Title: Gravitational Radiation from Preheating

Date: Jun 03, 2008 03:15 PM

URL: http://pirsa.org/08060134

Abstract: Parametric resonance, also known as preheating, is a plausible mechanism for bringing about the transition between the inflationary phase and a hot, radiation dominated universe. This epoch results in the rapid production of heavy particles far from thermal equilibrium and has the potential to source a significant stochastic background of gravitational radiation. Here, I present a numerical algorithm for computing the contemporary power spectrum of gravity waves generated in this post-inflationary phase transition for a large class of scalar-field driven inflationary models. I will present the results of this calculation for a number of inflationary models and discuss the (potential) observability of these models

Gravitational Radiation from Preheating

John T. Giblin, Jr

June 3rd, 2008

PASCOS

Perimeter Institute

Easther, Giblin, Lim: astro-ph/0612294
Easther, Giblin, Lim: arXiv/0712.2991

Outline

- Inflation and its end
- Parametric resonance
- Gravitational waves: How and Why
- Gravitational radiation from preheating
  - Methods and numerics
  - Results and parameter dependence

History / Prior Art

- Parametric resonance: Brandenberger & Traschen (1990) / Linde & co
- Khlebnikov and Tkachev: GUT scale inflation and Gravitational waves (1997)
- Yale (2005-) -- Energy Scale/Peak location correlation and lattice simulations
- Now: Garcia-Bellido et al., CITA group, UMW group

Inflation

- (Near) Exponential expansion of the universe
  - Realizable by "particle physics"
  - Energy density of the universe dominated by (scalar-field) potential
- Solves observational "issues" of Big Bang Cosmology
  - Isotropy, Homogeneity
- Quantum Fluctuations provide insight into structure formation

Parametric Resonance

- Consider the toy model

$$\mathcal{L} = \tfrac{1}{2}\partial_\mu\phi\partial^\mu\phi + \tfrac{1}{2}\partial_\mu\chi\partial^\mu\chi - g^2\phi^2\chi^2 - V(\phi)$$

- ϕ is the inflation and χ is a bosonic field. The mode equations for χ are

$$\ddot{\chi}_k + 3H\dot{\chi}_k + \left(\frac{k^2}{a^2} + g^2\phi^2\right)\chi_k = 0$$

- where

$$\tilde{\chi}(k,t) = \int d^3x\; \chi(x,t)e^{2\pi i k\cdot x}$$

Inflation

- (Near) Exponential expansion of the universe
  - Realizable by "particle physics"
  - Energy density of the universe dominated by (scalar-field) potential
- Solves observational "issues" of Big Bang Cosmology
  - Isotropy, Homogeneity
- Quantum Fluctuations provide insight into structure formation

Parametric Resonance

- Consider the toy model

$$\mathcal{L} = \frac{1}{2}\partial_\mu\phi\partial^\mu\phi + \frac{1}{2}\partial_\mu\chi\partial^\mu\chi - g^2\phi^2\chi^2 - V(\phi)$$

- ϕ is the inflation and χ is a bosonic field. The mode equations for χ are

$$\ddot{\chi}_k + 3H\dot{\chi}_k + \left(\frac{k^2}{a^2} + g^2\phi^2\right)\chi_k = 0$$

- where

$$\tilde{\chi}(k,t) = \int d^3x\, \chi(x,t)e^{2\pi i k\cdot x}$$

Parametric Resonance (II)

- If we assume that

$$V(\phi) = \frac{1}{2}m^2\phi^2$$

- *and* H=0 (non-expansion), then

$$\phi(t) = \Phi\sin(mt)$$

- and

$$\ddot{\chi}_k + \left(k^2 + g^2\Phi^2\sin^2(m_\phi t)\right)\chi_k = 0$$

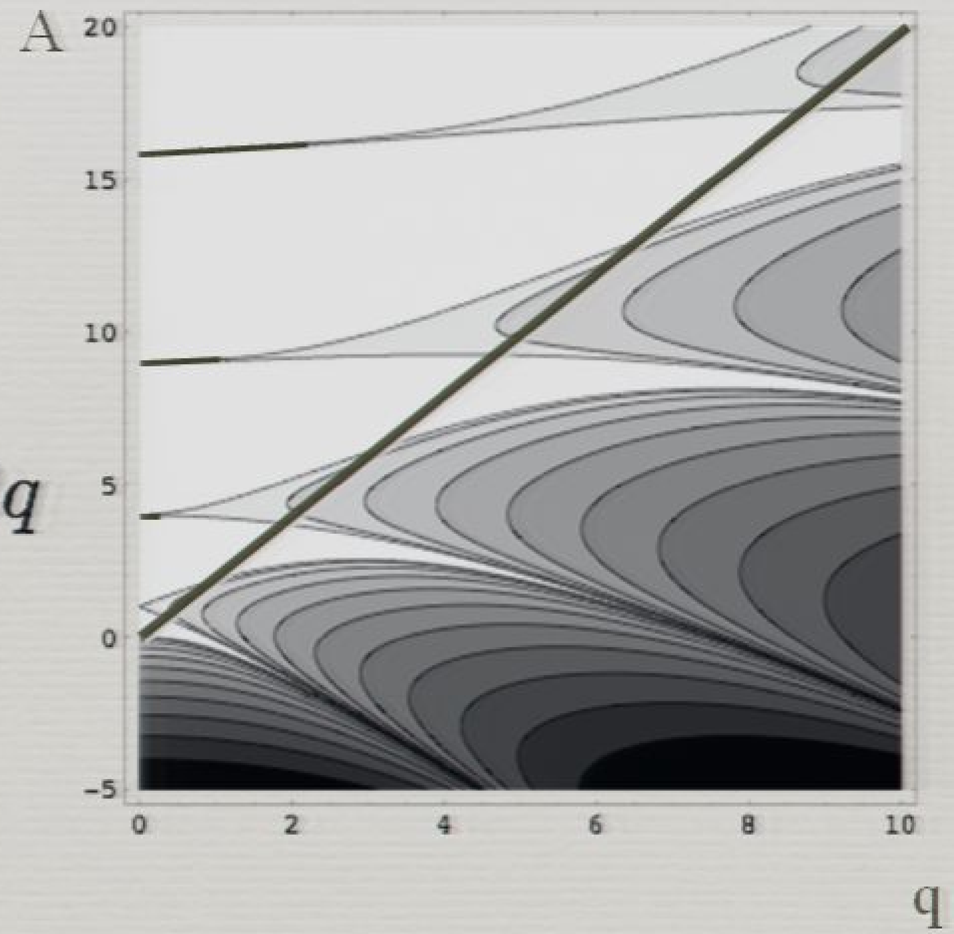
Mathieu Equation

$$z = mt$$

$$q = \frac{g^2\Phi^2}{4m^2}$$

$$A_k = \frac{k^2}{m^2} + \frac{g^2\Phi^2}{2m^2} = \frac{k^2}{m^2} + 2q$$

Solutions are either exponential or oscillatory



$$\chi_k'' + (A_k - 2q\cos(2z))\chi_k = 0$$

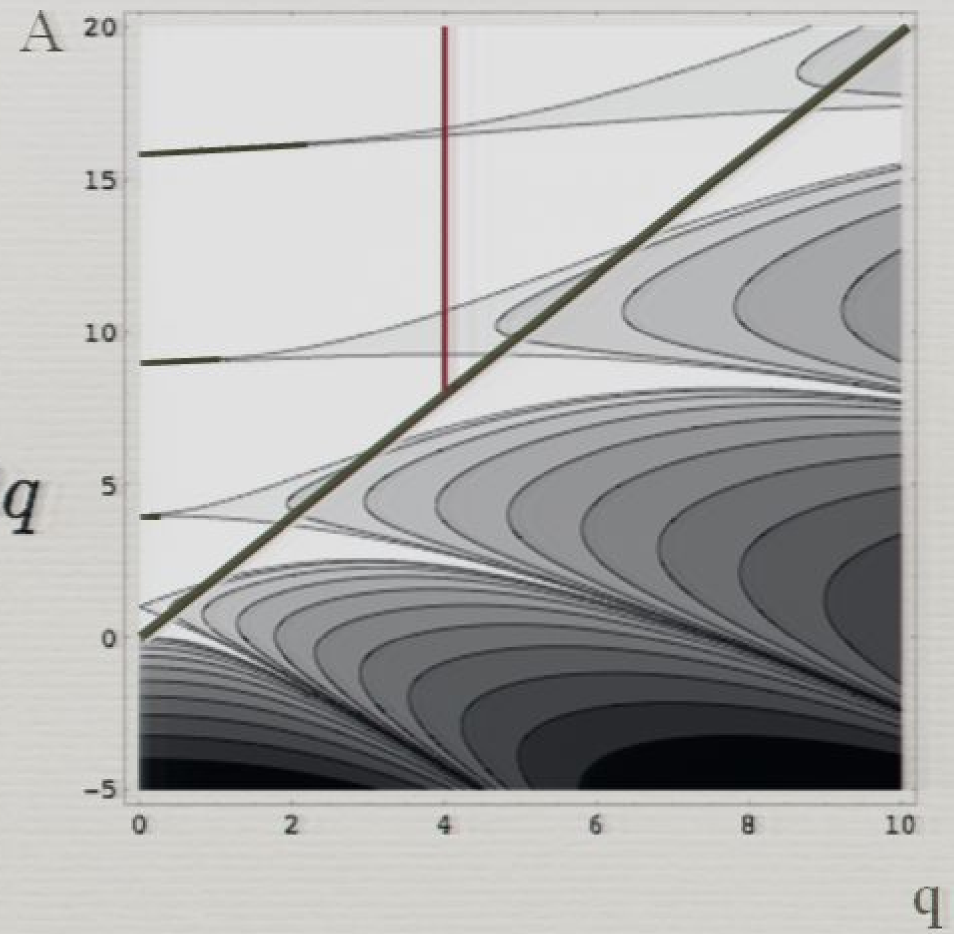
Mathieu Equation

$$z = mt$$

$$q = \frac{g^2\Phi^2}{4m^2}$$

$$A_k = \frac{k^2}{m^2} + \frac{g^2\Phi^2}{2m^2} = \frac{k^2}{m^2} + 2q$$

Solutions are either exponential or oscillatory



$$\chi_k'' + (A_k - 2q\cos(2z))\chi_k = 0$$

Implications

- Certain modes are exponentially excited
  - (copious) particle production of certain momentum states
  - out-of-equilibrium universe
- Large gradient energy
- Any model whose termination oscillates will (approximately) resonate
- Inhomogeneities lead to...

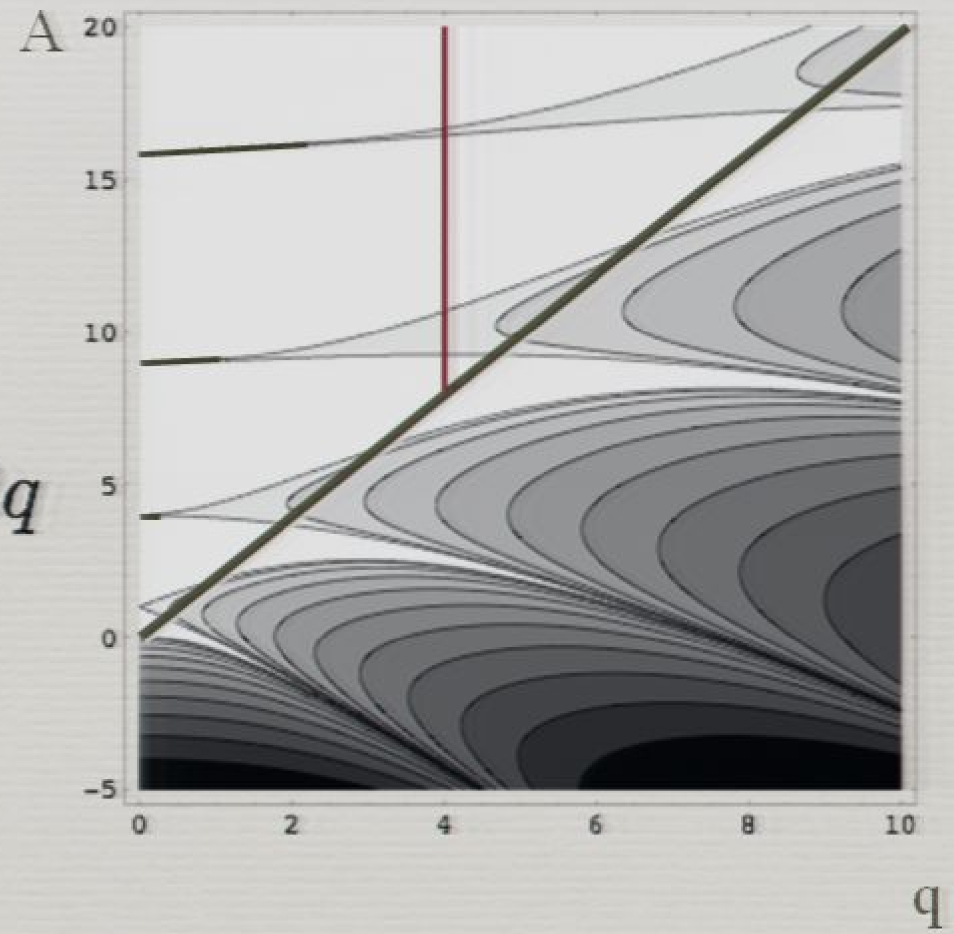
Mathieu Equation

$$z = mt$$

$$q = \frac{g^2\Phi^2}{4m^2}$$

$$A_k = \frac{k^2}{m^2} + \frac{g^2\Phi^2}{2m^2} = \frac{k^2}{m^2} + 2q$$

Solutions are either exponential or oscillatory



$$\chi_k'' + (A_k - 2q\cos(2z))\chi_k = 0$$

Implications

- Certain modes are exponentially excited
  - (copious) particle production of certain momentum states
  - out-of-equilibrium universe
- Large gradient energy
- Any model whose termination oscillates will (approximately) resonate
- Inhomogeneities lead to...

Gravity Waves

- Einstein's Equations

$$G_{\mu\nu} = 8\pi G T_{\mu\nu}$$

- With a metric of the form (in synchronous gauge)

$$ds^2 = dt^2 - a^2(t)\left[\delta_{ij} + h_{ij}\right] dx^i dx^j$$

- Where the perturbation is *transverse-traceless*

$$h^i_i = 0 \quad h^i_{j,i} = 0$$

Gravity Waves (II)

- We can use perturbation theory

$$\delta G_{\mu\nu}(x,t) = 8\pi G \delta T_{\mu\nu}(x,t)$$

- to write equations of motion for the metric perturbations

$$\ddot{h}_{ij} + 3\frac{\dot{a}}{a}\dot{h}_{ij} - \frac{1}{a^2}\nabla^2 h_{ij} = \frac{16\pi G}{a^2} S_{ij}$$

The Source

- The source

$$S_{ij} = T_{ij} - \frac{1}{3} T^k_k \delta_{ij}$$

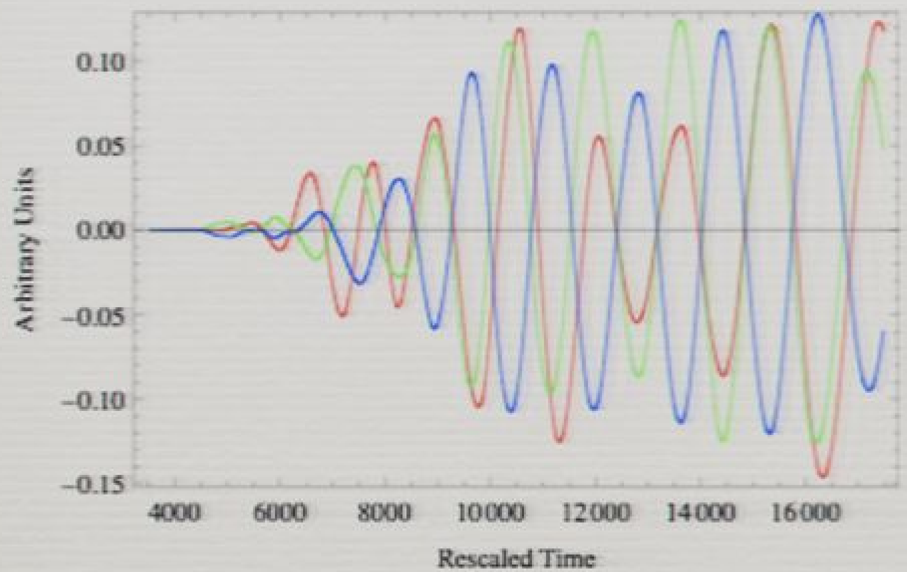
- must be made *transverse-traceless* by

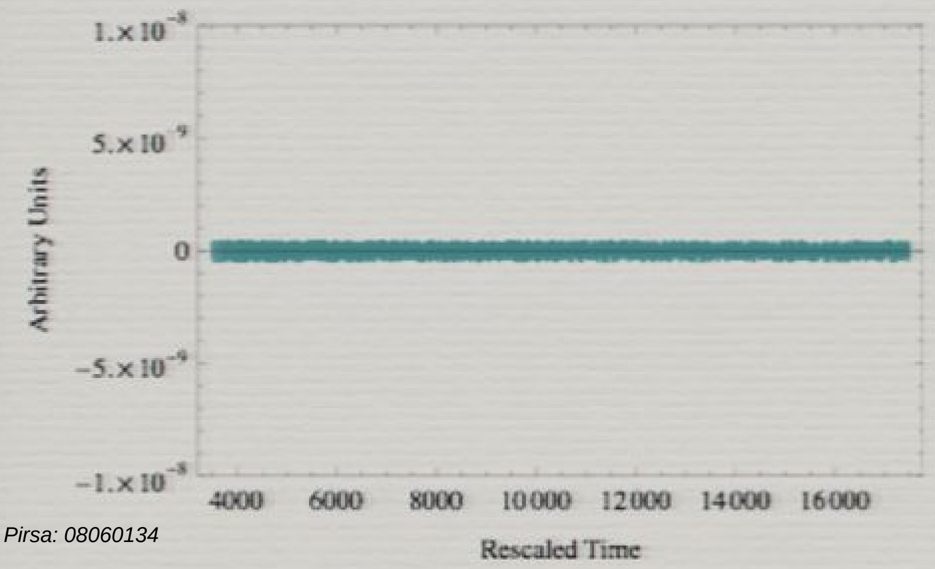
$$S^{TT}_{ij} = P_{ik} S_{kl} P_{lj} - \frac{1}{2} P_{ij} \left(P_{lm} S_{lm} \right)$$

- where

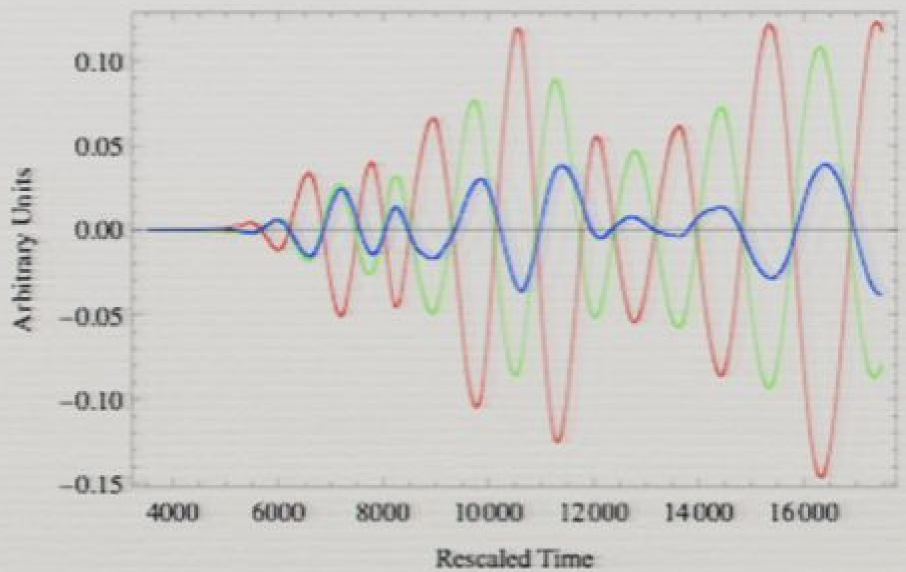
$$P_{ij} = \delta_{ij} - \frac{k_i k_j}{k^2}$$

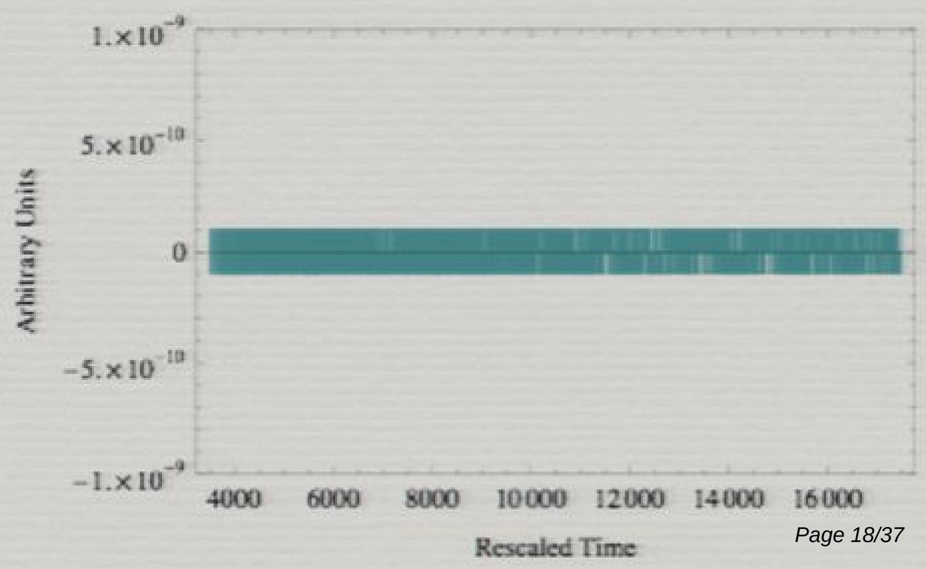
Transverse





Traceless

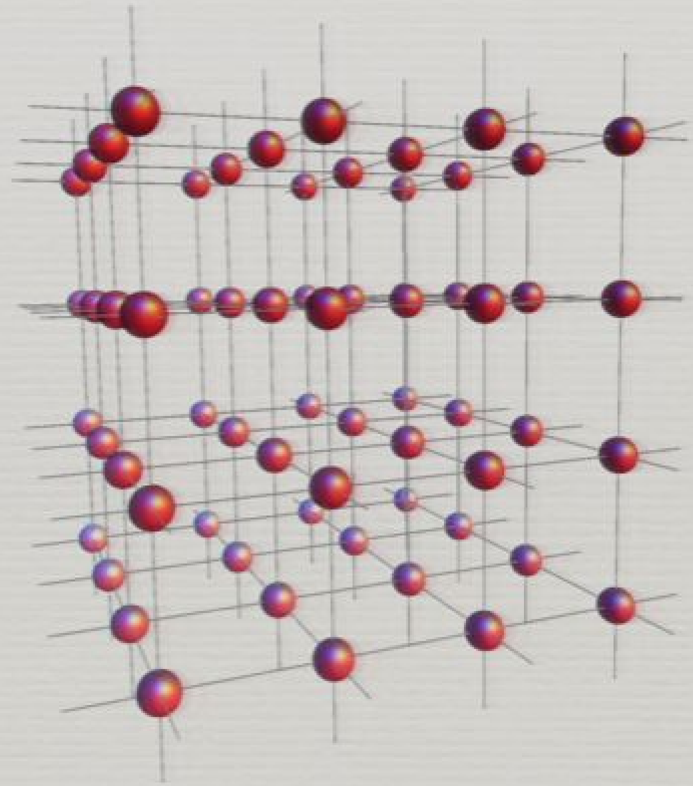




Computational Strategy

- We start by defining a 3-dimensional lattice with periodic boundary conditions and fill it with scalar fields,

$$\phi_i(\vec{x}, t)$$



- We fill the lattice with Scalar fields which obey

$$\ddot{\phi}_i + 3\frac{\dot{a}}{a}\dot{\phi}_i - \frac{1}{a^2}\nabla^2\phi_i + \frac{\partial V(\phi)}{\partial \phi_i} = 0$$

- We use LATTICEEASY to evolve fields *and* the Friedmann equations

$$\left(\frac{\dot{a}}{a}\right)^2 = \frac{8\pi G}{3}\rho$$

- We calculate the source term of

$$\ddot{h}_{ij} + 3\frac{\dot{a}}{a}\dot{h}_{ij} - \frac{1}{a^2}\nabla^2 h_{ij} = \frac{16\pi G}{a^2} S_{ij}$$

- in momentum space.
- Evolve the metric perturbations (in an expanding background) calculate

$$\rho_{gw} = \frac{1}{32\pi G}\left\langle h_{ij,0} h^{ij}_{\ \ ,0}\right\rangle = \sum_{i,j} \frac{1}{32\pi G}\left\langle h_{ij,0}^2\right\rangle$$

Transfer Function

- specifically, we calculate

$$\frac{d\Omega_{gw}}{d\ln k} = \frac{1}{\rho_{crit}}\frac{d\rho}{d\ln k} = \frac{\pi k^3}{3H^2L^2}\sum_{i,j}|h_{ij,0}(k)|^2$$

- and transfer that to the present day using

$$\Omega_{gw}h^2 = \Omega_r h^2 \frac{d\Omega_{gw}(a_e)}{d\ln k}\left(\frac{g_0}{g_*}\right)^{1/3}$$

$$f = 6\times 10^{10}\frac{k}{\sqrt{M_p H_e}}\ \mathrm{Hz}$$

The First Model

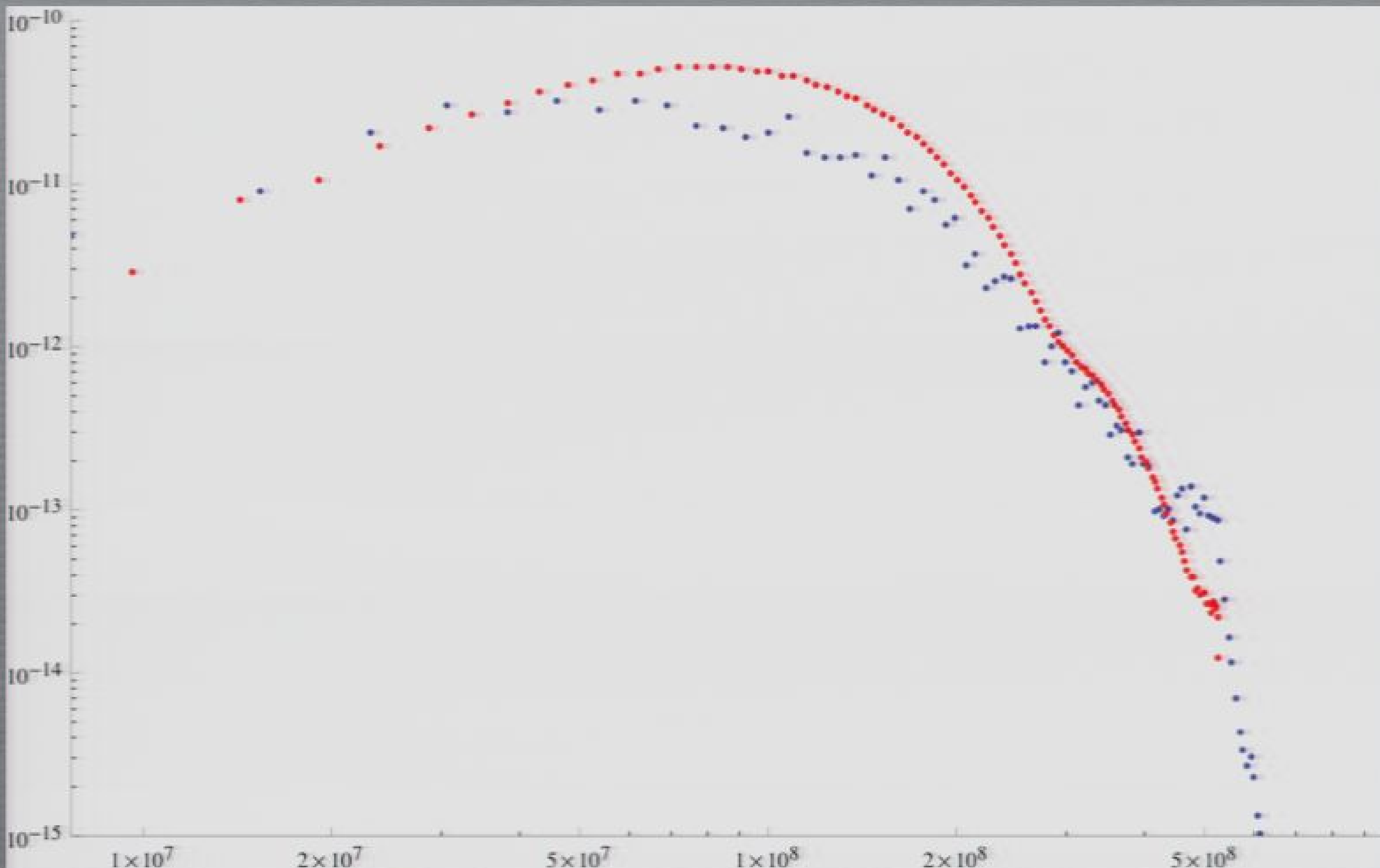
- We start by looking at

$$V(\phi) = \frac{1}{4}\lambda\phi^4$$

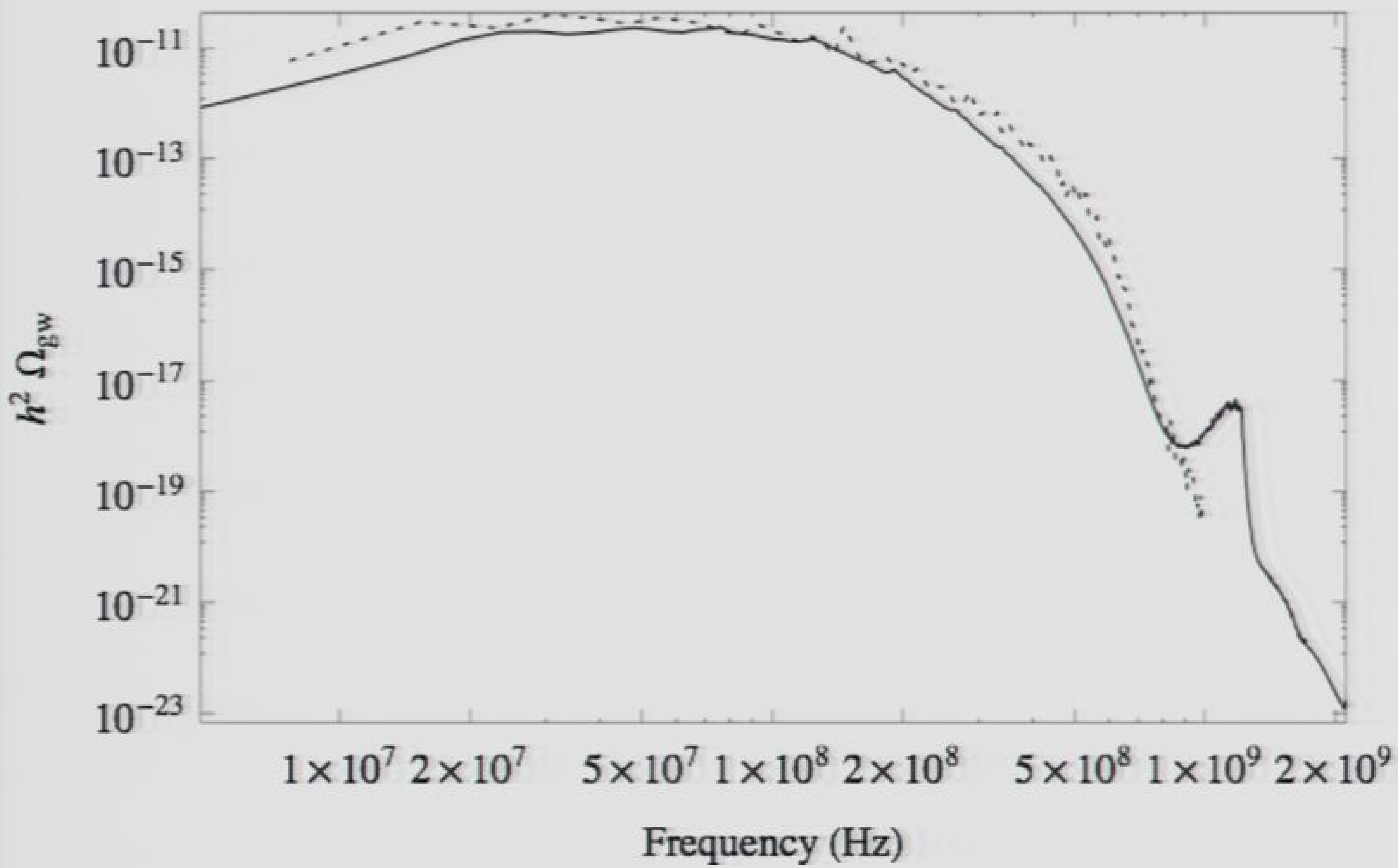
- First used by Khlebnikov and Tkachev
    - hep-ph/9701423
- Useful comparison to other work

State of Affairs

Group	Yale	CITA	Madrid	UWM
Algorithm	Compute tensor source, solve in momentum space	Green's function method	Compute tensor source, solve in position space	Green's function method
Scalar Dynamics	Staggered Leapfrog	Staggered Leapfrog	Staggered Leapfrog	Staggered Leapfrog
Approximations	No backreaction Numerical noise?	<< 1/Hubble "Sample" of k-modes.	Low Scale models have non-expanding background	Expansion is radiation or matter dominated
Comments	Few approximations; noisy sources.	Orthogonal to our approach	Same method as "Yale group" but (sometimes) no expansion	Exact method in radiation era



- Our Current Simulations
- CITA



Yale (solid) and UWM (dotted)

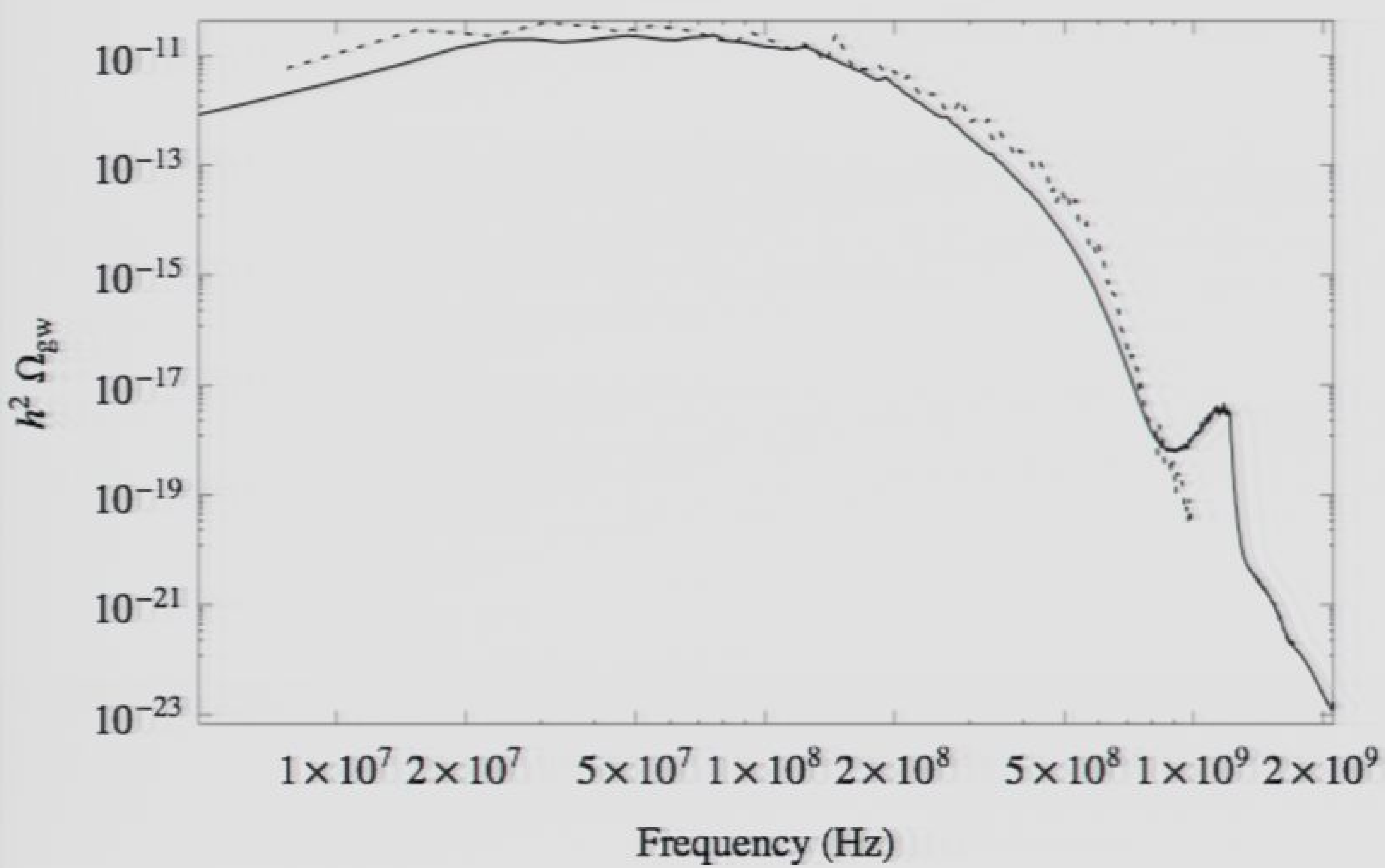
Scaling Argument

- Assuming that the largest possible wavelength to resonate corresponds to (approximately) the Hubble horizon, 1/H, where

$$H \sim \frac{\sqrt{V_e}}{m_{pl}}$$

- we can use our transfer function to calculate

$$f = 6 \times 10^{10} \frac{H_e}{\sqrt{m_{pl} H_e}} \propto V^{1/4}$$



Yale (solid) and UWM (dotted)

Scaling Argument

- Assuming that the largest possible wavelength to resonate corresponds to (approximately) the Hubble horizon, 1/H, where

$$H \sim \frac{\sqrt{V_e}}{m_{pl}}$$

- we can use our transfer function to calculate

$$f = 6 \times 10^{10} \frac{H_e}{\sqrt{m_{pl} H_e}} \propto V^{1/4}$$

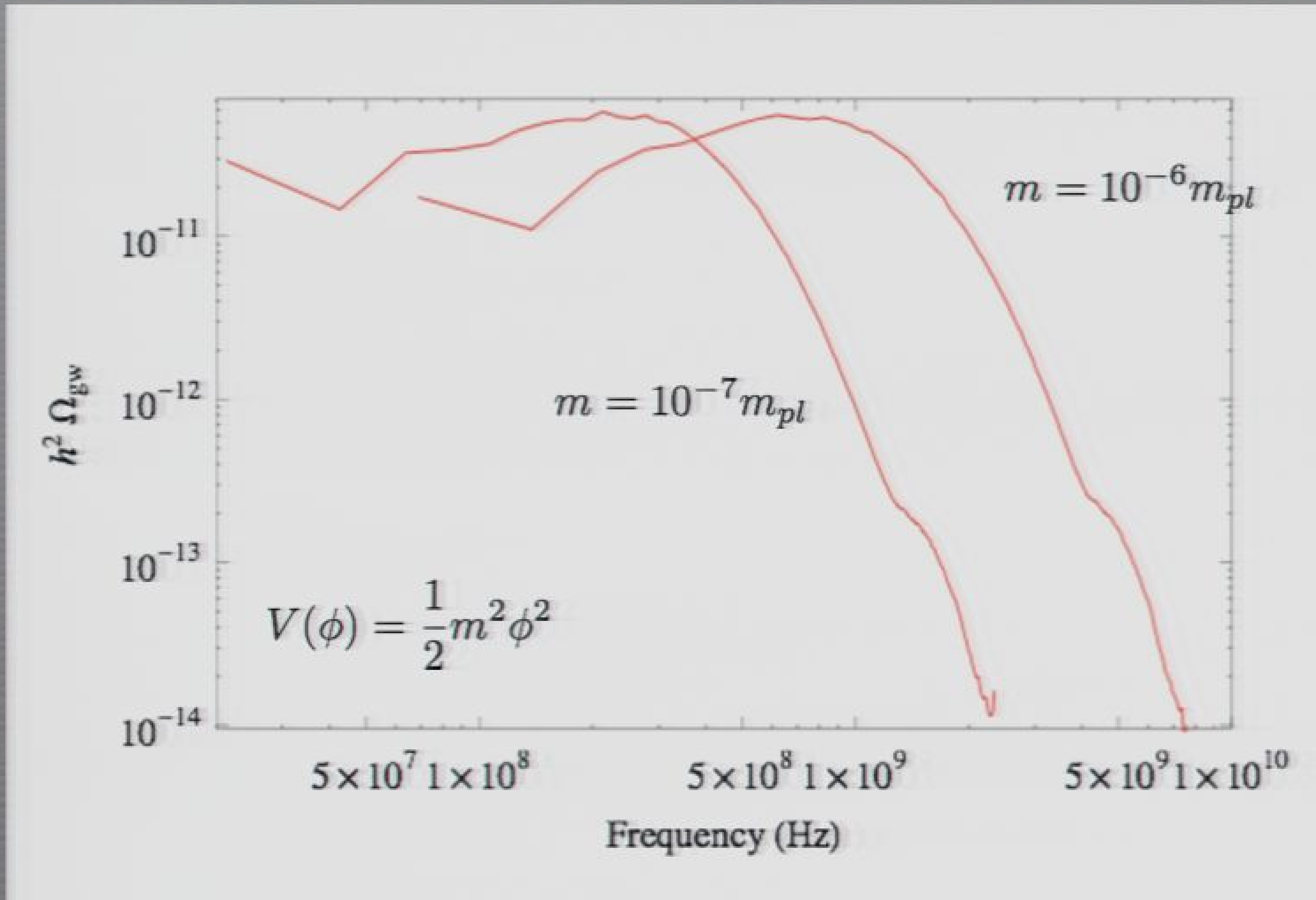
Other Models

- Consider

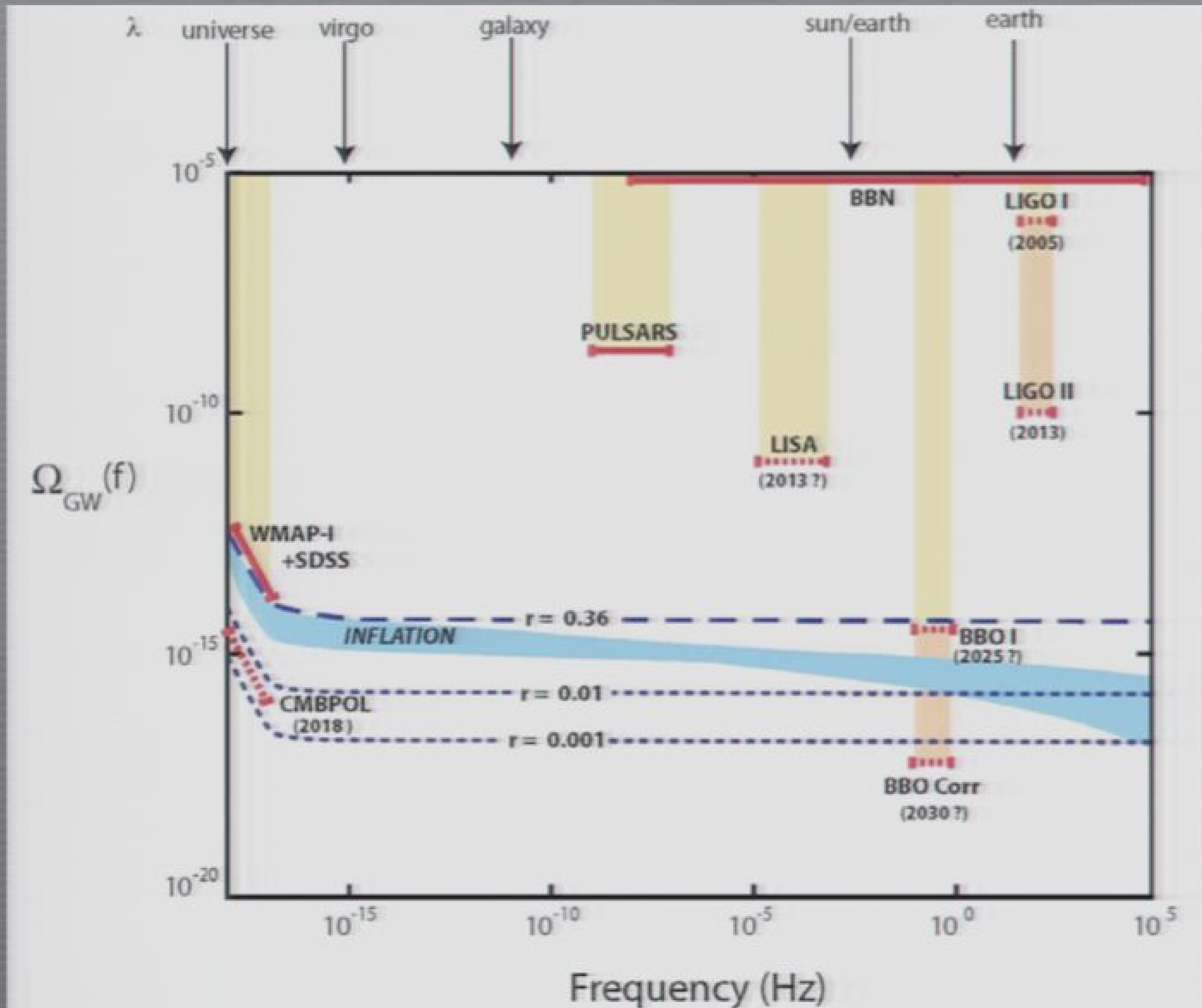
$$V(\phi) = \frac{1}{2}m^2\phi^2$$

- For single field inflation, the CMB places a bound that

$$m \approx 10^{-6} m_{pl}$$


$m = 10^{-6} m_{pl}$
$m = 10^{-7} m_{pl}$
$V(\phi) = \frac{1}{2} m^2 \phi^2$
$h^2\,\Omega_{gw}$
10^{-11}
10^{-12}
10^{-13}
10^{-14}
5×10^7 1×10^8
5×10^8 1×10^9
5×10^9 1×10^{10}
Frequency (Hz)

GUT Scale



The gravitational wave search

Hybrid Inflation

- Let's add a field:

$$V = \frac{(M^2 - \lambda\sigma^2)^2}{4\lambda} + \frac{m^2}{2}\phi^2 + \frac{h^2}{2}\phi^2\sigma^2$$

- For large values of ϕ, σ=0 is a stable point, however as ϕ decreases, σ is drawn to a minimum at

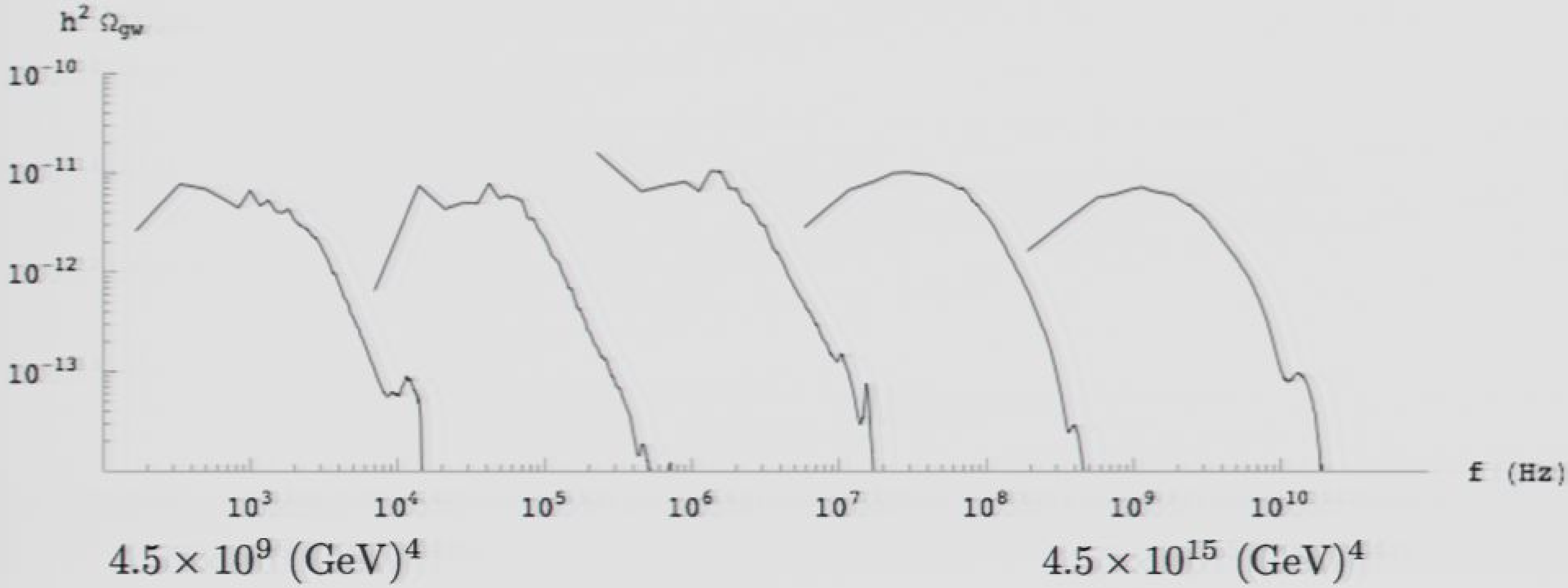
$$\sigma = \langle\sigma\rangle = \frac{M}{\sqrt{\lambda}}$$

Hybrid Inflation (II)

- Assume no fluctuations in σ
- We get an effective potential

$$V(\phi) = \frac{1}{2}\left(m^2 + \frac{h^2 M^2}{\lambda}\right)\phi^2$$

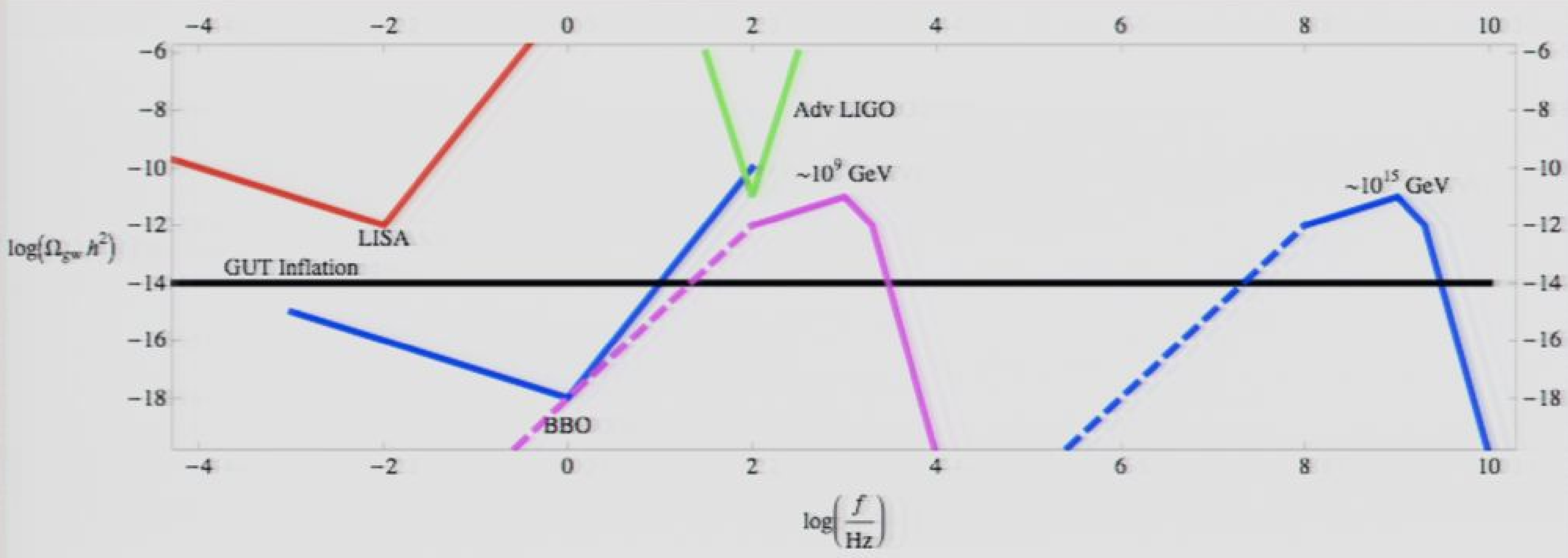
- which is (dynamically) identical to chaotic inflation.



$h^2\ \Omega_{gw}$
10^{-10}
10^{-11}
10^{-12}
10^{-13}
f (Hz)
10^{3}
10^{4}
10^{5}
10^{6}
10^{7}
10^{8}
10^{9}
10^{10}
$4.5 \times 10^{9}\ (\mathrm{GeV})^4$
$4.5 \times 10^{15}\ (\mathrm{GeV})^4$

Low Scale

Where this fits in...



Different from Primordial Spectrum

Primordial	**Preheating**
Quantum source	Classical source
Scale Invariant	Peaked near observable range
Low H : Low amplitude	Low H : redder peak
Always generated	Strongly model dependent
Amplitude bounded by CMB	Amplitude possibly large
$\Omega_{gw,inf}\, h^2 < 10^{-14}$	$\Omega_{gw}\, h^2 \lesssim 10^{-10}$

Conclusions

- Direct detection of gravitational waves "inevitable"?
- Preheating provides a frequency-dependent, constant amplitude probe
- Detection could serve as a "model" selector
- Preheating a new window on inflation
  - Particularly for low scales ($<10^9$GeV)