

Title: Schwarzschild's Interior Solution, Gravastars and Echoes

Date: Nov 10, 2017 09:20 AM

URL: <http://pirsa.org/17110093>

Abstract:

Gravitational Condensate Stars

Negative Pressure & Surface Tension of the Interior Schwarzschild Solution *Echoes, GWs and Imaging*

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w. **P. O. Mazur**

→ **Class. Quant. Grav. 32, 215024 (2015)**

Proc. Natl. Acad. Sci., 101, 9545 (2004)

arXiv:gr-qc/0109035

Acta Phys. Pol. B 41, 2031 (2010)

REVIEW: LPT OF Gravity.

Irreducible Mass and 'No Hair' Theorem

- All Gen. Rel. Black Holes specified by their mass, angular momentum and electric charge: M, J, Q
- Rotating Kerr Black Holes have all higher multipoles *determined completely* by M, J (no "hair")
- Irreducible Mass M_{irr} *increases monotonically classically* (Christodoulou, 1972)

$$M^2 = (M_{\text{irr}} + Q/4M_{\text{irr}})^2 + J^2/4M_{\text{irr}}^2$$

$$M_{\text{irr}}^2 = (\text{Area})/16\pi G \quad \Delta M_{\text{irr}}^2 \geq 0$$

- Smarr Formula: *Purely Classical*

$$dM = \frac{\kappa}{8\pi G} dA + \Omega dJ + \Phi dQ$$

Surface Gravity $\kappa = \frac{1}{4GM} \left\{ 1 - \frac{16\pi^2 G^2}{A^2} (Q^4 + 4J^2) \right\}$



Black 'Holes'... or Not Singularity Theorems

Black Holes 'inevitable' in Gen. Rel. if

- A Trapped Surface forms
- One of Energy Conditions:
- **Weak Energy Condition** (Penrose 1965)

$$\rho + p_i \geq 0 \quad i = 1, 2, 3$$

Violated by Quantum Fields, e.g. by Casimir Effect

- **Strong Energy Condition** (Hawking-Penrose 1970)

$$\rho + \sum_{i=1}^3 p_i \geq 0$$

Violated by Hadronic 'Bag', Cosmological Dark Energy,

Inflation $V(\varphi)$: $p_i = -\rho < 0$

Negative Pressure → Defocusing → Effective Repulsion

Gravitational Vacuum Condensates

- Gravity is a theory of spin-2 **bosons**
- Its interactions are **attractive**
- The interactions become **strong** near $r = R_S$
- Energy of any **scalar** order parameter must couple to gravity with the **vacuum** eq. of state,
$$p_V = -\rho_V = -V(\phi)$$
- Relativistic Entropy Density s is (for $\mu = 0$),
$$Ts = p + \rho = 0 \text{ if } p = -\rho$$
- Zero entropy density for a **single** macroscopic quantum state, $k_B \ln \Omega = 0$ for $\Omega = 1$
- This eq. of state **violates** the energy condition,
$$\rho + 3p \geq 0 \text{ (if } \rho_V > 0)$$
 needed to prove the classical singularity theorems
- Dark Energy acts as a **repulsive** core

A GBEC phase transition can stabilize
a high density, compact cold stellar
remnant to further gravitational collapse

Gravitational Vacuum Condensate Star Proposed (2001)

Today: It's Realized in Schwarzschild Soln.II (1916)

Static, Spherical Symmetry

- 2 Metric Fns.

$$f(r), h(r) \equiv 1 - \frac{2Gm(r)}{r} \quad \leftarrow \text{Misner-Sharp Mass}$$

- 3 Stress Tensor Fns.

$$T^{\mu}_{\nu} = \begin{pmatrix} -\rho & 0 & 0 & 0 \\ 0 & p & 0 & 0 \\ 0 & 0 & p_{\perp} & 0 \\ 0 & 0 & 0 & p_{\perp} \end{pmatrix}$$

- 2 Einstein Eqs.

$$\frac{dm}{dr} = 4\pi r^2 \rho \quad \frac{h}{2f} \frac{df}{dr} = \frac{Gm}{r^2} + 4\pi Gpr$$

- 1 Conservation Eq.

$$\nabla_{\mu} T^{\mu}_{r} = \frac{dp}{dr} + \frac{\rho + p}{2f} \frac{df}{dr} + \frac{2(p - p_{\perp})}{r} = 0$$

Buchdahl Bound (1959)

Assuming classical Einstein eqs. &

- **Static Killing time:**

$$K^\mu \frac{\partial}{\partial x^\mu} = \frac{\partial}{\partial t}$$

- **Spherical Symmetry:**

$$ds^2 = -f(r) dt^2 + \frac{dr^2}{h(r)} + r^2 d\Omega^2$$

- **Isotropic Pressure:**

$$p_i = p(r)$$

- **Positive Monotonically Decreasing Density:** $\frac{d\rho}{dr} \leq 0$

- **Metric Continuity at Surface of Star $r=R$**

- **Then** $R > \frac{9}{8} R_s = \frac{9}{4} GM$

or the pressure must diverge in the Interior

Note this R is outside horizon

Schwarzschild Interior (1916)

- Importance of Buchdahl Bound is:

Under **Adiabatic Compression** **Something Happens**
Inside-- Before the Event Horizon is Reached

- Holds for **Any** isotropic equation of state $p_{\perp} = p$
- Bound is Saturated by Schwarzschild Interior Soln.
- Constant Density

$$\frac{d\rho}{dr} = 0$$

$$\rho(r) = \bar{\rho} \equiv \frac{3M}{4\pi R^3}$$

- Solve for Pressure $p(r)$, Metric Functions $f(r), h(r)$

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Schwarzschild Interior Solution

- Constant Density
 $\rho' = 0$
 Saturates
 Buchdahl Bound

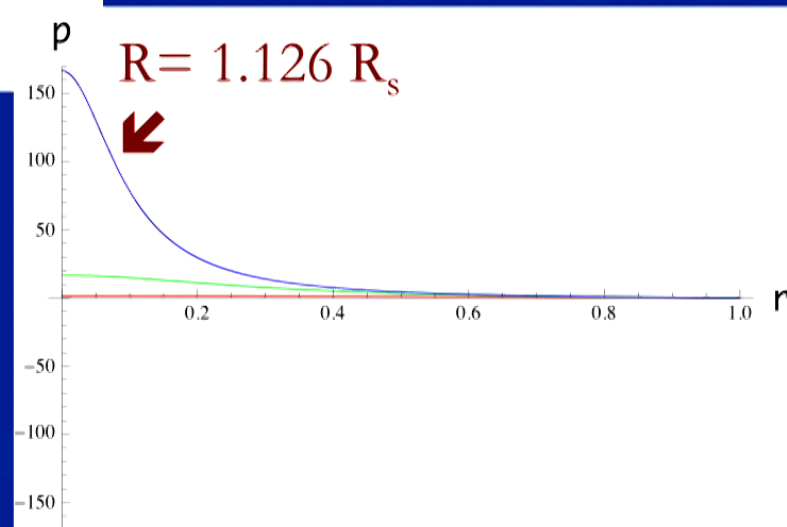
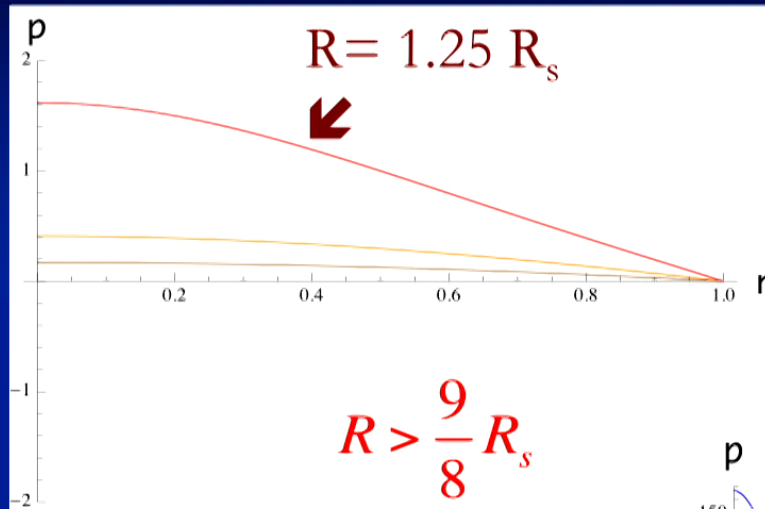
$$m(r) = \frac{4\pi}{3} \bar{\rho} r^3 = \frac{M}{R^3} r^3$$

$$h(r) = 1 - H^2 r^2$$

$$H^2 = \frac{8\pi G}{3} \bar{\rho} = \frac{2GM}{R^3}$$
- Pressure

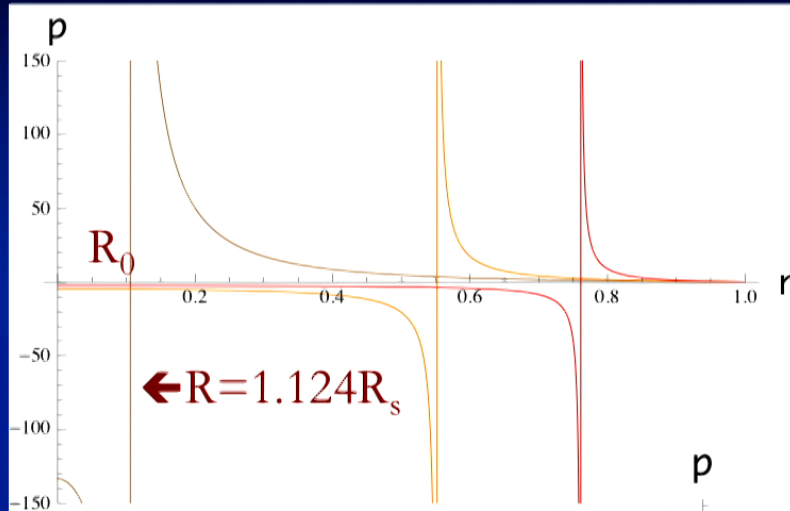
$$p(r) = \bar{\rho} \left[\frac{\sqrt{1 - H^2 r^2} - \sqrt{1 - H^2 R^2}}{3\sqrt{1 - H^2 R^2} - \sqrt{1 - H^2 r^2}} \right]$$
- Diverges at $R_0 = 3R \sqrt{1 - \frac{8}{9} \frac{R}{R_s}}$ iff $R < \frac{9}{8} R_s = \frac{9}{4} GM$
- $f(r) = \frac{1}{4} \left[3\sqrt{1 - H^2 R^2} - \sqrt{1 - H^2 r^2} \right]^2$ Vanishes at R_0
- Pressure becomes negative for $0 < r < R_0$

Interior Pressure



As $R \rightarrow \frac{9}{8} R_s$ from above
central pressure diverges

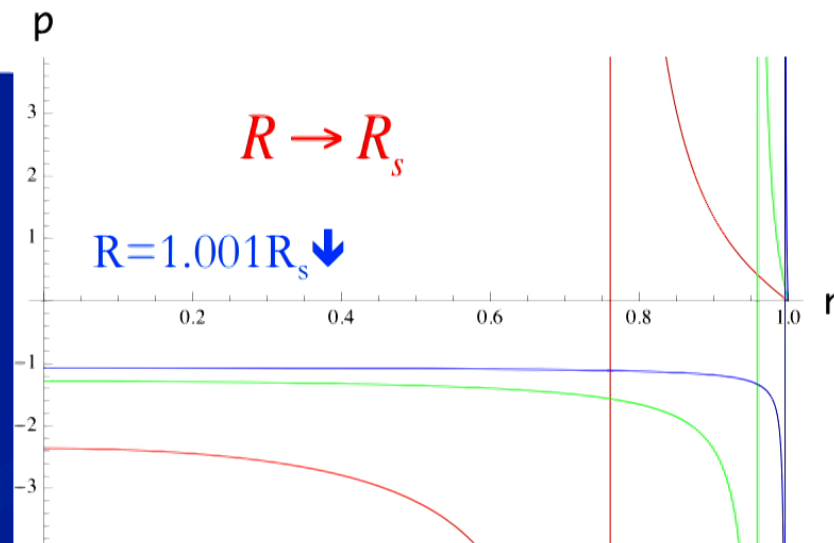
Interior Pressure



$$R < \frac{9}{8} R_s$$

Negative Pressure soln.
opens up for $R < R_0$

As $R \rightarrow R_s$ from above
 $R_0 \rightarrow R_s$ from below and
negative pressure region
fills entire interior with
 $p = -\rho$, $w = -1$



$$R \rightarrow R_s$$

$$R = 1.001 R_s \downarrow$$

Schwarzschild Interior Solution

- Constant Density

$$\rho' = 0$$

Saturates

Buchdahl Bound

$$m(r) = \frac{4\pi}{3} \bar{\rho} r^3 = \frac{M}{R^3} r^3$$

$$h(r) = 1 - H^2 r^2$$

$$H^2 = \frac{8\pi G}{3} \bar{\rho} = \frac{2GM}{R^3}$$

- Pressure

$$p(r) = \bar{\rho} \left[\frac{\sqrt{1 - H^2 r^2} - \sqrt{1 - H^2 R^2}}{3\sqrt{1 - H^2 R^2} - \sqrt{1 - H^2 r^2}} \right]$$

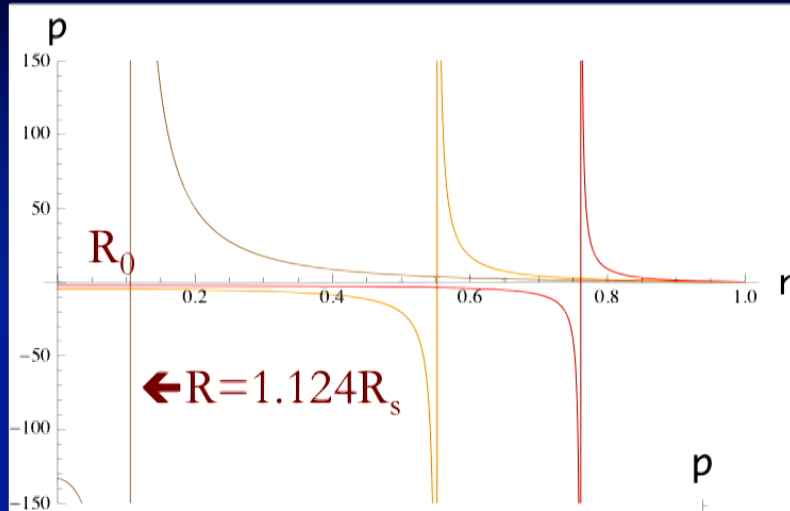
- Diverges at

$$R_0 = 3R \sqrt{1 - \frac{8}{9} \frac{R}{R_s}} \quad \text{iff} \quad R < \frac{9}{8} R_s = \frac{9}{4} GM$$

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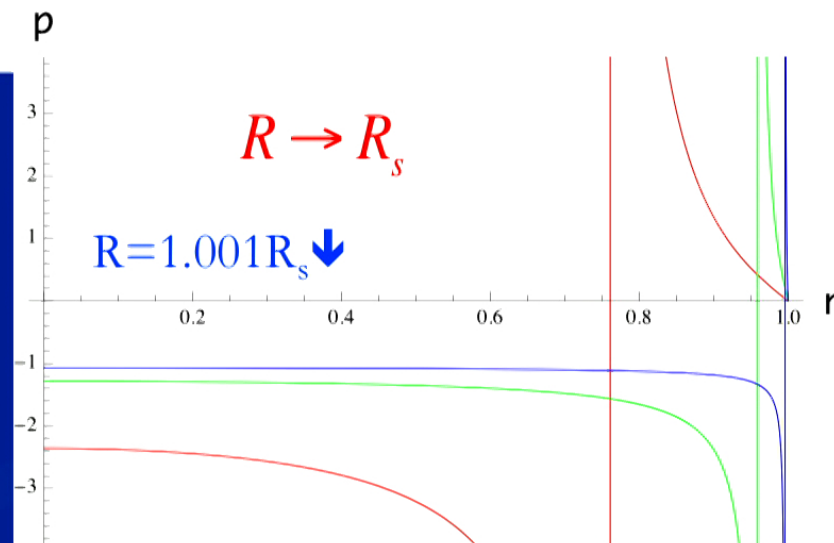
Interior Pressure



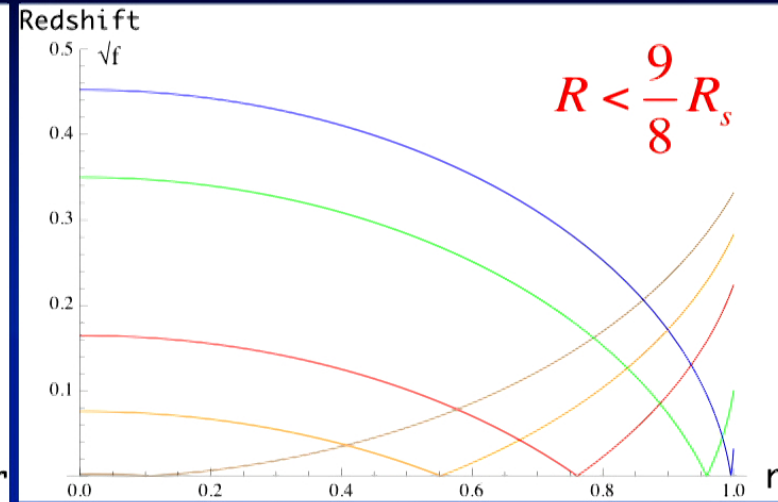
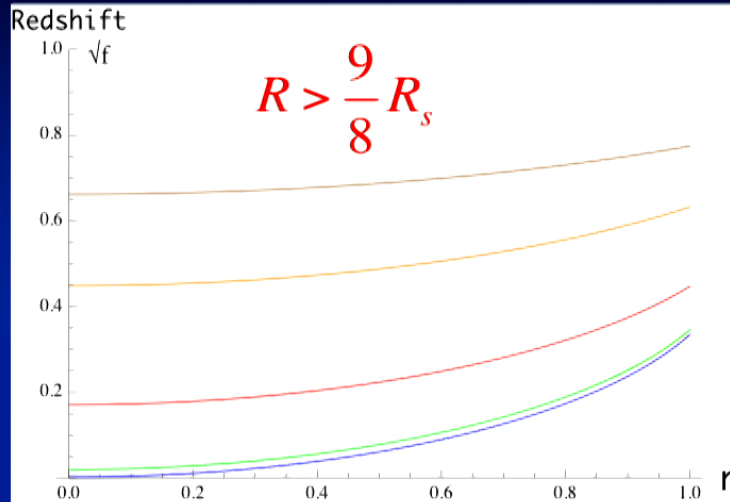
$$R < \frac{9}{8} R_s$$

Negative Pressure soln.
opens up for $R < R_0$

As $R \rightarrow R_s$ from above
 $R_0 \rightarrow R_s$ from below and
 negative pressure region
 fills entire interior with
 $p = -\rho$, $w = -1$



Interior Redshift



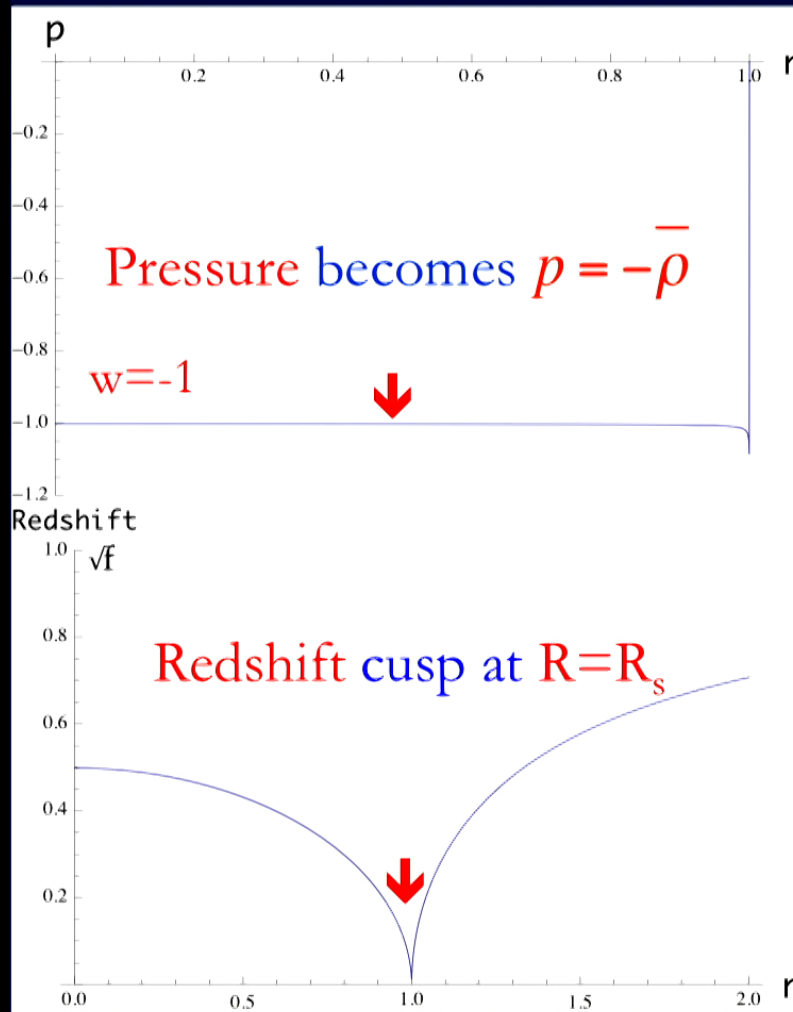
$$f(r) = \frac{1}{4} D^2 \geq 0 \quad \text{Non-negative (no trapped surface)}$$

$D \equiv 3\sqrt{1 - R_s/R} - \sqrt{1 - R_s r^2/R^3}$ vanishes at **same** radius

$$R_0 = 3R\sqrt{1 - \frac{8}{9} \frac{R}{R_s}} \quad \text{where } p \text{ diverges} \rightarrow \text{Integrable}$$

Redshift $\sqrt{f(r)} = \frac{1}{2} |D|$ has cusp-like behavior

$R=R_s$ Limit is Grav. Condensate Star (2001)



No divergence in p

$$p = -\bar{\rho}$$

$$h(r) = 1 - H^2 r^2$$

$$f(r) = \frac{1}{4} h(r)$$

$$H = 1/R_s$$

but non-analytic cusp

Discontinuity (classically)

Interior is de Sitter space in
(modified) static coordinates
(Time runs slower inside)

Komar Mass-Energy Flux (1959-62)

$$\frac{1}{G} \frac{d}{dr} (r^2 \kappa) = 4\pi \sqrt{\frac{f}{h}} r^2 (\rho + p + 2p_{\perp})$$

$$\kappa(r) = \frac{1}{2} \sqrt{\frac{h}{f}} \frac{df}{dr} \rightarrow \frac{GM}{r^2} \quad \text{Surface Gravity}$$

Total Mass: 'Gauss' Law' for Static Gravity

$$M = 4\pi \int_0^{R_s} dr \sqrt{\frac{f}{h}} r^2 (\rho + p + 2p_{\perp})$$

Transverse Pressure

Cusp in Redshift & Step produces Transverse Pressure

$$r \frac{d}{dr} \left[(p + \bar{\rho}) f^{\frac{1}{2}} \right] = 2 (p_{\perp} - p) f^{\frac{1}{2}}$$

δ Localized at $r = R_0$ Surface

$$8\pi \sqrt{\frac{f}{h}} r^2 (p_{\perp} - p) = \frac{8\pi}{3} \bar{\rho} R_0^3 \delta(r - R_0)$$

Integrable Surface Energy

$$E_s = \frac{8\pi}{3} \bar{\rho} R_0^3 = 2M \left(\frac{R_0}{R} \right)^3 \rightarrow 2M$$

$$M = E_v + E_s$$

Surface Tension

Discontinuity in Surface Gravities

$$\kappa_{\pm} \equiv \lim_{r \rightarrow R_0^{\pm}} \kappa(r) = \pm \frac{4\pi G}{3} \bar{\rho} R_0$$
$$\Delta\kappa \equiv \kappa_+ - \kappa_- = \frac{R_s R_0}{R^3} \rightarrow \frac{1}{R_s}$$

is (redshifted) surface tension

$$\tau_s = \frac{E_s}{2A} = \frac{\Delta\kappa}{8\pi G} \rightarrow \frac{1}{8\pi G R_s} > 0$$

Interior is not analytic continuation of exterior

First Law

Classical Mechanical Conservation of Energy

$$dM = dE_v + \tau_s dA$$

Gibbs Relation

$$p + \rho = sT + \mu N = 0$$

Schw. Interior Soln. in $R \rightarrow R_s$ Limit describes

a Zero Entropy/Zero Temperature Condensate

Discontinuity in κ implies non-analytic behavior

No Trapped Surface, Truly Static, t is a Global Time

Surface Area is Surface Area not Entropy

Surface Gravity is Surface Tension not Temperature

Rindler-like C^1 Coordinates

$$ds^2 = -\frac{\xi^2}{4R_s^2} dt^2 + q^2(\xi) d\xi^2 + r^2(\xi) d\Omega^2$$
$$q(\xi) = \begin{cases} \frac{R_s}{r} = \left(1 - \frac{\xi^2}{R_s^2}\right)^{-\frac{1}{2}} & -R_s < \xi \leq 0 \\ \frac{r^2}{R_s^2} = \left(1 - \frac{\xi^2}{4R_s^2}\right)^{-2} & 0 \leq \xi < 2R_s \end{cases}$$

Surface is at $\xi = 0$ $\xi^2 \geq 0$ ξ remains real

Generalizes Israel Jcn. Condition to Null Boundary
Faithful to Einstein Equivalence Principle as originally
conceived for real coordinate transformations

Surface Oscillations

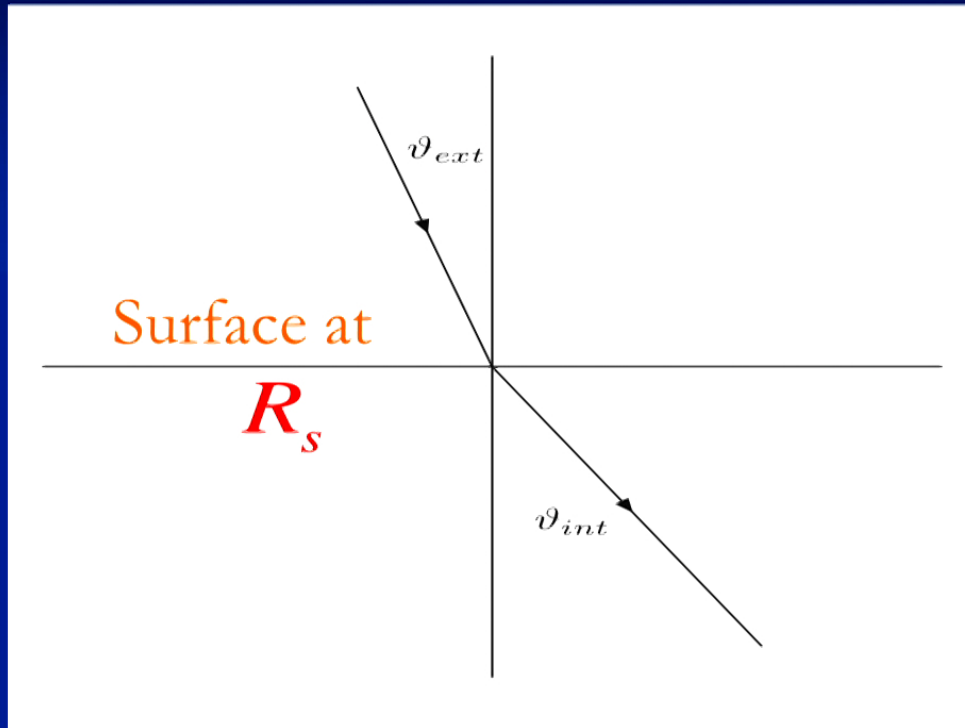
$$dM = dE_v + \tau_s dA$$

- Energy Minimized by minimizing A for fixed Volume
- Surface Tension acts as a restoring force
- Surface Oscillations are **Stable**
- Surface Normal Modes are **Discrete**
- Characteristic Frequency

$$\omega \sim \frac{c}{4\pi R_s} = 8.1 \left(\frac{M_\odot}{M} \right) \text{ kHz}$$

- Discrete Gravitational Wave Spectrum
- Striking Signature for LIGO/VIRGO for
 $M \sim 10^{1-2} M_\odot$

Refraction of Null Rays at Surface



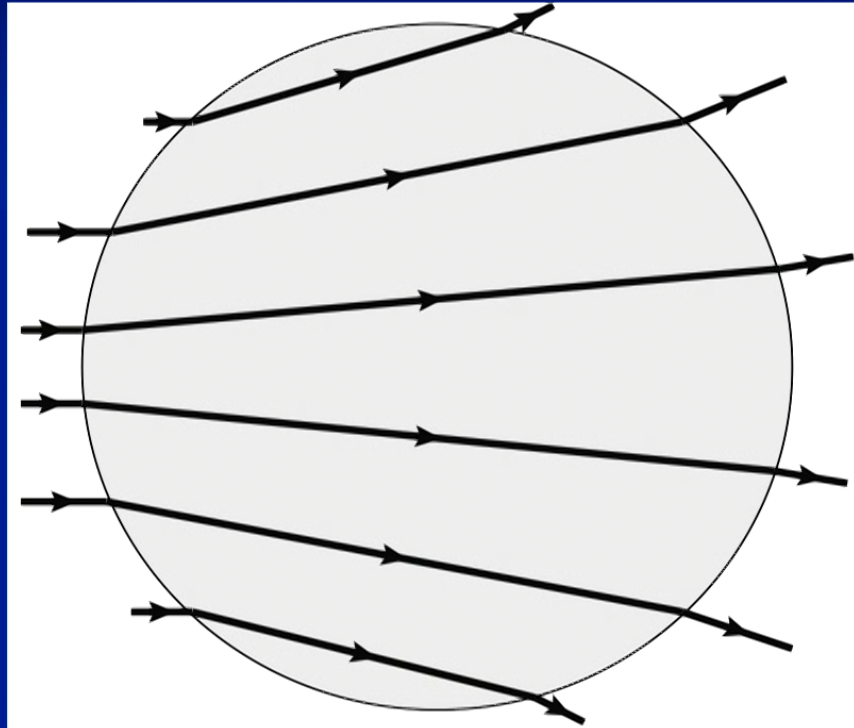
Solve
Geodesic Eq.
→ Snell's Law
 $n < 1$

Impact
Parameter
 b

$$\sin \vartheta_{ext} \sqrt{1 + \frac{4R_s^2}{b^2}} = \sin \vartheta_{int} \sqrt{1 + \frac{R_s^2}{b^2}}$$

Defocusing of Null Rays

No Horizon → Light Rays Penetrate Interior



Completely Different Imaging from a Black Hole

Remarks

- **Non-Analyticity** Characteristic of a Phase Boundary
Freezing of Time \rightarrow Critical Slowing Down
$$-g_{tt} = -K_\mu K^\mu = f(r) = c_{eff}^2 \geq 0$$
- K_μ is everywhere **timelike**: Hamiltonian exists & is **Hermitian** w. proper b.c. at 0 and ∞
- Quantum State of Test Field \sim Boulware State (No Flux)
- Energy Density $p = \frac{\rho}{3} \rightarrow -\frac{C}{M^4 f^2} < 0$
- **Violates** Both Weak & Strong Energy Conditions
- **Significant Backreaction** when $f(r) \geq \epsilon \sim \frac{L_{pl}}{R_s}$
- Occurs at **Physical Length**
- Also time delay $\sim \ln(\epsilon)$
- All **Regulated** by finite ϵ

$$\ell \sim \frac{\Delta r}{\sqrt{h(r)}} \sim \sqrt{L_{pl} R_s}$$

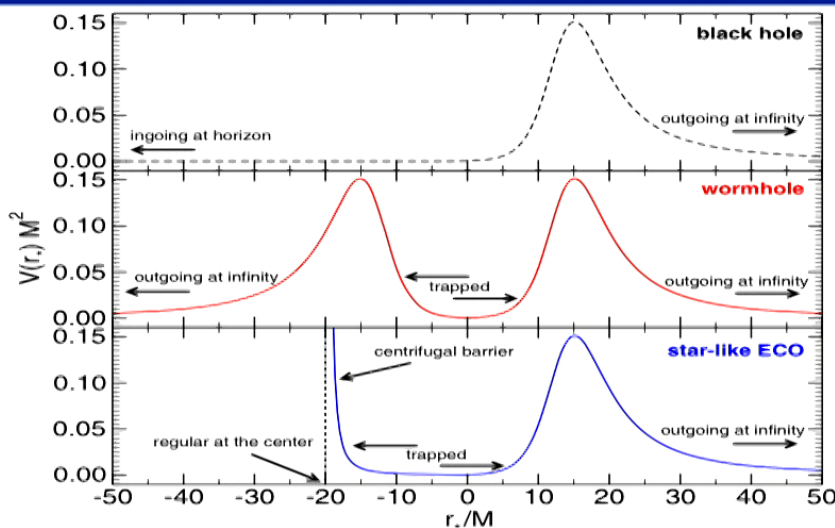
Gravastar Interior Echoes

- Transmission (with log time delay) through surface
- Regge-Wheeler radial coordinate
- Scalar Wave or Grav. Wave Eq.

$$dr^* = \frac{dr}{\sqrt{fh}}$$

$$\left(-\frac{d^2}{dr^{*2}} + V_\ell\right) \phi_{\omega\ell} = \omega^2 \phi_{\omega\ell}$$

- Reflection from de Sitter interior barrier: 'Echo'



$$V_\ell = \frac{1}{4} (1 - H^2 r^2) \left(\frac{\ell(\ell+1)}{r^2} - 2H^2 \right)$$

$$r < 2M = H^{-1}$$

$$V_\ell = \left(1 - \frac{2M}{r} \right) \left(\frac{\ell(\ell+1)}{r^2} + \frac{2M}{r^3} \right)$$

$$r > 2M$$

Summary

- **Buchdahl Bound** → Interior Pressure Divergence Develops (inside out) before Event Horizon forms
- Constant Density Interior Schwarzschild Solution $R > \frac{9}{8} R_s = \frac{9}{4} \frac{GM}{c^2}$ **Saturates Bound** & illustrates the generic behavior
- Pressure Singularity is **Integrable-Sensible Boundary Layer**
- Implies Formation of a **δ -fn. Surface** & **Surface Tension** & a **Non-Singular** (de Sitter) Interior
- **Gravitational Condensate Star** $p = -\rho$ **negative pressure** already realized/inherent in Classical Gen. Rel. (1916)
- **Cold Quantum Final State** of Gravitational Collapse
- No Thermal Radiation/ Not a Firewall

Summary

- Area Term is Classical **Mechanical Surface Energy**
not Entropy
- QM, Unitarity ✓ **No 'Information Paradox'**
- **No large number of microstates: One Vacuum State**
- Both Echoes and Discrete Surface Modes
- **Transmission/Refraction** → different imaging for EHT
- Has been extended to slow rotation C. Posada-Aguirre,
MNRAS 468 (2017) 2128
- Full **Non-Singular Soln.** Requires Quantum
Effective Theory of the Conformal Anomaly
- **Dynamical Vacuum Condensate Energy**
- Regulated **Finite Thickness** Boundary layer