Title: Inflationary Constraints on String Theory

Date: Jun 06, 2008 12:00 PM

URL: http://pirsa.org/08060185

Abstract: It is an important task to embed inflation in a fundamental microphysical theory such as string theory. Since string theory possesses a vast landscape of 4-dimensional theories, we would like to know which portions contain inflation and which do not. I prove a no-go theorem that inflation and de Sitter vacua are forbidden in an exponentially large number of infinite families of simple and well understood compactifications of type IIA string theory. I also mention more complicated and less well understood compactifications, which may have the ingredients for our cosmology.

Pirsa: 08060185 Page 1/61

Inflationary Constraints on String Theory

PASCOS 08

Mark Hertzberg, MIT

- astro-ph/0709.0002: MH, Tegmark, Kachru, Shelton, Ozcan
- hep-th/0711.2512: MH, Kachru, Taylor, Tegmark

Pirsa: 08060185 Page 2/61

Pirsa: 08060185 Page 3/61

- Cosmologists are asking questions like:

What is the inflaton potential?

What is the inflaton/s?

Pirsa: 08060185 Page 4/61

- Cosmologists are asking questions like:

What is the inflaton potential?

What is the inflaton/s?

- String theorists are asking questions like:

Where in the space of string models is our universe?

What can be predicted?

Pirsa: 08060185 Page 5/61

- Cosmologists are asking questions like:

What is the inflaton potential?

What is the inflaton/s?

- String theorists are asking questions like:

Where in the space of string models is our universe?

What can be predicted?

Cosmologists and String theorists can learn from eachother!

Pirsa: 08060185 Page 6/61

Stringlish to English

Symbol	Name	Approximate meaning
α'	Regge parameter	Inverse string tension
l_s	String length	$=2\pi\sqrt{\alpha'}$ (in our convention)
κ ₁₀	10-d gravitational strength	$=\sqrt{8\pi G_{10}}=l_s^4/\sqrt{4\pi}$, gravitational strength in 10 dims
\bar{m}_{P}	(Reduced) Planck mass	$=1/\sqrt{8\pi G}$, mass scale of quantum gravity in 4 dims
φ	Dilaton	Scalar field that rescales the strength of gravity
a_i	Axions	Pseudo-scalars that appear in the 4-d theory
b_i	Geometric moduli	Scalar fields describing ϕ and the size & shape of the compact space
	- Dilaton modulus ^a	$\sim e^{-\phi}$ (explicit form is model dependent)
	- Kähler moduli	Scalar fields that specify the size of the compact space
	- Complex structure moduli	Scalar fields that specify the shape of the compact space
ψ_i	Complex moduli	$=a_i+ib_i$
ψ	Complex inflaton vector	$=(\psi_1,,\psi_n)$, the complex moduli-vector that can evolve during inflation
φ	Real inflaton vector	$=(a_1,b_1,,a_n,b_n)$, the real moduli-vector that can evolve during inflation
g_s	String coupling	$=e^{\phi}$, the string loop expansion parameter
F_p	p-form field strength	Generalized electromagnetic field strength carrying p-indices
f_p	Flux	$\propto \int F_p$, (normally integer valued) equivalent to a generalized electric or magnetic charge, but can arise purely due to non-trivial topology
g_{10}/R_{10}	10-d string metric/Ricci scalar	Metric/Ricci scalar in the fundamental 10-d action in string frame
g_4/R_4	4-d string metric/Ricci scalar	Metric/Ricci scalar in the effective 4-d action in string frame
g_E/R_E	4-d Einstein metric/Ricci scalar	Metric/Ricci scalar in the effective 4-d action after a conformal transfor- mation to Einstein frame
Pirsa: 080601	Metric on compact space 6-d volume of compact space	2nd block of $g_{10} = \text{diag}(g_4, g_6)$, describing the geometry of compact space $g_{10} = \int_{c_8} d^6x \sqrt{g_6}$ (cs \equiv compact space)
776	6.1 +	A C d manifold that is Dismann flat defined by pariedic identifications

Pirsa: 08060185 Page 8/61

- Review Conditions for Inflation
- Type IIA String Theory
- Compactification

Pirsa: 08060185 Page 9/61

- Review Conditions for Inflation
- Type IIA String Theory
- Compactification

String Theory provides Ingredients for Inflation!

Pirsa: 08060185 Page 10/61

- Review Conditions for Inflation
- Type IIA String Theory
- Compactification

String Theory provides Ingredients for Inflation!

- Fluxes and Potential
- No-Go Theorem

Pirsa: 08060185 Page 11/61

- Review Conditions for Inflation
- Type IIA String Theory
- Compactification

String Theory provides Ingredients for Inflation!

- Fluxes and Potential
- No-Go Theorem

Inflation Constrains String Theory!

Pirsa: 08060185 Page 12/61

- Review Conditions for Inflation
- Type IIA String Theory
- Compactification

String Theory provides Ingredients for Inflation!

- Fluxes and Potential
- No-Go Theorem

Inflation Constrains String Theory!

- Outlook: Evading the No-Go Theorem

Pirsa: 08060185 Page 13/61

Pirsa: 08060185 Page 14/61

- To explain horizon problem, flatness, etc

Pirsa: 08060185 Page 15/61

- To explain horizon problem, flatness, etc
- Introduce scalar field ϕ

$$\mathcal{L} = \frac{1}{16\pi G} R - \frac{1}{2} (\partial_{\mu} \phi)^2 - V(\phi)$$

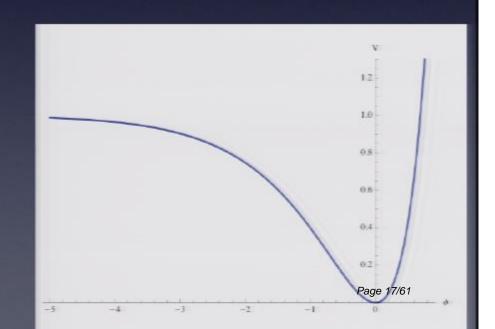
Pirsa: 08060185 Page 16/61

- To explain horizon problem, flatness, etc
- Introduce scalar field ϕ

$$\mathcal{L} = \frac{1}{16\pi G}R - \frac{1}{2}(\partial_{\mu}\phi)^{2} - V(\phi)$$

- Slow-roll:

$$\epsilon = \frac{\bar{m}_p^2}{2} \left(\frac{\partial \ln V}{\partial \phi} \right)^2 \ll 1$$



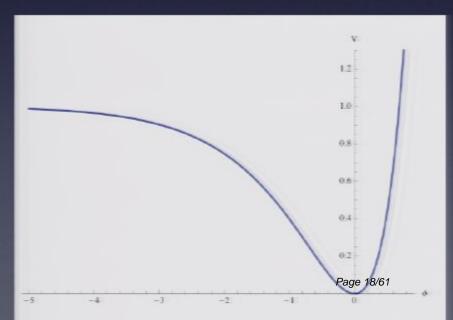
- To explain horizon problem, flatness, etc
- Introduce scalar field ϕ

$$\mathcal{L} = \frac{1}{16\pi G} R - \frac{1}{2} (\partial_{\mu} \phi)^2 - V(\phi)$$

- Slow-roll:

$$\epsilon = \frac{\bar{m}_p^2}{2} \left(\frac{\partial \ln V}{\partial \phi} \right)^2 \ll 1$$

$$\epsilon = \frac{\bar{m}_p^2}{2} \left(\left(\frac{\partial \ln V}{\partial \phi_1} \right)^2 + \left(\frac{\partial \ln V}{\partial \phi_2} \right)^2 + \dots \right)$$



Pirsa: 08060185 Page 19/61

- Fields: $g_{AB}, \phi, H_3, F_0, F_2, F_4, \ldots, D6/O6$

Pirsa: 08060185 Page 20/61

- Fields: $g_{AB}, \phi, H_3, F_0, F_2, F_4, \ldots, D6/O6$

$$S = \frac{1}{2\kappa_{10}^2} \int d^{10}x \sqrt{-g} e^{-2\phi} \left(R + 4(\partial_{\mu}\phi)^2 - \frac{1}{2} |H_3|^2 - e^{2\phi} \sum_p |F_p|^2 \right)$$
$$-\mu_6 \int_{D6} d^7 \xi \sqrt{-g} e^{-\phi} + 2\mu_6 \int_{O6} d^7 \xi \sqrt{-g} e^{-\phi}$$

- Fields: $g_{AB}, \phi, H_3, F_0, F_2, F_4, \ldots, D6/O6$

$$S = \frac{1}{2\kappa_{10}^2} \int d^{10}x \sqrt{-g} e^{-2\phi} \left(R + 4(\partial_{\mu}\phi)^2 - \frac{1}{2} |H_3|^2 - e^{2\phi} \sum_p |F_p|^2 \right)$$
$$-\mu_6 \int_{D6} d^7 \xi \sqrt{-g} e^{-\phi} + 2\mu_6 \int_{O6} d^7 \xi \sqrt{-g} e^{-\phi}$$

Why study this?

- Fields: $g_{AB}, \phi, H_3, F_0, F_2, F_4, \ldots, D6/O6$

$$S = \frac{1}{2\kappa_{10}^2} \int d^{10}x \sqrt{-g} e^{-2\phi} \left(R + 4(\partial_{\mu}\phi)^2 - \frac{1}{2} |H_3|^2 - e^{2\phi} \sum_p |F_p|^2 \right)$$
$$-\mu_6 \int_{D6} d^7 \xi \sqrt{-g} e^{-\phi} + 2\mu_6 \int_{O6} d^7 \xi \sqrt{-g} e^{-\phi}$$

Why study this?

 DeWolfe, Giryavets, Kachru, Taylor (DGKT) in 2005 stabilized all fields in 4D, under parametric control, on CYs

- Fields: $g_{AB}, \phi, H_3, F_0, F_2, F_4, \ldots, D6/O6$

$$S = \frac{1}{2\kappa_{10}^2} \int d^{10}x \sqrt{-g} e^{-2\phi} \left(R + 4(\partial_{\mu}\phi)^2 - \frac{1}{2} |H_3|^2 - e^{2\phi} \sum_p |F_p|^2 \right)$$
$$-\mu_6 \int_{D6} d^7 \xi \sqrt{-g} e^{-\phi} + 2\mu_6 \int_{O6} d^7 \xi \sqrt{-g} e^{-\phi}$$

Why study this?

- DeWolfe, Giryavets, Kachru, Taylor (DGKT) in 2005 stabilized all fields in 4D, under parametric control, on CYs
- IIA with D6/O6 branes on CYs are "semi-realistic"

 e.g., can build MSSM

Compactification I0D to 4D

Pirsa: 08060185 Page 25/61

Compactification I0D to 4D

- We proved 4D theory has the Lagrangian:

$$\mathcal{L} = \frac{1}{16\pi G} R_E - \left(\frac{1}{2} (\partial_\mu \phi_d)^2 + \frac{1}{2} (\partial_\mu \phi_v)^2 + \dots\right) - V(\phi_i)$$

Pirsa: 08060185 Page 26/61

Compactification I0D to 4D

- We proved 4D theory has the Lagrangian:

$$\mathcal{L} = \frac{1}{16\pi G} R_E - \left(\frac{1}{2} (\partial_\mu \phi_d)^2 + \frac{1}{2} (\partial_\mu \phi_v)^2 + \ldots\right) - V(\phi_i)$$

Einstein gravity

Dilaton 6D Volume

All other "moduli" size & shape

Pirsa: 08060185 Page 28/61

- Potential typically needs some fine tuning

$$V(\phi) = \frac{1}{2}m^2\phi^2 + \phi^4 \sum_{n\geq 1} c_n \left(\frac{\phi}{m_p}\right)^n$$

Pirsa: 08060185 Page 29/61

- Potential typically needs some fine tuning

$$V(\phi) = \frac{1}{2}m^2\phi^2 + \phi^4 \sum_{n\geq 1} c_n \left(\frac{\phi}{m_p}\right)^n$$

 So perhaps we should search for some symmetry e.g., Silverstein & Westphal (monodromy)

Pirsa: 08060185 Page 30/61

- Potential typically needs some fine tuning

$$V(\phi) = \frac{1}{2}m^2\phi^2 + \phi^4 \sum_{n\geq 1} c_n \left(\frac{\phi}{m_p}\right)^n$$

 So perhaps we should search for some symmetry e.g., Silverstein & Westphal (monodromy)

Or perhaps just get lucky...

Pirsa: 08060185 Page 31/61

Pirsa: 08060185 Page 32/61

- Typically 100's of fields ϕ_i (100D field space to explore)

Pirsa: 08060185 Page 33/61

- -Typically 100's of fields ϕ_i (100D field space to explore)
- Typically 100's of fluxes $\int F_p = f_p$ (N^100 different potentials to explore)

Pirsa: 08060185 Page 34/61

- -Typically 100's of fields ϕ_i (100D field space to explore)
- Typically 100's of fluxes $\int F_p = f_p$ (N^100 different potentials to explore)
- At least millions of CYs (Many topologies to explore)

Pirsa: 08060185 Page 35/61

- -Typically 100's of fields ϕ_i (100D field space to explore)
- Typically 100's of fluxes $\int F_p = f_p$ (N^100 different potentials to explore)
- At least millions of CYs (Many topologies to explore)

In Summary; we explore an exponentially large number of infinite (10^500) families of 4D models

The Potential V

Pirsa: 08060185 Page 37/61

The Potential V



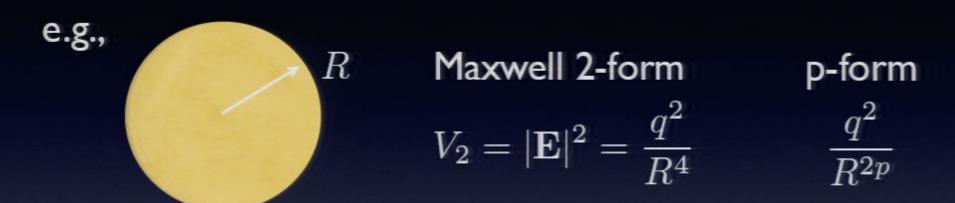
Maxwell 2-form

$$V_2 = |\mathbf{E}|^2 = \frac{q^2}{R^4}$$

p-form

$$\frac{q^2}{R^{2p}}$$

The Potential V



- We proved rigorously (using K and W, or direct dimensional reduction) that in the Einstein frame: $(R \Leftrightarrow \psi_v)$

$$\begin{split} V &= V_3 + \sum_p V_p + V_{D6} + V_{O6} \\ &= \frac{A_3(\phi_j)}{\psi_d^2 \psi_v^3} + \sum_p \frac{A_p(\phi_j)}{\psi_d^4 \psi_v^{p-3}} + \frac{A_{D6}(\phi_j)}{\psi_d^3} - \frac{A_{O6}(\phi_j)}{\psi_{d_{Page} \, 39/6}^3} \end{split}$$

Pirsa: 08060185

Pirsa: 08060185 Page 40/61

$$V = \frac{A_3(\phi_j)}{\psi_d^2 \psi_v^3} + \sum_p \frac{A_p(\phi_j)}{\psi_d^4 \psi_v^{p-3}} + \frac{A_{D6}(\phi_j)}{\psi_d^3} - \frac{A_{O6}(\phi_j)}{\psi_d^3}$$

$$\psi_d = \exp\left(\frac{\phi_d}{\sqrt{2}\,m_p}\right)$$

Where:
$$\psi_d = \exp\left(\frac{\phi_d}{\sqrt{2}\,m_p}\right)$$
 $\psi_v = \exp\left(\frac{\sqrt{2}\,\phi_v}{\sqrt{3}\,m_p}\right)$

$$V = \frac{A_3(\phi_j)}{\psi_d^2 \psi_v^3} + \sum_p \frac{A_p(\phi_j)}{\psi_d^4 \psi_v^{p-3}} + \frac{A_{D6}(\phi_j)}{\psi_d^3} - \frac{A_{O6}(\phi_j)}{\psi_d^3}$$

Where:
$$\psi_d = \exp\left(\frac{\phi_d}{\sqrt{2}\,m_p}\right)$$
 $\psi_v = \exp\left(\frac{\sqrt{2}\,\phi_v}{\sqrt{3}\,m_p}\right)$

Fact:
$$-3\psi_d\frac{\partial V}{\partial\psi_d}-\psi_v\frac{\partial V}{\partial\psi_v}=9V+\sum_p p\,V_p\geq 9V$$

Pirsa: 08060185

$$V = \frac{A_3(\phi_j)}{\psi_d^2 \psi_v^3} + \sum_p \frac{A_p(\phi_j)}{\psi_d^4 \psi_v^{p-3}} + \frac{A_{D6}(\phi_j)}{\psi_d^3} - \frac{A_{O6}(\phi_j)}{\psi_d^3}$$

Where:
$$\psi_d = \exp\left(\frac{\phi_d}{\sqrt{2}\,m_n}\right)$$
 $\psi_v = \exp\left(\frac{\sqrt{2}\,\phi_v}{\sqrt{3}\,m_n}\right)$

$$\psi_v = \exp\left(\frac{\sqrt{2}\,\phi_v}{\sqrt{3}\,m_p}\right)$$

Fact:
$$-3\psi_d \frac{\partial V}{\partial \psi_d} - \psi_v \frac{\partial V}{\partial \psi_v} = 9V + \sum_p p \, V_p \ge 9V$$

$$m_p \left| 3\sqrt{2} \left(\frac{\partial \ln V}{\partial \phi_d} \right) + \sqrt{\frac{3}{2}} \left(\frac{\partial \ln V}{\partial \phi_v} \right) \right| \ge 9$$

$$\epsilon \ge \frac{27}{13}$$
 whenever $V > 0$

- First slow-roll parameter always large so no inflation!

Pirsa: 08060185 Page 44/61

$$V = \frac{A_3(\phi_j)}{\psi_d^2 \psi_v^3} + \sum_p \frac{A_p(\phi_j)}{\psi_d^4 \psi_v^{p-3}} + \frac{A_{D6}(\phi_j)}{\psi_d^3} - \frac{A_{O6}(\phi_j)}{\psi_d^3}$$

$$\psi_d = \exp\left(\frac{\phi_d}{\sqrt{2}\,m_p}\right)$$

Where:
$$\psi_d = \exp\left(\frac{\phi_d}{\sqrt{2}\,m_p}\right)$$
 $\psi_v = \exp\left(\frac{\sqrt{2}\,\phi_v}{\sqrt{3}\,m_p}\right)$

Fact:
$$-3\psi_d \frac{\partial V}{\partial \psi_d} - \psi_v \frac{\partial V}{\partial \psi_v} = 9V + \sum_p p V_p \ge 9V$$

$$m_p \left| 3\sqrt{2} \left(\frac{\partial \ln V}{\partial \phi_d} \right) + \sqrt{\frac{3}{2}} \left(\frac{\partial \ln V}{\partial \phi_v} \right) \right| \ge 9$$

$$\epsilon \ge \frac{27}{13}$$
 whenever $V > 0$

- First slow-roll parameter always large so no inflation!

Pirsa: 08060185 Page 46/61

- First slow-roll parameter always large so no inflation!
- Many many vacua \neq inflation (structure)

Pirsa: 08060185 Page 47/61

- First slow-roll parameter always large so no inflation!
- Many many vacua \neq inflation (structure)
- Corollary: no de Sitter solutions (extreme slow-roll)

Pirsa: 08060185 Page 48/61

- First slow-roll parameter always large so no inflation!
- Many many vacua \neq inflation (structure)
- Corollary: no de Sitter solutions (extreme slow-roll)
- Cosmology: Field vector rolls quickly to AdS vacuum

Pirsa: 08060185 Page 49/61

- First slow-roll parameter always large so no inflation!
- Many many vacua \neq inflation (structure)
- Corollary: no de Sitter solutions (extreme slow-roll)
- Cosmology: Field vector rolls quickly to AdS vacuum

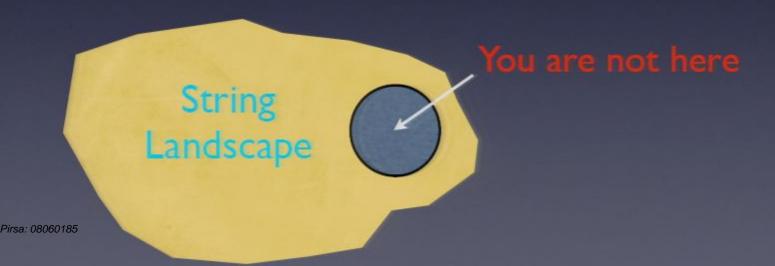


- First slow-roll parameter always large so no inflation!
- Many many vacua \neq inflation (structure)
- Corollary: no de Sitter solutions (extreme slow-roll)
- Cosmology: Field vector rolls quickly to AdS vacuum



- First slow-roll parameter always large so no inflation!
- Many many vacua \neq inflation (structure)
- Corollary: no de Sitter solutions (extreme slow-roll)
- Cosmology: Field vector rolls quickly to AdS vacuum

Page 52/61



- First slow-roll parameter always large so no inflation!
- Many many vacua \neq inflation (structure)
- Corollary: no de Sitter solutions (extreme slow-roll)
- Cosmology: Field vector rolls quickly to AdS vacuum



- First slow-roll parameter always large so no inflation!
- Many many vacua \neq inflation (structure)
- Corollary: no de Sitter solutions (extreme slow-roll)
- Cosmology: Field vector rolls quickly to AdS vacuum



You are not here

Inflation Constrains
String Theory!

Page 54/61

Pirsa: 08060185 Page 55/61

$$-3\psi_d \frac{\partial V}{\partial \psi_d} - \psi_v \frac{\partial V}{\partial \psi_v} = 9V + \sum_p p \, V_p \ge 9V$$

In order to inflate, a IIA compactification must contain some additional structure which gives a term whose scaling leads to a RHS with a coefficient < 9 if positive or > 9 if negative

Pirsa: 08060185 Page 56/61

$$-3\psi_d \frac{\partial V}{\partial \psi_d} - \psi_v \frac{\partial V}{\partial \psi_v} = 9V + \sum_p p \, V_p \ge 9V$$

In order to inflate, a IIA compactification must contain some additional structure which gives a term whose scaling leads to a RHS with a coefficient < 9 if positive or > 9 if negative

e.g., NS5-branes, (non)geometric flux, D8, O4 etc

- Guiding us to tackle less well understood ingredients

Pirsa: 08060185 Page 57/61

$$-3\psi_d \frac{\partial V}{\partial \psi_d} - \psi_v \frac{\partial V}{\partial \psi_v} = 9V + \sum_p p \, V_p \ge 9V$$

In order to inflate, a IIA compactification must contain some additional structure which gives a term whose scaling leads to a RHS with a coefficient < 9 if positive or > 9 if negative

e.g., NS5-branes, (non)geometric flux, D8, O4 etc

- Guiding us to tackle less well understood ingredients



You are not here

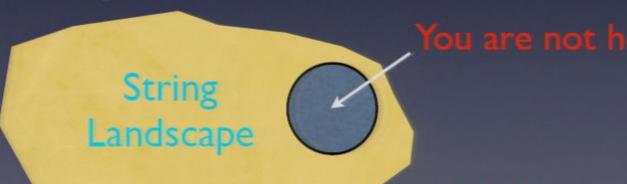
Pirsa: 08060185 Page 58/61

$$-3\psi_d \frac{\partial V}{\partial \psi_d} - \psi_v \frac{\partial V}{\partial \psi_v} = 9V + \sum_p p \, V_p \ge 9V$$

In order to inflate, a IIA compactification must contain some additional structure which gives a term whose scaling leads to a RHS with a coefficient < 9 if positive or > 9 if negative

e.g., NS5-branes, (non)geometric flux, D8, O4 etc

- Guiding us to tackle less well understood ingredients



Maybe you are here

- First slow-roll parameter always large so no inflation!
- Many many vacua \neq inflation (structure)
- Corollary: no de Sitter solutions (extreme slow-roll)
- Cosmology: Field vector rolls quickly to AdS vacuum



$$-3\psi_d \frac{\partial V}{\partial \psi_d} - \psi_v \frac{\partial V}{\partial \psi_v} = 9V + \sum_p p \, V_p \ge 9V$$

In order to inflate, a IIA compactification must contain some additional structure which gives a term whose scaling leads to a RHS with a coefficient < 9 if positive or > 9 if negative

e.g., NS5-branes, (non)geometric flux, D8, O4 etc

- Guiding us to tackle less well understood ingredients



You are not here

Pirsa: 08060185 Page 61/61