

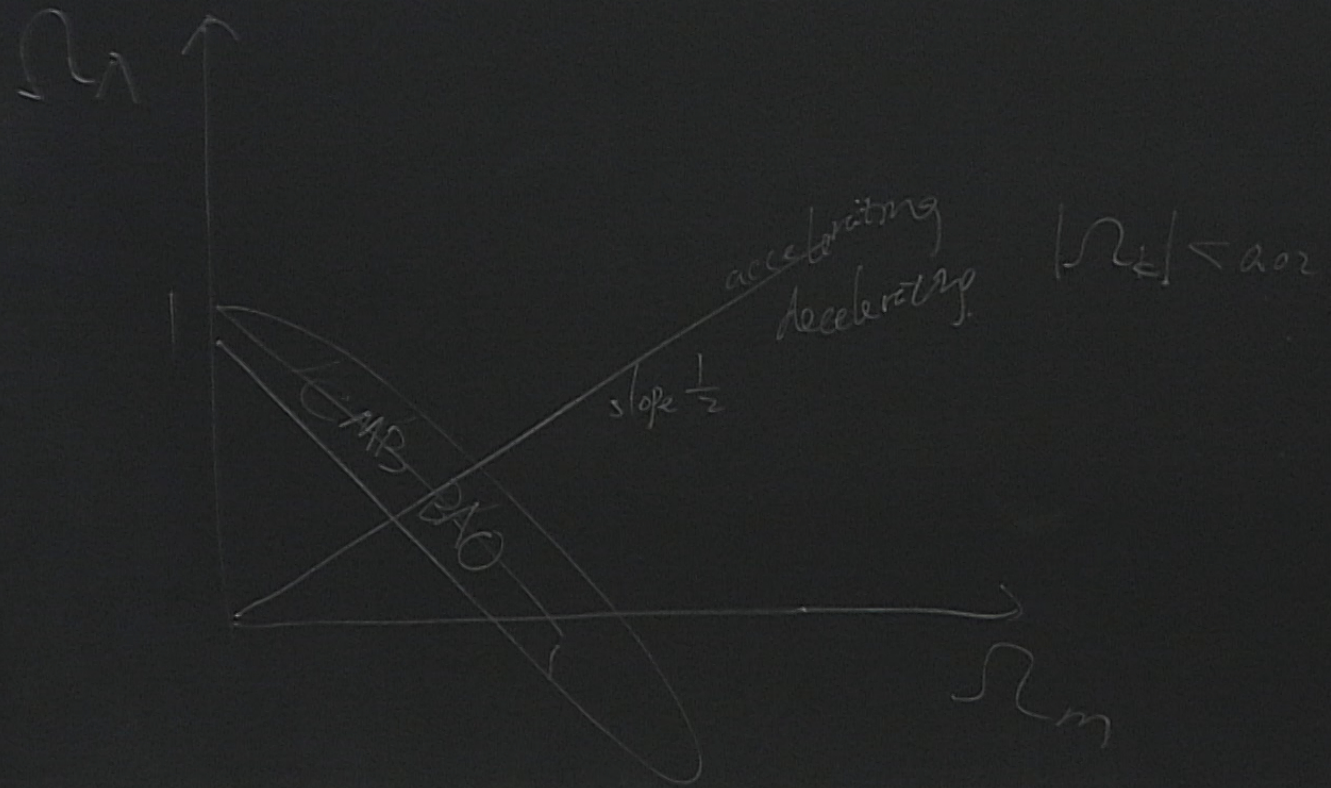
Title: PSI 2018/2019 - Cosmology Review - Lecture 10

Date: Jan 18, 2019 09:00 AM

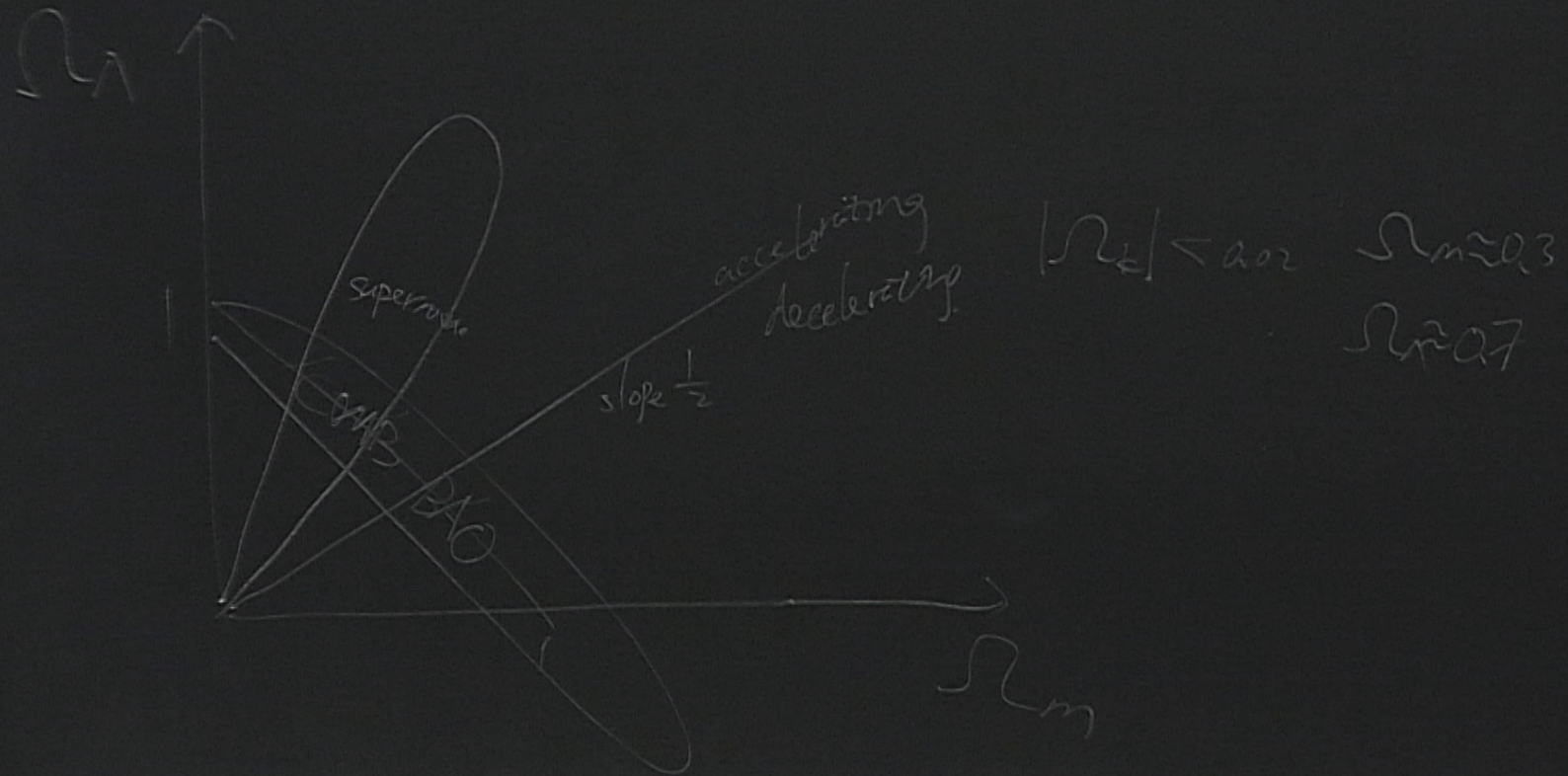
URL: <http://pirsa.org/19010011>

Abstract:

Dark Matter



Dark Matter



Structure formation

baryonic is not
enough

DM starting early

side

MOND

$$\sqrt{g} f(R)$$

side

MOND

$$\sqrt{g} f(r)$$

History

1933 Zwicky

dark matter

measures

velocity

Star + gas $\Rightarrow$ mass

galaxies

big clusters

Zwicky
matter

measures

velocity

Star & gas $\Rightarrow$ mass

Galaxies

N galaxy mass m $M = Nm$ average speed v

$$K = \frac{1}{2} M v^2$$

$$U \approx - \frac{GM^2}{R}$$

Virial theorem

$$K = - \frac{1}{2} U$$

$$I = M r^2$$

$$\ddot{I} = 0$$

Zwicky
matter

measures

velocity

Star & gas $\Rightarrow$ mass

Galaxies

N galaxy mass m $M = Nm$ average speed v

$$K = \frac{1}{2} M v^2$$

$$v^2 \sim - \frac{GM^2}{R}$$

Virial theorem

$$K = -\frac{1}{2} U$$

$$I = M r^2$$

$$\dot{I} = 0$$

$$M v^2 = \frac{G M^2}{R}$$

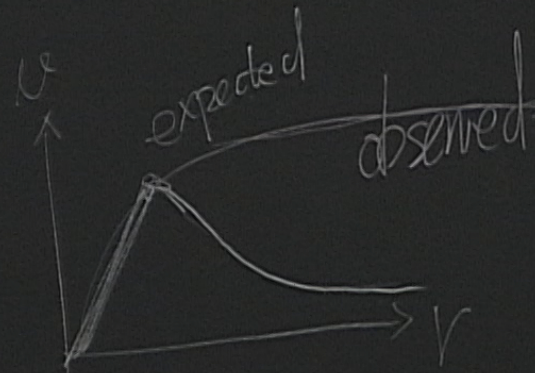
$$M = \frac{v^2 R}{G}$$

off by a factor of 10

$$\Omega_{cluster,0} = 0.2$$

$$\Omega_m > 0.2$$

1973 Vera Rubin



$$\frac{GM}{r^2} = \frac{v^2}{r}$$

$$v \propto \frac{1}{\sqrt{r}}$$

$$\rho \propto r^{-x}$$

$$M \propto \int r^{-x} r^2 dr = \frac{r^{3-x}}{3-x}$$

$$v^2 \propto \frac{M}{r} \propto r^{2-x}$$

$$v \propto r^{1-\frac{x}{2}}$$

$$x=2$$

$$\rho \propto r^{-2}$$

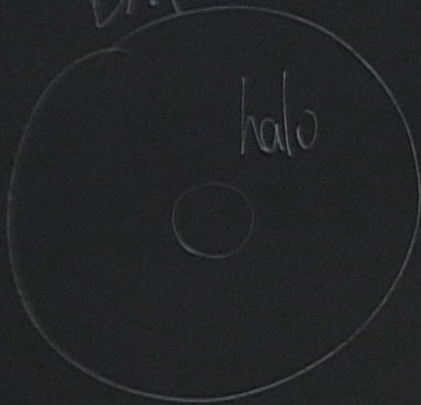
$$\rho_c$$

$$\frac{\text{mass}}{\text{luminosity}}$$

$$\frac{M_0}{L_0} = 1$$

mass $\uparrow$ $L \uparrow$ faster than linear

DM



$$\frac{M}{L} < 1 \quad M > M_0$$

$$\frac{M}{L} > 1$$

$M < M_0$ milky way 2.2

clusters : 200-300

mass
luminosity $\frac{M}{L_0} = 1$ mass $\uparrow$ $L \uparrow$ faster than linear

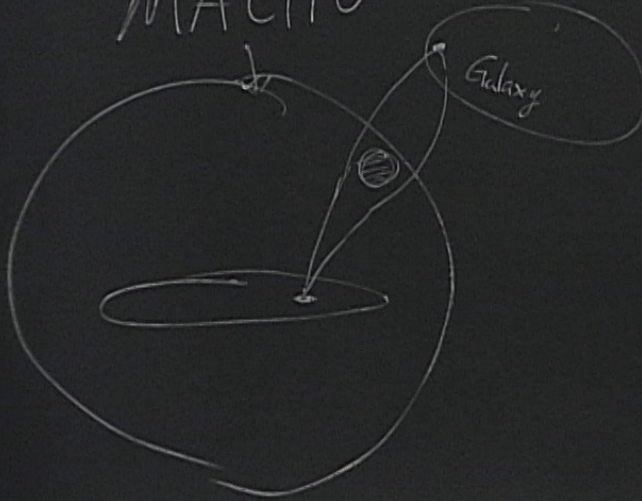
$$\frac{M}{L} < 1 \quad M > M_0$$

$$\frac{M}{L} > 1 \quad M < M_0 \quad \text{milky way } 2.2$$

clusters : 200-300 $0.14 < \Omega_m < 0.45$

BDM

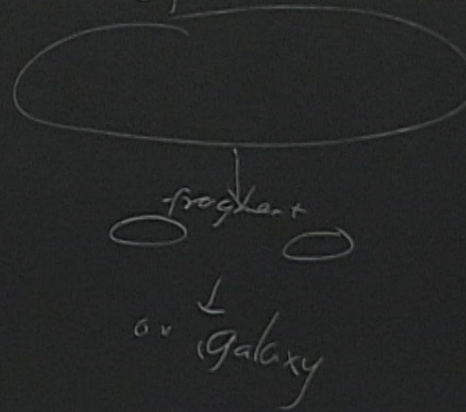
MACHO



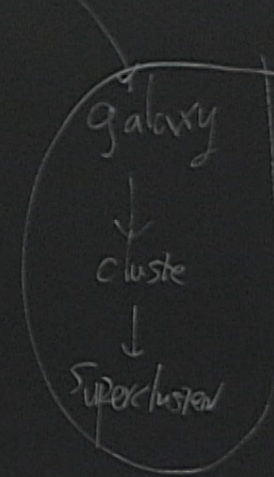
40% disc
70% gas

nonBDM

hot
relativistic
top down



CDM
cold
non relativistic
→ decouple



CDM ~ WIMP

cold

χ ← collider production

$\bar{\chi}$

SM

~ weak scale

~ 100 GeV

SM

scattering
heavy nuclei

→ F_{DM}

$G_{\mu\nu}^a G_{\rho\sigma}^a \epsilon^{\mu\nu\rho\sigma}$

