Title: Was Einstein Right Handed?

Date: Nov 25, 2009 02:00 PM

URL: http://pirsa.org/09110132

Abstract: Of all four forces only the weak interaction has experimentally exhibited parity violation. At the same time observations suggest that general relativity may require modification to account for dark matter and dark energy. Could it be that this modification involves gravitational parity violation? Many of the dominant approaches to quantum gravity, such as string theory and loop quantum gravity, point to an effective parity violating extension to general relativity known as Chern-Simons General Relativity (CSGR). In this colloquium I will discuss the uniqueness and phenomenological implications of parity violation in CSGR. In particular, I will discuss how CSGR can work together with inflation to generate the cosmic baryon asymmetry in a most economical fashion, through gravitational waves; and its predictions for upcoming CMB polarization experiments. I will also discuss current predictions of CSGR on binary pulsars, neutron stars and prospects for the LISA/LIGO gravitational wave detectors. While we focus on a specific theory for concreteness, some of the results presented in this colloquium can be seen as model independent.

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WAS EINSTEIN LEFT-HANDED?

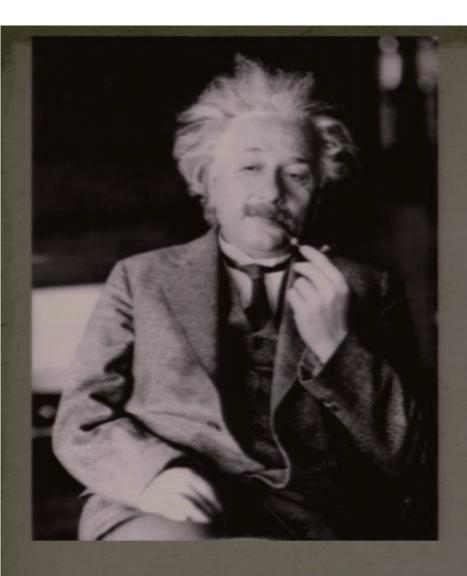
Stephon Alexander Haverford College

&

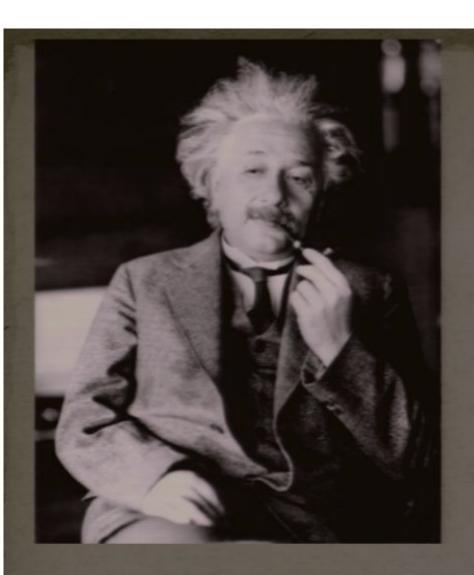
The Institute for Gravitation and the Cosmos Penn State

> Perimeter Institute Colloquium

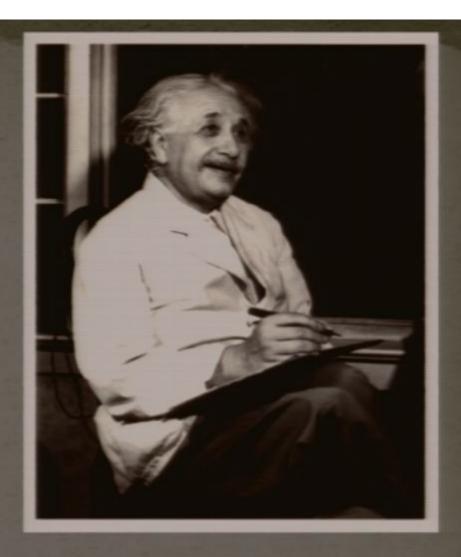




Left-handed !!



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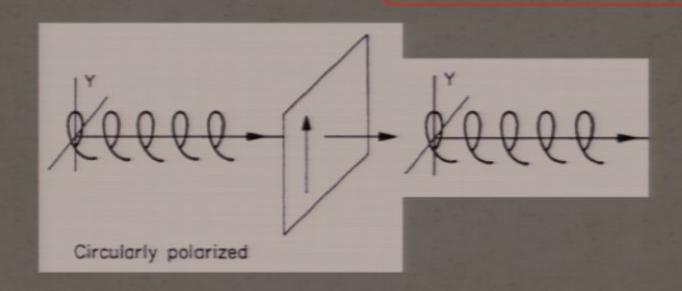
Right-handed?

$$S \sim \int (E^2 + B^2)$$

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P[B] = -B but P[S] = + S !!

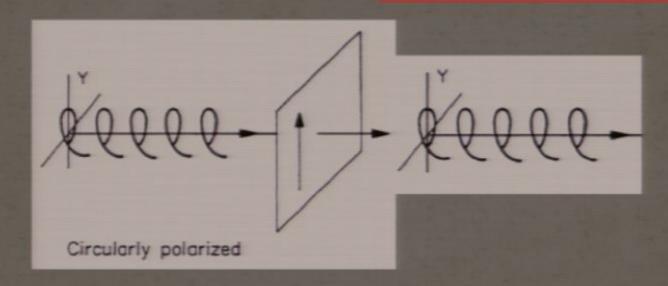
Maxwell is ambidextrous



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A right-polarized EM wave tays right-polarized under reflection

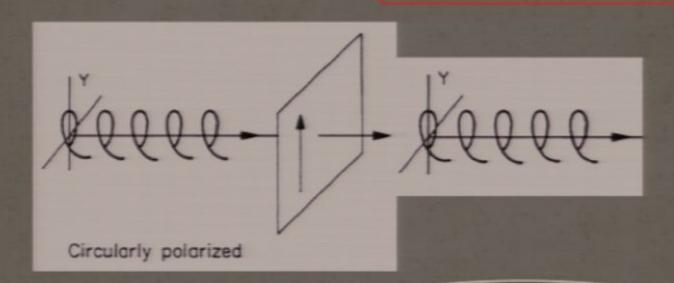
Obviously, since



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Obviously, since $ec{E} \cdot ec{B} = 0$

$$S \sim \int (R + \mathcal{L}_M) \rightarrow G_{ab} \sim T_{ab}$$

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The Levi-Civita tensor would break Parity, but what contraction can we construct that is non-vanishing?

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But what happens if we add this term to the action?

IMPORTANT FACTS:

- Chern-Simon's General Relavitivty is motivated by
- The standard-model (Adler, Bell, Jackiw; Salam, Isham, Duff, Deser)
- String Theory (Green Schwarz, S.A, Gates)
- Loop Quantum Gravity (Smolin; Nesti Peracci, S.A, Perez, Friedel, Mercuri)
- The Only modified Gravity theory known to have specific signatures of parity violation and circular polarization of GW generation/propigation
- LET'S EXPLORE...

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Part I: The Action

 $*RR \sim e^{...}R....R.$ is the Pontryagin density.

$$R^{\beta}_{\ \alpha}^{\ \gamma\delta}\left(\frac{1}{2}\epsilon_{\sigma\tau\gamma\delta}R^{\alpha}_{\ \beta}^{\ \sigma\tau}\right)$$
 Dual of Riemann tensor

$$S \propto \int R + \theta *RR + (\nabla \theta)^2 + V[\theta]$$

- $: \theta$ is either
 - a fixed, externally prescribed par.
 - a dynamical field that evolves.

Non-Dynamical Framework

Dynamical Framework

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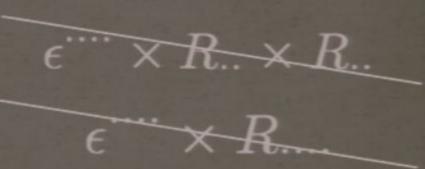
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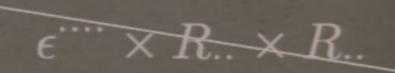
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Road Map

- Chern-Simons Basics:
- Cosmology: BaryogenesisChern-Simons Observables:
- Chern-Simons Tests:

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Dynamical Framework

$$G_{ab} + C_{ab} = \kappa T_{ab}$$

• "C-tensor":

$$C^{\cdots} \sim (\nabla . \theta) \epsilon^{\cdots} \nabla . R^{\cdot} . + (\nabla . \nabla . \theta) \epsilon^{\cdots} R^{\cdots} .$$

$$C^{\mu\nu} = \frac{1}{2} \left[(\partial_{\sigma}\theta) \left(\epsilon^{\sigma\mu\alpha\beta} \nabla_{\alpha} R^{\nu}_{\beta} + \epsilon^{\sigma\nu\alpha\beta} \nabla_{\alpha} R^{\mu}_{\beta} \right) + \nabla_{\tau} (\partial_{\sigma}\theta) \left(\tilde{R}^{\tau\mu\sigma\nu} + \tilde{R}^{\tau\nu\sigma\mu} \right) \right]$$
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In non-dyn framework

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- In non-dyn framework *RR = 0
- Theta evolution preserves SEP: $\nabla^{\cdot}T..=0$
- Pirsa: 09110132 does this modification change any GR results? Page 26/80

Basics: Stringy Realization :10D string-action for Heterotic String $S \propto \int R + H_{(3)}^2$

$$*dB = d\vartheta$$

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Basics: Stringy Realization

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H is the "Kalb-Ramond" three-form, $H_{(3)} = dB_{(2)} + \omega_{(3)}$ where B is a 2-form field and w is the Chern-Simons 3-form: *dH = *dw = R*R

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Compactification to 4D leads to (* $dB=d\vartheta$)

$$S \sim \int d^6y \int d^4x \ (dB_{(2)} + w_{(3)})^2$$

 $\sim \int d^4x \ [dB \wedge *dB + *dB \wedge w + \cdots]$
 $\sim \int d^4x \ [(\partial \vartheta)^2 - \vartheta \ R^*R]$ [S.A.S.) Gates,

Part II: Chern-Simons Gravity and

Leptogenesis

SA, Peskin, S Jabbari PRL 05, S. Weinberg 08, Lyth et. al 06

In largest scales we see no anti-galaxies(matter)

$$\frac{n_B}{n_\gamma} = (6.5 \pm 0.4) \times 10^{-10}$$

lucleosynthesis and WMAP measure

Baryon/Lepton violation CP Violation

The must happen out of equilibrium

A QUICK DETOUR....

R

$$S = \int d^4x \sqrt{-g} \left(-\frac{1}{16\pi G} R + \frac{\ell}{12} \theta R \tilde{R} \right)$$
 Action

$$\begin{split} R\tilde{R} &= \frac{4i}{a^3} \left[\left(\partial_z^2 h_R \ \partial_z \partial_t h_L + a^2 \partial_t^2 h_R \ \partial_t \partial_z h_L \right. \right. \\ &\left. + \frac{1}{2} \partial_t a^2 \partial_t h_R \ \partial_t \partial_z h_L \right) - \left(L \leftrightarrow R \right) \right] \end{split}$$

$$\Box h_L = -2i\frac{\Theta}{a}\dot{h}_L' \; , \qquad \Box h_R = +2i\frac{\Theta}{a}\dot{h}_R'$$

Birefringent Gravity Waves (Circular Dichrohism)

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Birefringent Gravity Waves (Circular Dichrohism)

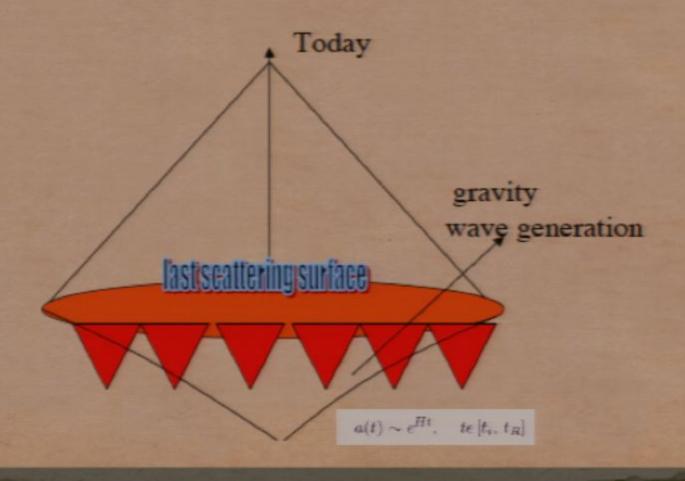
Question: Can nature make use of primordial gravitational waves?

- Scalar perturbations source the gravitational potential for structure formation.
- But we still need asymmetry in matter over antimatter (otherwise we'll have anti-galaxies).

We argue that birefringent gravity waves source matter asymmetry, giving them a physical role in the early universe.

Inflation Basics

Inflation is an epoch wherein
The scale factor has positive acceleration



The Mechanism

Inflaton sourced GW production: (CP) "CP Violation in Inflaton's phase Inflation (B) Gravitational Anomaly Rapid Expansion: Pirsa: 09110132 -> lepton production

Gravity Waves and Chirality

Left and Right handed GWs can 'stir' leptons Out of the empty vacuum

Inflation amplifies this Process

OPPORTUNITY:
CMB Polarization may
Detect this event.



We Need one Important Ingredient

Lepton Number Violation

Gravitational Chiral Anomaly

In the interaction the global lepton current is classically conserved

$$\partial_{\mu}J_{\mu 5} = \partial_{\mu}\bar{\Psi}\gamma_{\mu}\gamma_{5}\Psi = 0$$

One Loop Graviton Quantum correction ...

$$\partial_{\mu}J_{l}^{\mu}=rac{1}{16\pi^{2}}R ilde{R}$$
 $J_{i}^{\mu}=ar{l}_{i}\gamma^{\mu}l_{i}+ar{
u}_{i}\gamma^{\mu}
u_{i}$ Page 42/80

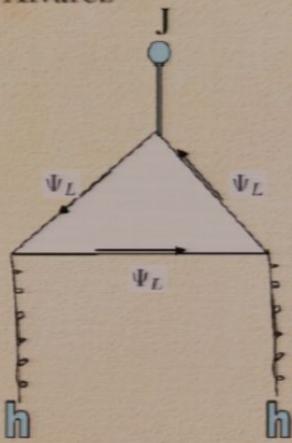
Chiral-Current Anomaly

Duff, Deser, Isham 88; Alvarez-

Gaume, Witten 90

Alvarez-Gaume, Nelson, '85

Cardoso, Ovrut,



Hence the quantum expectation value:

$$\langle R\tilde{R} \rangle = \frac{16}{a} \int \frac{d^3k}{2\pi^3} \frac{H^2}{2k^3 M_{Pl}^2} (k\eta)^2 . k^4 \Theta$$

We pick up only the leading behavior for $k\eta >> 1$ Which corresponds to UV Sub-Horizon modes

A reminder: The above expression is non-zero because of the effect of inflation in producing CP asymmetry out of equilibrium.

WE ARE FINALLY READY TO COMPUTE LEPTON NUMBER:)

DO WE GET ENOUGH?

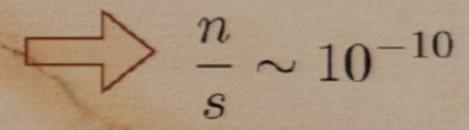
SEMI-FINAL RESULT

We arrive at the final result for the baryon to entropy ratio. We can find μfor a range of Hubble that is acceptable by CMB constraints.

$$\frac{n}{s} \sim 1 \times 10^{-5} \cdot \left(\frac{H}{M_{\rm Pl}}\right)^{-1/2} \left(\frac{\mu}{M_{\rm Pl}}\right)^{5}$$
$$10^{-30} \lesssim H/M_{\rm Pl} < 10^{-4}$$

The range is $3 \times 10^{14} < \mu \lesssim 10^{17}$

Which is the scale of the right handed neutrino! From this we can get



Caveat: Still some fine tuning Page 45/80

CMB Polarization

 Lue, Wang and Kamionkowski showed that birefringent GW lead to parity violtating cross-correlations

$$C_l^{XX'} \equiv \langle a_{(lm)}^X (a_{(lm)}^{X'})^* \rangle$$

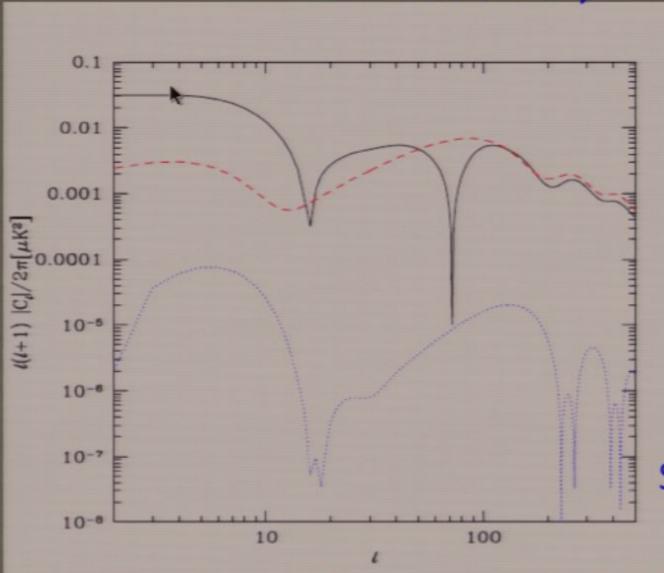
$$X = \{T, G, C\}$$

$$C_l^{TC} = 2(2l+1)^{-1}A_l^T(k)A_l^C(k)$$

$$a_{(lm)}^{T} = \begin{cases} (\delta_{m,2} + \delta_{m,-2}) A_l^T(k) & \text{even l} \quad (+), \\ -i(\delta_{m,2} - \delta_{m,-2}) A_l^T(k) & \text{odd l} \quad (\times), \end{cases}$$

$$a_{(lm)}^{C} = \begin{cases} (\delta_{m,2} + \delta_{m,-2}) A_l^C(k) & \text{even l} \quad (\times), \\ -i(\delta_{m,2} - \delta_{m,-2}) A_l^C(k) & \text{odd l} \quad (+), \end{cases}$$

TB > BB Generically



Black: TB

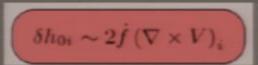
Red: BB

Blue: EB

Contaldi, Majueijo Smolin(PRL '08)

This signal will be searched for in BICEP &
Pirsa: 09110132CK sattellite data (Keating, Miller Shimon; Partridge (Page 47/80 at e

Part III: Testing Chern-Simons Gravity in our solar system



LEADS TO FRAME DRAGGING OF GYROSCOPES

CHERN SIMONS PREDICTS A COMPLETELY NEW PPN PARAMETER

 $\delta h_{0i} \sim 2\dot{f} \left(\nabla \times V \right)_i$

LEADS TO FRAME DRAGGING
OF GYROSCOPES

Introducing Gravitomagnetism

Start with a perturbed metric:
$$g_{\mu
u} = \stackrel{\it Flat}{\eta_{\mu
u}} + \stackrel{\it Small perturbation}{h_{\mu
u}}$$

Define the vector potential:
$$A_{\mu} \equiv -\frac{1}{4} \left(h_{0\mu} - \frac{1}{2} \eta_{0\mu} h
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Define the mass current density:
$$J_{\mu} \equiv -T_{\mu 0} = (- \stackrel{\star}{\rho}, \vec{J} \)$$

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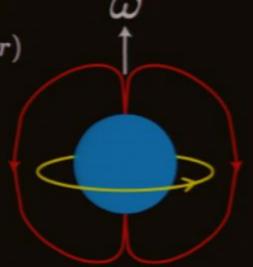
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abla} imesec{B}-rac{\partial E}{\partial t}-rac{1}{m_{cs}}\Boxec{B}=4\pi Gec{J}$$
Standard GR New parity-
violating terms
 $m_{cs}\equivrac{-3}{8\pi G\ell\dot{ heta}}$

Solving the modified Ampère's Law and imposing continuity of the vector potential yields

$$\vec{B} = \vec{B}_{GR} + \vec{B}_{CS}$$

 $m{\theta}_{CS}$ is ocsillatory: $B_{CS} \propto y_{1,2}(m_{cs}r)$

• While \vec{B}_{GR} is purely poloidal, \vec{B}_{CS} has poloidal and toroidal components. Toroidal fields violate parity.



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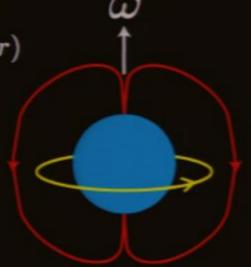
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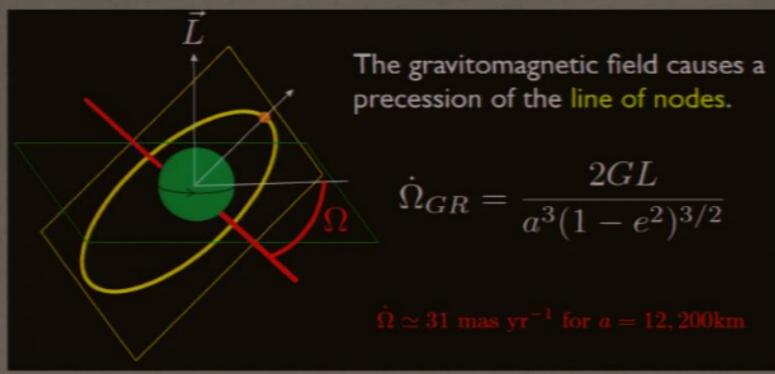
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Modified Precession





$$\left(\vec{B} = \vec{B}_{GR} + \vec{B}_{CS}\right) \Rightarrow \left(\dot{\Omega} = \dot{\Omega}_{GR} + \dot{\Omega}_{CS}\right)$$

Kamionkowski, Erickcek, Smith PRD 08, S.A., Yunes 08 PRL

Generically, stationary, spinning solutions are CS modified in the gravitomagnetic sector only! $g_{ti} = g_{ti}^{(GR)} + g_{ti}^{(CS)}$

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Such a correction leads to a modification in the geodesic equations $\vec{a} = -\vec{k} - 4\vec{v} \times \vec{B}$

A modified acceleration vector produces an anomalous change in orbital precession

Generically, stationary, spinning solutions are CS modified in the gravitomagnetic sector only! $g_{ti} = g_{ti}^{(GR)} + g_{ti}^{(CS)}$

Such a correction leads to a modification in the geodesic equations $\vec{a} = -\vec{k} - 4\vec{v} \times \vec{B}$

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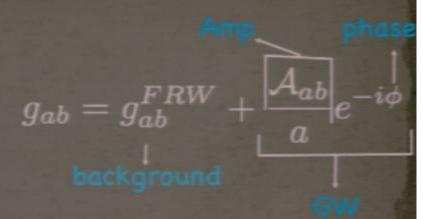
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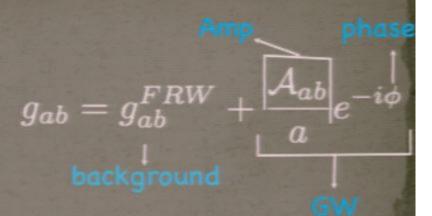
Smith et. al, RD 77 (2008), Sa.A, Yunes PRL (2008)]

$$\frac{\langle \dot{\omega}_{CS} \rangle}{\langle \dot{\omega}_{GR} \rangle} \sim \frac{r^2}{R^2} \frac{\dot{\theta}}{R} \qquad \begin{array}{c} \text{distance to source} \\ \text{source} \end{array}$$

Focus on the propagation of GWs:

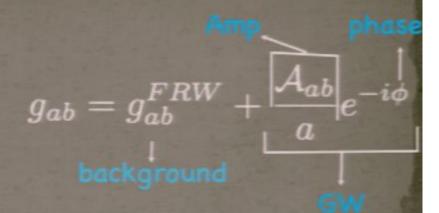


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Modified Dispersion Relation



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$$\Omega_{R,L}^2 \sim \kappa^2 + i \; S_{R,L}[\partial \theta]$$
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GWs still travel at the speed of light, but their amplitudes obey different evolution equations.

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Tests: Solar System

Most Solar System tests passed (Schw. still a sol.)

Kerr is not a solution!

LAGEOS (or GP B) can search for frame-dragging effects:

- LAGEOS (I&II) are laser-ranged satellites, whose orbit can be mapped accurately.
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(SA, Finn & Yunes, RD 78 (2008)1

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Pirsa: 09110132

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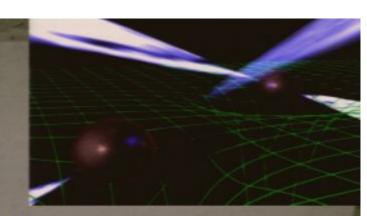
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[S.A & Yunes, PRL 99 (2007), Kamionkowski, Smith et. al, PRI 77 (2008)]

First test of local, non-dynamical CS effects for canonical CS field

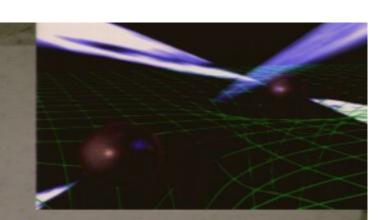
Binary pulsars, sensitive to precession.



CS correction larger for binary pulsar because:

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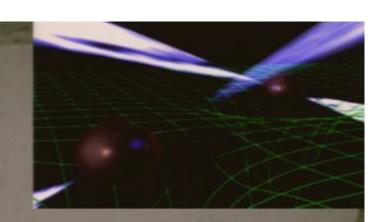
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AGEOS:

$$\frac{R_+ + h}{R_+} \sim 1$$
Earth
Radius

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Satellite altitude $\frac{R_+ + h}{R_+} \sim 1$ Earth Radius

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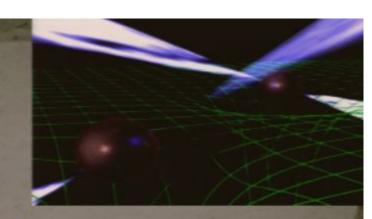
Binary Pulsar:

NS-NS

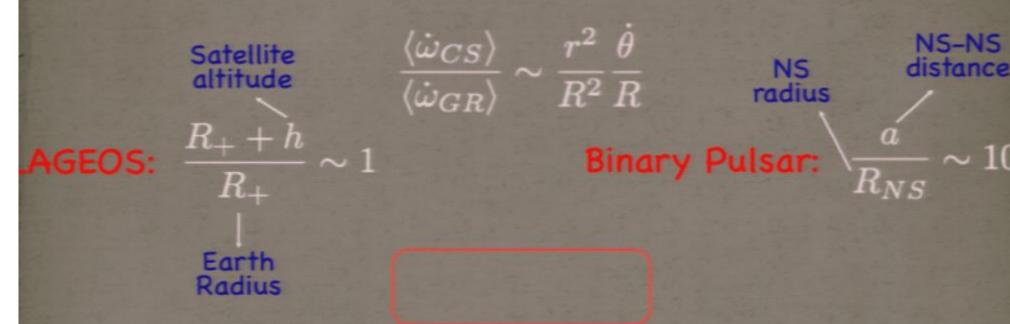
$$\frac{a}{R_{NS}} \sim 10^4$$

AGEOS:

Binary pulsars, sensitive to precession.



CS correction larger for binary pulsar because:



Binary Pulsar test improves Solar System one by at least eleven orders of magnitude!!

Recent improvement from CS neutron stars Upcoming work: Loeb, Yunes et al

LISA is capable of detecting GWs from EMRIs far away.



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mplitude Birefringence Test:

Could measure evolution history of CS scalar, one order of magnitude better than Solar System experiments.



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- Only avenue to test dynamical CS theory is the strong field because dynamical CS field highly 1/r suppressed.

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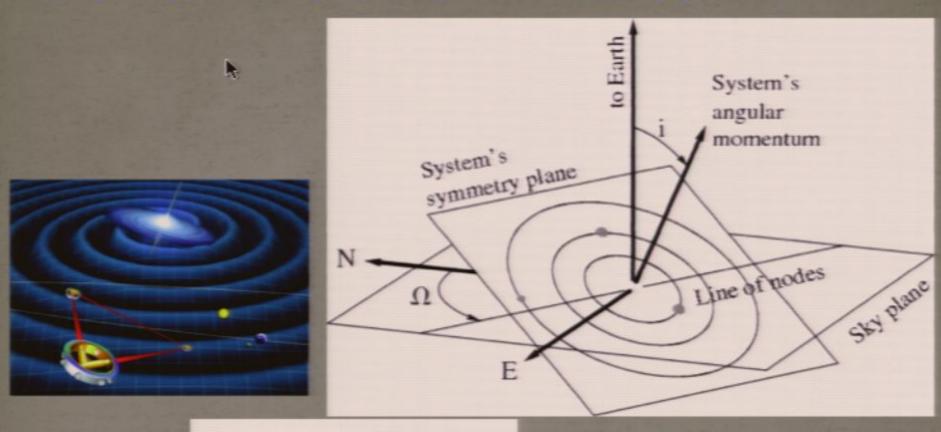
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Pirsa: 09110132 A could be a probe into effective quantum gra Page 77/80

LISA SPACE SATTELITE TEST CONT'D



$$rac{h_{
m R}}{h_{
m L}} \; = \; rac{h_{
m R}^{GR}}{h_{
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ight] \sim \cos \iota + rac{k_0(t)\xi(z)}{H_0} \; \sin^2 \iota \; .$$

where we have defined the redshift-dependent CS form-factor $\xi(z)$ via

$$\xi(z) = \frac{H_0^2}{2} \int_0^z dz \, (1+z)^{5/2} \left[\frac{7}{2} \frac{d\theta}{dz} + (1+z) \frac{d^2\theta}{dz^2} \right]$$

UPCOMING WORK

- •We must understand how GWs behave in dynamical CS, thus, define the following program:
 - . I. Find a CS modified relevant background.
 - · II. Calculate the modified Quadrupole formula.
 - III. Study EMRIs in this background.
- : IV. Study Binary BH mergers in full generality :(Program being carried out at Princeton by Pretorius and Yunes) So far, completed (I), a/M<<1 & $\xi/M^4\ll 1$

$$ds^2 \sim ds_{\rm Kerr}^2 + \frac{5}{8} \frac{\xi}{r^4} \left(1 + \frac{12}{7} \frac{M}{r} + \frac{27}{10} \frac{M^2}{r^2} \right) \sin^2 \theta dt d\phi,$$

$$\theta \sim \frac{5}{8} \sqrt{\xi} \frac{a}{M} \frac{\cos(\theta)}{r^2} \left(1 + \frac{2M}{r} + \frac{18M^2}{5r^2} \right)$$

- •CS gravity is an effective theory of quantum gravity that encodes gravitational parity violation.
- CS affects: frame-dragging and gravitational waves (circular-dichrohism)
- Frame-dragging tests (Solar Sys., Bin. Pul.) have strongly constrained the non-dynamical model.
- Gravitational waves could constraint the dynamical framework.