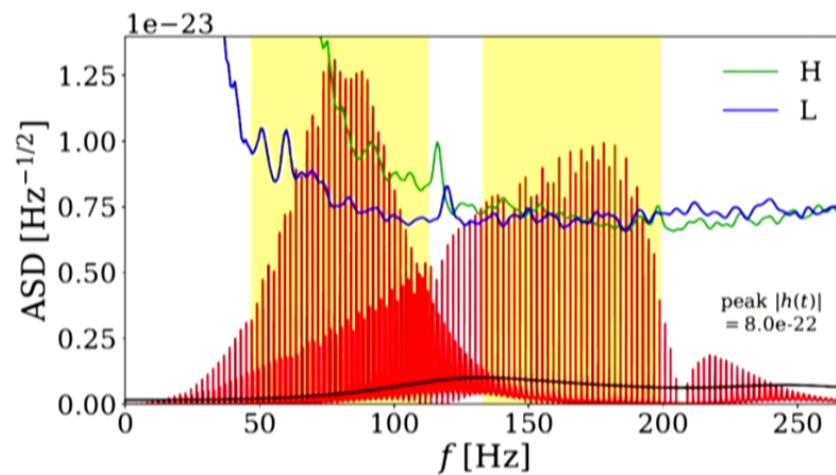
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GW170729

N_E	180/170
Δf	2.0441/2.0443Hz
$\frac{\Delta t}{M}$	1240
χ	0.81

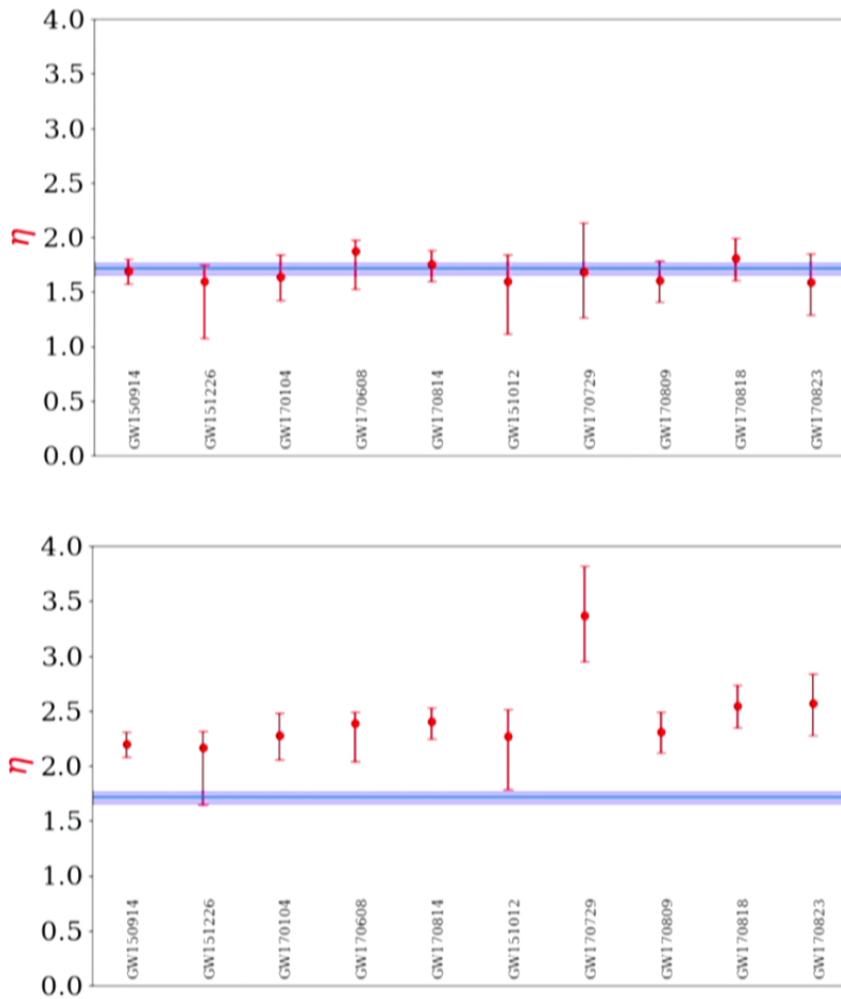


- Δt for a truncated Kerr black hole has known dependence on mass M , spin χ and redshift z

$$\frac{\Delta t}{M} = 4 \eta \log\left(\frac{M}{\ell_{\text{Pl}}}\right) \left(\frac{1 + (1 - \chi^2)^{-\frac{1}{2}}}{2}\right) (1 + z)$$

- a resonance signal determines Δt with negligible error
- LIGO measures M, χ and z of final BH with uncertainties
- thus determine an η for each event
- η determines δr , the distance from the would-be horizon to where the Schd metric applies

$$\delta r \approx \left(\frac{\ell_{\text{Pl}}}{M}\right)^{(\eta-1)} \cdot \ell_{\text{Pl}} \approx \left(\frac{M}{\ell_{\text{Pl}}}\right)^{(2-\eta)} \cdot \bar{\ell}_{\text{Pl}} \quad \begin{aligned} \ell_{\text{Pl}} &\equiv \text{Planck length} \\ \bar{\ell}_{\text{Pl}} &\equiv \text{proper Planck length} \end{aligned}$$



- same as previous plot

- remove the spin and red-shift factors from the formula for η
- draws attention to GW170729

p-values

GW150914	0.011	GW151226	0.014
GW170104	0.33	GW170814	0.098
GW170608	0.038	GW170809	0.081
GW151012	0.0015	GW170823	0.026
GW170818	0.0094	GW170729	0.0010 & 0.0006

- there are two popular methods to combine independent p-values, Fisher and Stouffer
- they give 5×10^{-11} and 2×10^{-12} respectively
- even when adding 0.005 to all 11 p-values, the combined p-values are 1×10^{-8} and 2×10^{-10}

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Conclusions

- LIGO is sensitive to not quite black holes
- this sensitivity to Planck scale physics is under appreciated
- signal waveform may be too variable and complicated for template analysis
- equally spaced resonance pattern is robust and relatively easy to search for
 - there is still a two component structure to explore
- evidence is building from all 10 events
 - consistency of $\Delta t/M$ values with spin and red-shift dependence
 - consistency of bandpasses with predicted spectrum
 - p-values and very small combined p-value